

# The Compton Spectrometer and Imager

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**Abstract:** The Compton Spectrometer and Imager (COSI) is a NASA Small Explorer (SMEX) satellite mission in development with a planned launch in 2027. COSI is a wide-field gamma-ray telescope designed to survey the entire sky at 0.2–5 MeV. It provides imaging, spectroscopy, and polarimetry of astrophysical sources, and its germanium detectors provide excellent energy resolution for emission line measurements. Science goals for COSI include studies of 0.511 MeV emission from antimatter annihilation in the Galaxy, mapping radioactive elements from nucleosynthesis, determining emission mechanisms and source geometries with polarization measurements, and detecting and localizing multimessenger sources. The instantaneous field of view for the germanium detectors is >25% of the sky, and they are surrounded on the sides and bottom by active shields, providing background rejection as well as allowing for detection of gamma-ray bursts and other gamma-ray flares over most of the sky. In the following, we include an overview of the COSI mission, including the science, the technical design, and the project status.

## 1. Overview

The low-energy part of the gamma-ray regime – the "MeV bandpass" – has not been well-explored due to observational challenges, such as high background levels. However, this is a part of the electromagnetic spectrum where we know that there are signals of great scientific interest, including nuclear lines from radioactive elements and the electron-positron annihilation line at 0.511 MeV. In addition, accreting black holes produce emission at these energies, and this is also a key regime for multimessenger astrophysics (MMA).

Future observations made by the COSI satellite mission will address all of these topics, and its key scientific goals are related to positrons, nuclear lines, polarimetry of accreting black holes, and MMA [18]. COSI is a Compton telescope, and the capabilities that make it unique are its combination of excellent energy resolution, which it gets from cryogenically-cooled germanium detectors, and its large field of view. Thus, COSI is optimized to make all-Galaxy and all-sky emission line images in the MeV bandpass, which will advance our understanding of creation and destruction of matter in our Galaxy.

Here, we provide an update on the COSI mission, focusing on developments since our previous ICRC presentation in July 2021 [18]. COSI was selected by NASA to move to Phase B in October 2021. This was followed by maturing of the requirements and the payload design, and COSI completed a combined System Requirements Review and Mission Definition Review during the first calendar quarter of 2023.

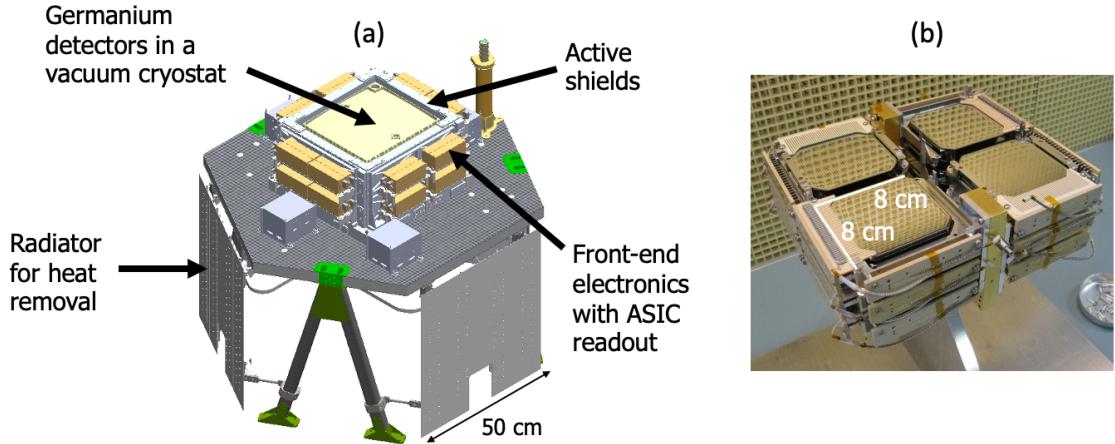
In this proceedings paper, we describe the COSI mission concept and design in Section 2. A summary of how the COSI design enables the key science goals is provided in Section 3. Examples of the extended portfolio of science that will be enabled by COSI are given in Section 4.

