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ABSTRACT

Accurately localizing fast radio bursts (FRBs) is essential for understanding their birth environments and for their use as cosmological probes. Recent advances in radio interferometry, particularly with MeerKAT, have enabled the localization of individual bursts with arcsecond precision. In this work, we present the localization of 15 apparently non-repeating FRBs detected with MeerKAT. Two of the FRBs, discovered in 2022, were localized in 8 s images from the projects that MeerTRAP was commensal to, while eight were localized using the transient buffer (TB) pipeline, and another one through SeeKAT, all with arcsecond precision. Four additional FRBs lacked TB triggers and sufficient signal, limiting their localization only to arcminute precision. For eight of the FRBs in our sample, we identify host galaxies with greater than 90 per cent confidence, and one with 80 per cent confidence, while two FRBs have ambiguous associations. We measured spectroscopic redshifts for six host galaxies, ranging from 0.33 to 0.85, demonstrating MeerKAT's sensitivity to high-redshift FRBs. We modelled the spectral energy distributions of host galaxies with sufficient photometric coverage to derive their stellar population and star formation properties. This work represents one of the largest uniform samples of well-localized distant FRBs to date, laying the groundwork for using MeerKAT FRBs as cosmological probes and understand how FRB hosts evolve at high redshift.

Key words: methods: data analysis – methods: observational – techniques: interferometric – fast radio bursts.

1 INTRODUCTION

Fast radio bursts (FRBs) are radio transient signals lasting from micro to milliseconds, with luminosities high enough to be seen

from galaxies beyond the Milky Way (MW) (J. M. Cordes & S. Chatterjee 2019; E. Petroff, J. W. T. Hessels & D. R. Lorimer 2019, 2022). The first FRBs were discovered using the 64 m single-dish Parkes radio telescope, and their extragalactic origin was inferred from their unusually large dispersion measures (DMs), which far exceeded the predictions from MW models, as well as from their

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isotropic distribution across the sky (D. R. Lorimer et al. 2007; D. Thornton et al. 2013). Early FRB discoveries were made with single-dish instruments, including both non-repeaters and the first repeater, FRB 20121102A, discovered with the Arecibo radio telescope (L. G. Spitler et al. 2014, 2016). However, interferometric observations with radio telescope arrays became crucial for precisely localizing FRBs to their host galaxies and thereby confirming their extragalactic origin. The repetitions from FRB 20121102A enabled follow-up observations with the Karl G. Jansky Very Large Array (VLA), leading to its localization to a dwarf star-forming galaxy at $z \sim 0.2$ (S. Chatterjee et al. 2017; S. P. Tendulkar et al. 2017), and definitively establishing the cosmological origin of at least some FRBs.

In the past decade, the advent of new radio telescope arrays with dedicated FRB surveys has rapidly pushed the field forward. This results from a combination of increased collecting area, which improves instantaneous sensitivity, multibeam receivers, which increase the field of view (FoV), and longer baselines, which facilitate precise localization. Consequently, thousands of FRBs have been detected to date (CHIME/FRB Collaboration 2021), more than 50 have been found to repeat (CHIME/FRB Collaboration 2023), and nearly a hundred FRBs have been localized to their host galaxies (K. E. Heintz et al. 2020; S. Bhandari et al. 2022; M. Caleb et al. 2023; A. C. Gordon et al. 2023; M. Bhardwaj, J. Lee & K. Ji 2024; L. N. Driessen et al. 2024; K. M. Rajwade et al. 2024; R. M. Shannon et al. 2025; K. Sharma et al. 2024).

Recent discoveries have further revealed crucial information about the FRB nature, which strengthens the link between FRBs and young magnetized neutron stars (e.g. C. D. Bochenek et al. 2020; CHIME/FRB Collaboration 2020). Other properties of both one-off and repeating FRBs, such as polarization (R. Mckinven et al. 2024) and intrinsic structure (I. Pastor-Marazuela et al. 2025), further resemble those observed in the Galactic pulsar and magnetar population.

Studying the host galaxy properties from which FRBs are emitted can further expose the environments where they reside, their age, and the nature of their progenitors. Repeating sources facilitate the follow-up with very long baseline interferometry (VLBI) to obtain milliarcsecond precision localizations. This has shown, for instance, that while some repeaters are located right outside star-forming regions (B. Marcote et al. 2020), others are located in globular clusters (F. Kirsten et al. 2022), or are associated with very old stellar populations (T. Eftekhari et al. 2025; V. Shah et al. 2025), and thus are likely to have very different ages or formation channels. However, only ~ 2.6 per cent of the currently detected FRBs seem to be repeaters (CHIME/FRB Collaboration 2023), thus emphasizing the importance of instantaneous localization of any detected FRBs. For this purpose, many FRB surveys have the capacity to store voltage data when a new FRB is detected, which is then imaged in order to determine the FRB position. With this method, ASKAP (D. R. Scott et al. 2023), DSA-110 (V. Ravi et al. 2023), and MeerTRAP (K. M. Rajwade et al. 2024) can currently achieve sub-arcsecond localization, and while CHIME/FRB can only attain sub-arcminute localization on its own (D. Michilli et al. 2022), the CHIME/FRB Outriggers has recently started to localize FRBs with milliarcsecond precision (A. E. Lanman et al. 2024; CHIME/FRB Collaboration 2025b; FRB Collaboration 2025a).

The identification of FRB host galaxies via optical follow-up observations has enabled the study of host galaxy demographics. These studies have revealed that FRB host galaxies are generally star-forming, with a wide range of stellar masses and star formation rates (SFRs; K. E. Heintz et al. 2020; S. Bhandari et al. 2022; A. C. Gordon et al. 2023). While there is as yet no significant difference

between the hosts of repeating and one-off FRBs, repeaters may occur in galaxies with lower stellar masses (e.g. S. Bhandari et al. 2022; A. C. Gordon et al. 2023). Additionally, some studies have suggested that FRBs primarily trace star formation in galaxies (N. Loudas et al. 2025) with a possible preference towards metal-rich massive star-forming galaxies (K. Sharma et al. 2024), suggesting a bias towards environments favourable to magnetar formation through core-collapse supernovae or stellar mergers. However, other studies have posited that FRBs do not trace star formation alone, and stellar mass must also play a role, indicative of more delayed progenitor formation channels relative to star formation (A. Horowitz & B. Margalit 2025). Current interpretations are limited by small-number statistics, particularly at higher redshifts where only a handful of FRBs have been accurately localized. Increasing the sample of FRB hosts is hence essential for constraining their progenitor channels.

In addition to providing key insights into the nature of FRBs, accurate localization and redshift determination facilitates their use as cosmological probes. One of the first such applications was using FRBs to measure the baryonic content of the intergalactic medium (IGM) through the Macquart relation (J.-P. Macquart et al. 2020). Beyond this, other proposed applications include new measurements of the Hubble constant and other cosmological parameters (S. Hagstotz, R. Reischke & R. Lilow 2022; C. W. James et al. 2022), the tomographic reconstruction of the cosmic web (S. Simha et al. 2020; K.-G. Lee et al. 2022; I. S. Khrykin et al. 2024), and probing the epoch of reionization (see S. Bhandari & C. Flynn 2021, and references therein). Gravitationally lensed FRBs could also provide powerful cosmological probes (I. Pastor-Marazuela 2025). These applications, however, rely on the detection and localization of FRBs at relatively high redshift, a population that remains limited.

In this paper, we present the localization of 15 new FRBs discovered with the MeerKAT radio telescope. Of those, 11 have been localized with arcsecond precision, thus allowing for host galaxy identification. Given their moderate to high DMs, these FRBs hold significant potential for the aforementioned cosmological applications, which will be presented in Caleb, Rajwade et al. (in preparation). The burst properties will be presented in Pastor-Marazuela et al. (in preparation). The paper is divided as follows: in Section 2, we present the MeerKAT observations in which the bursts were detected. In Section 3, we describe our host galaxy data analysis methods and new observations. In Section 4, we detail the localization of our FRB sample and their host galaxy properties. Finally, we conclude in Section 4.

2 OBSERVATIONS

2.1 MeerTRAP real-time pipeline

The MeerKAT (More (meer) Karoo Array Telescope) radio telescope, located in the Northern Cape Province in South Africa, is an array of 64 13.5 m diameter antennas with a maximum baseline of 7.7 km (J. Jonas & the MeerKAT Team 2018). The antennas have three receivers that allow MeerKAT to observe in the UHF (544–1088 MHz), *L* band (856–1712 MHz), or *S* band (five different 875 MHz bandwidths from 1750 to 3500 MHz). The MeerTRAP (More (meer) TRAnsients and Pulsars) project performs a commensal survey for short radio transients, simultaneous to other ongoing MeerKAT observations (S. Sanidas et al. 2017; K. Rajwade et al. 2021). To search for radio transient signals, the Filterbanking Beamformer User Supplied Equipment (FBFUSE; E. D. Barr 2017; W. Chen et al. 2021) generates one incoherent beam (IB) and up to 768 coherent beams (CBs) formed by coherently adding the signals