## 2. COSI mission concept and design

COSI operates as a Compton telescope when multiple photon interactions are detected in the instrument. This occurs when a gamma-ray enters the instrument and undergoes one or more Compton scattering events before the interactions end in a final photoabsorption event. When the full energy of a gamma-ray is deposited in the germanium detectors, there are two quantities that are obtained: the energy of the gamma-ray is determined by adding up the individual interaction energies; and the Compton angle measured from an axis defined by the first and second interactions defines an event circle on the sky ([21] and references therein). Thus, it is essential that the detectors provide precise energy measurements and 3-dimensional interaction positions.

Figure 1a highlights the main subsystems of the COSI payload: 1. the germanium detectors enclosed in a vacuum cryostat; 2. the front-end electronics; 3. the active shields; and 4. the heat removal subsystem.

Semiconducting germanium detectors are used where gamma-ray interactions liberate electrons and holes which, under high-voltage, drift to opposite sides of the detectors. COSI uses double-sided strip detectors with 64 strips per side and 1.162 mm strip pitch to measure the x and y interaction positions. The depth (z) position is determined by measuring the time difference between the rise of the electron and hole signals. In addition, the strips are surrounded by a guard ring electrode, which acts as a veto signal. The detectors are cooled to between 80 and 90 K using a mechanical cryocooler. Inside the cryostat, there are 16 detectors (4 stacks of 4), and each detector is 8 x 8 x



**Figure 1:** (a) The COSI payload design. (b) The 12 germanium detectors (4 stacks of 3) in the COSI-balloon instrument.

1.5 cm<sup>3</sup>. Figure 1b shows the 12 detectors (4 stacks of 3) that flew on a balloon platform. Special processing of the detectors is required [1].

The heart of the front-end electronics is an Application Specific Integrated Circuit (ASIC). The COSI ASIC has 32 channels with separate circuits for timing, which is required to measure the interaction depth, and spectroscopy [20]. Each detector side is read out using three ASICs: two for the 64 signal strips and one for the guard ring. Thus, each detector requires six ASICs, and COSI's 16 detectors use 96 ASICs.

The detectors are surrounded by active shields on four sides and underneath. The scintillator material used is bismuth germanium oxide (BGO), and they are read out by silicon photomultipliers (SiPMs). The shields have several functions, including passive background attenuation, active anticoincidence for background radiation coming from the side and bottom, anticoincidence of gamma-rays that escape from the germanium detectors, and finally as a gamma-ray transient monitor.

The largest consumer of power in the satellite is for the cryocooler to keep the germanium detectors cold. The cryocooler is attached underneath the instrument. The cryostat heat removal system (CHRS) includes the cryocooler, heat pipes for conducting the heat from the cryocooler, and the radiators.

### 3. Measurement requirements for key science goals

The key science goals for COSI are: A. Uncover the origin of Galactic positrons; B. Reveal Galactic Element Formation; C. Gain insight into extreme environments with polarization; and D. Probe the physics of multimessenger events. Descriptions of how these goals connect to Astro2020 decadal survey white papers can be found in [18] and [19], and a condensed summary follows.

The primary measurements for science goals A and B involve making full Galaxy images at 0.511 MeV and at four nuclear line energies. The specific spectral resolution, angular resolution, and line sensitivity requirements for making these measurements are provided in Table 1. These measurements will allow COSI to determine if the 0.511 MeV emission from the Galactic bulge has substructure and will measure the latitudinal extension of the Galactic disk. They will also provide

information about the electron-positron annihilation conditions across the Galaxy, including the width of the narrow line and the positronium fraction (by comparing the strength of the narrow line to the orthopositronium continuum). The 1.157 MeV measurements provide a search for  $^{44}\text{Ti}$  from young (<1000 years old) supernova remnants. The 1.173 and 1.333 MeV measurements will provide the first Galactic image of  $^{60}\text{Fe}$ , which traces core-collapse supernovae that have occurred over the past millions of years. Finally, the 1.809 MeV measurements of  $^{26}\text{Al}$  from massive stars will provide advances over past COMPTEL and INTEGRAL maps in terms of sensitivity and angular resolution. The COSI required line sensitivities are compared to INTEGRAL/SPI and COMPTEL in Figure 2a.

Science goal C requires measurements of polarization, which COSI is sensitive to because the azimuthal component of the Compton scattering angle depends on the polarization of the incident gamma-ray. The COSI requirement focuses on the measurement of polarization for accreting black holes, and the flux limit given in Table 1 corresponds to being able to measure polarization in the 0.2-0.5 MeV band for the brightest three known Active Galactic Nuclei (AGN): Cen A, 3C 273, and NGC 4151.

**Table 1:** COSI Requirements

Parameter	Requirement
Sky Coverage	25%-sky instantaneous field of view (all-sky every day in survey mode)
Spectral Resolution (FWHM)	6.0 keV at 0.511 MeV and 9.0 keV at 1.157 MeV
Angular Resolution (FWHM)	4.1° at 0.511 MeV and 2.1° at 1.809 MeV
Line Sensitivity <sup>a</sup>	$1.2 \times 10^{-5}$ photons $\text{cm}^{-2} \text{s}^{-1}$ at 0.511 MeV $3.0 \times 10^{-6}$ photons $\text{cm}^{-2} \text{s}^{-1}$ at 1.157 MeV ( $^{44}\text{Ti}$ ) $3.0 \times 10^{-6}$ photons $\text{cm}^{-2} \text{s}^{-1}$ at 1.173 MeV ( $^{60}\text{Fe}$ ) $3.0 \times 10^{-6}$ photons $\text{cm}^{-2} \text{s}^{-1}$ at 1.333 MeV ( $^{60}\text{Fe}$ ) $3.0 \times 10^{-6}$ photons $\text{cm}^{-2} \text{s}^{-1}$ at 1.809 MeV ( $^{26}\text{Al}$ )
Flux limit for polarization <sup>b</sup>	$1.4 \times 10^{-10}$ erg $\text{cm}^{-2} \text{s}^{-1}$ (0.2-0.5 MeV)
Reporting short GRB detections	<1 hour reporting time $\sim 1^\circ$ localizations (accuracy depends on GRB fluence) 100 ms absolute time accuracy

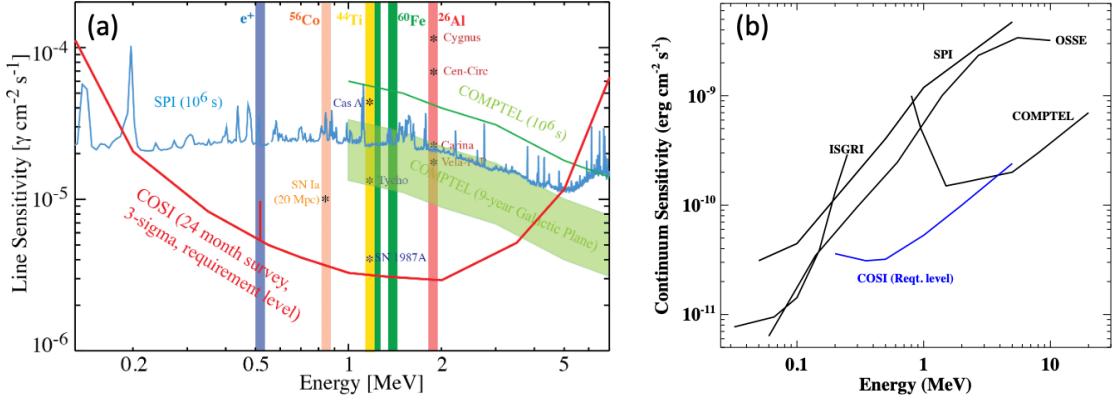
<sup>a</sup>3 $\sigma$  narrow line point source sensitivity in 2-years of survey observations.