from up to 40 of the inner MeerKAT antennas (N_{ant}). This CB tiling significantly increases the sensitivity, which scales as N_{ant} , though at the cost of a smaller FoV. In contrast, the IB, which sums the signal from up to 64 antennas without phasing them, offers a larger FoV but with a sensitivity scaling as $\sqrt{N_{\text{ant}}}$. In the L band, the CBs cover $\sim 0.4 \text{ deg}^2$, whereas the IB spans a larger area of $\sim 1 \text{ deg}^2$. This set-up effectively results in two complementary FRB surveys: one using the CBs that is roughly five times more sensitive than the IB but has a smaller FoV, and the other using the IB with a broader FoV but lower sensitivity (F. Jankowski et al. 2023). The number of CBs (≤ 768) is determined by the computational resources of the FBFUSE and the Transient User Supplied Equipment cluster (K. Rajwade et al. 2021; K. M. Rajwade et al. 2022), which searches for dispersed single pulses in the FBFUSE output in near-real time. To mitigate false detections due to radio frequency interference (RFI) occurring during the observations, we applied the outlier detection algorithm named inter-quartile range mitigation (V. Morello, K. M. Rajwade & B. W. Stappers 2021), developed for and tested on data from the Lovell radio telescope (UK) and MeerKAT. After RFI mitigation, dispersed single pulse searches are carried out with *AstroAccelerate* (K. Adámek & W. Armour 2020; W. Armour et al. 2020), a GPU software that incoherently dedisperses the signal up to a DM of $\sim 4000 \text{ pc cm}^{-3}$ in the L band, with a sampling interval of $306 \mu\text{s}$ and searching up to a maximum boxcar width of 0.67 s . The burst search parameters, including DM ranges and boxcar widths, are optimized in a way that allows for the data to be searched in real-time with the available computing resources, and to have enough time to trigger the transient buffer (TB) data while the burst is held in memory. Although some highly dispersed FRBs could be missed due to this DM cut-off, the most dispersed MeerTRAP FRB found so far is $> 1000 \text{ pc cm}^{-3}$ away from this limit, making the bias against such bursts small. We have included Table A1 in Appendix A detailing the limits that were used for each band at different times.

An overview of the MeerTRAP FRB population, including detection rates in the CBs and IB, was presented in F. Jankowski et al. (2023). While we have briefly discussed relevant aspects of the pulse search pipeline and configuration, a full analysis of the selection effects, detection biases (e.g. between CB and IB detections), and the impact of RFI on false positives and trigger recovery rates is beyond the scope of this paper and will be addressed in future work.

2.2 Transient buffer data

If a radio transient candidate is identified by the real-time search pipeline within 45 s of its detection time, it triggers the storage of the TB data via a custom VOEVENT (E. Petroff et al. 2017) emission software (F. Jankowski et al. 2022). The TB data that are stored contain 300 ms of channelized complex data from all MeerKAT antennas that were available at the time of detection (up to 64), centred at the time of the burst candidate. To ensure the full burst is recovered, even in the case of large dispersion delays, the TB data are incoherently dedispersed following the dispersion curve of the burst. This allows us to reconstruct the full burst profile within the available TB duration and prevents temporal truncation (K. M. Rajwade et al. 2024). The full observing bandwidth of the TB data is divided into 64 frequency subbands of equal bandwidth, and each subband is written to a different storage node. Some subbands might be missing due to a malfunction of the storage nodes, but this does not usually exceed 10 per cent of the full bandwidth. Because the TB data contain the signal from all available antennas that we add coherently, the resulting S/N of the burst is significantly higher than

in the original CB or IB detection. A typical loss of 10 per cent of the bandwidth due to missing subbands does not have a substantial impact on the detection of the bursts in the image domain or on the astrometric accuracy. In practice, additional frequency subbands are flagged due to RFI, which commonly exceeds the fraction due to storage node malfunctions. The loss of subbands due to node failures is hence not the dominant limitation on burst signal recovery or localization precision. The TB data are used to localize the FRBs and to obtain high spectro-temporal resolution polarimetric data. While this paper focuses only on the localization achieved with the TB data, the full TB pipeline has been detailed in K. M. Rajwade et al. (2024), and the pulse properties of the FRBs in this work will be presented in Pastor-Marazuela et al. (in preparation).