<sup>b</sup>For 50% minimum detectable polarization in 2-years of survey observations.

While COSI's gamma-ray measurements have connections to all of the other messengers (gravitational waves, neutrinos, and cosmic rays), it is the detection and reporting of short gamma-ray bursts (GRBs) from merging neutron stars that drives COSI requirements related to Goal D. COSI's large field of view and sensitivity are important for detecting the short GRBs. In addition, the requirement that COSI provides public reports of short GRB localizations within an hour is met by automatically downlinking germanium detector data when gamma-ray transients are detected in the BGO shields.

#### 4. Extended portfolio of science enabled

Beyond the key science goals, COSI's capabilities will enable a large number of additional scientific studies. We describe a few illustrative examples here, which are taken from an array of investigations planned by the COSI science team.



**Figure 2:** Narrow-line (a) and continuum (b) sensitivities based on COSI’s requirements compared to current and previous instruments. The sensitivity curves are for point sources at the  $3\sigma$  level during 2-years of COSI survey time. Due to the all-sky coverage that COSI obtains, these sensitivities will be reached for every source in the sky.

Another multimessenger association is between high-energy neutrinos and gamma-rays. This link was made when the blazar TXS 0506+056 was seen to produce high-energy neutrinos along with gamma-ray flares [10]. Under photo-hadronic processes, the flux of neutrinos detectable by IceCube corresponds to photon emission in the MeV range [13]. The fact that blazar gamma-ray flares typically last for more than a week and that COSI provides daily all-sky coverage in survey mode, COSI will obtain MeV gamma-ray coverage for all high-energy neutrinos seen by IceCube. Thus, we will perform a correlation study between COSI and IceCube to determine which high-energy neutrino events are associated with gamma-rays and to localize the gamma-ray detections to search for blazar associations.

Observing gamma-rays at the MeV bandpass also benefits the study of cosmic dark matter (DM). Attractive DM candidates are predicted in the broad mass range of  $10^{-22}$  eV to  $10^{35}$  g, and COSI has the potential to search for many of them. In the ultralight mass region where DM is much lighter than 1 eV, the axion-like particle (ALP) is an attractive candidate. COSI will be sensitive to the effects of ALPs on the flux and polarization of MeV gamma-rays emitted from, e.g., blazars [6]. In the light mass region where DM is heavier than 1 eV and lighter than the electroweak (EW) scale, many influential DM candidates have been proposed, including the sterile neutrino, dark photon, and light weakly interacting massive particle (WIMP). Those are predicted around the MeV mass scale, and COSI will observe MeV gamma-ray signals emitted by their decay and annihilation [2, 3]. Moreover, the candidates also emit low-energy positrons, which are captured by ambient electrons, form positronium, and contribute to 0.511 MeV emissions. In the heavy mass region where DM is heavier than the EW scale, the heavy WIMP is known to be a well-motivated candidate. Though it primarily emits GeV/TeV gamma-rays by its annihilation, MeV gamma-rays are also produced via inverse Compton scattering caused by energetic electrons from the annihilation. In the ultraheavy mass region where DM is heavier than the Planck scale, primordial black holes (pBH) with mass around  $10^{17}$  g could still compose all of the DM in the present universe. Such pBH would emit both MeV gamma-rays and positrons via Hawking radiation [3].

The prompt emission spectrum of GRBs is often described by the Band function. But the underlying radiation mechanisms are still debated. Synchrotron emission in either ordered or random magnetic fields can contribute to the prompt emission, while photospheric emission may also explain the observed spectrum. COSI can distinguish these mechanisms by measuring the MeV polarization, since they produce distinct polarization signatures [5, 7, 14, 17]: synchrotron polarization degree depends on the orderness of the magnetic field, while the photospheric emission may produce polarization degree from 0 to 100% depending on the geometry of the Compton scattering, which is probably anisotropic in the prompt emission region. By observing the polarization degree of  $>30$  GRBs, COSI will be able to tell synchrotron in an ordered magnetic field apart from synchrotron in a random magnetic field and photospheric models, since the latter two are characterized by a different number distribution of GRB prompt polarization. Additionally, COSI may observe the time-dependent polarization signatures for a few bright GRBs during its mission span, which can further probe the underlying physical conditions in those bright GRBs.

Thermonuclear supernovae (SN Ia) are important tools of modern cosmology and significant contributors to galactic nucleosynthesis and energetics, yet their progenitor systems and explosion mechanisms are poorly understood. As shown in Figure 2a, the expected peak flux of the prominent 0.847 MeV line of  $^{56}\text{Co}$  decay is below the  $10^6$  s COSI sensitivity for a SN Ia at distance 20 Mpc. However, with COSI's complete coverage of the  $\sim 10^7$  s light curve peak and sensitivity to other lines, e.g., 1.238 MeV, COSI should detect  $^{56}\text{Co}$  from SN Ia at this distance, within which there is roughly one per year. For nearer SN Ia, like the previous SN 2011fe and SN 2014J (see, e.g., [4]), COSI could monitor the line light curves and resolve their velocity profiles, yielding independent determinations of the synthesized  $^{56}\text{Ni}$  mass, its distribution within the ejecta, and total ejecta mass and kinetic energy [11]. These would provide important constraints on the thermonuclear explosions and white dwarf progenitor systems.

Classical nova explosions are caused by thermonuclear runaways of material accreted by white dwarfs in a close binary systems. The material ejected at a velocity of several thousands of km/s is enriched with freshly synthesized nuclei. Among these, the unstable  $^{13}\text{N}$  and  $^{18}\text{F}$  decay with a lifetimes of 10 min and 1.8 hr, respectively, by emitting positrons that rapidly annihilate in the dense ejecta in expansion. The resulting annihilation photons and their Comptonization produce a prompt emission ( $\sim 2$  hr) of a 511 keV line and a gamma-ray continuum  $< 511$  keV [8, 9], which has not been detected yet [15, 16]. The long-lived radioactive nuclei synthesized in the explosion are  $^7\text{Be}$  in CO white dwarfs and  $^{22}\text{Na}$  in ONe white dwarfs. Their lifetimes (53 days and 2.6 yr, respectively) are long enough for them to decay in a mostly transparent ejecta by emitting gamma-ray lines at 478 keV and 1275 keV, respectively. The detection of the prompt emission with COSI would provide information on the outburst conditions and on the properties of the ejecta, while the detection of 478 keV and 1275 keV lines would provide the type (CO or ONe) of the underlying white dwarf and the mass of released radioactive isotopes, which give unique constraints to explosive nucleosynthesis models.

COSI's continuum sensitivity will also provide new constraints on MeV-band emission from the Fermi Bubbles [12] and the Galactic Diffuse Continuum Emission (GDCE). The GDCE in the COSI energy band is dominated by inverse Compton radiation, produced by cosmic ray (CR) electrons up-scattering low-energy photons of the interstellar radiation field. There is also expected to be subdominant contributions from Bremsstrahlung radiation, unresolved point sources, and

plausibly other sources as well. COSI's continuum sensitivity and wide field of view will enable measurements of a highly precise spectrum of the GDCE over the entire sky. Such measurements will play a major role in disentangling the different components of the emission, including searches for new physics, and they will likely lead to new insights regarding the sources of CR electrons in the Galaxy. More details regarding COSI's prospects for probing the GDCE can be found in the ICRC2023 proceedings paper by Karwin et al.

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