3 DATA ANALYSIS

3.1 Localization of FRBs with TB data

The FRBs that triggered TB dumps upon detection were localized following the method described in K. M. Rajwade et al. (2024). To summarize, the TB data are incoherently dedispersed to the detection DM and written to disc. Next, they are correlated with *xGPU* (M. A. Clark, P. C. La Plante & L. J. Greenhill 2011), and converted into visibility files with the appropriate metadata in a FITS-Interferometry Data Interchange format, and then we use the COMMON ASTRONOMY SOFTWARE APPLICATIONS package (CASA; CASA Team 2022) to convert the visibility files into measurement sets (MSs). Since the TB data divide the full observing bandwidth into 64 frequency subbands that follow the dispersion delay, we obtain up to 64 MSs.

Once the MSs are obtained, we create images with *WSClean* (A. R. Offringa et al. 2014). Initially, we perform a simple dirty clean step to quickly inspect the images and identify when and where the FRB appears. We create an image with the full 300 ms integration time for each of the 64 frequency subbands, visually inspecting them to flag out those affected by RFI. We then produce a frequency-averaged image from the remaining subbands. Next, we generate frequency-averaged images of 11 time bins of each data set around the expected burst arrival time, with each time bin having an integration time of the order of the burst duration as seen in the detection filterbank file. By subtracting the full integration time images, we generate difference images, in which the FRB should appear as a new source in one or more of the central time bins. After identifying the time bins where the FRB is detected, we produce new images with more advanced cleaning parameters (see K. M. Rajwade et al. 2024, section 2.4.2). These new images include a full integration time image, an ‘on’ image corresponding to the time intervals where the FRB is detected, and an ‘off’ image from intervals where the source is not detected (usually the first time intervals of the MSs) with the same duration as the ‘on’ image. The images are generally produced with 8192 pixels per side, and 1 arcsec^2 pixel size.

The next step involves performing an astrometric correction of the resulting images and source positions, following the procedure outlined in L. N. Driessen et al. (2022) and L. N. Driessen et al. (2024). First, we run the Python Blob Detector and Source Finder (PyBDSF; N. Mohan & D. Rafferty 2015) algorithm on the full integration, ‘on’ and ‘off’ images to identify the image source positions and associated errors. PyBDSF determines the positions and associated errors by fitting 2D Gaussian models to the radio images. The positional errors hence reflect the precision of the fitted centroid, which depends on the S/N of the source and the beam size, rather than the apparent angular extent of the source. As a result,

the positional errors can be significantly smaller than the observed sources.

Next, we query various radio source catalogues to find reference sources for the astrometric correction. Although catalogues with high positional accuracy are preferred, such as the Long Baseline Array (LBA) Calibrator Survey (LCS1; L. Petrov et al. 2011), or the Australian Telescope Compact Array Parkes-MIT-NRAO (ATPMN; D. McConnell et al. 2012), these are often sparsely populated and may lack sufficient sources within the MeerKAT FoV for a reliable correction. The Very Large Array Sky Survey (VLASS; M. Lacy et al. 2020) Epoch 2 source catalogue goes down to declinations of -40° , it usually contains tens of sources within the MeerKAT FoV, and the sources have a typical positional accuracy of 0.2 arcsec. We thus use VLASS unresolved sources as our default reference for FRBs at declinations $> -40^\circ$, and assume a systematic error of 0.2 arcsec.

For declinations $\leq -40^\circ$, the Rapid ASKAP Continuum Survey (RACS; C. L. Hale et al. 2021) typically contains tens to hundreds of sources within the FoV. However, due to insufficient VLBI reference positions in the Southern hemisphere, the RACS-mid (1367.5 MHz; S. W. Duchesne et al. 2023) source positions have systematic astrometric offsets of $\sim 1 - 2$ arcsec. To overcome this, we use the Radio Fundamental Catalog (RFC¹) as a reference to correct the RACS-mid source positions. The RFC provides milliarcsecond accuracy, and often contains more sources within the FoV than LCS1 or ATPMN, though not enough for stand-alone use. Therefore, we first use RFC to correct the RACS-mid source positions locally over a large FoV of 6° in radius, then apply these corrected positions to perform the astrometric correction of the MeerTRAP images. We select RACS-mid sources classified as point-like with positional uncertainties < 0.5 arcsec in both right ascension (RA) and declination (Dec.), and we perform the astrometric correction with the Python module `astroalign` (M. Beroiz, J. B. Cabral & B. Sanchez 2020).

In order to avoid astrometric distortions from sources that are far from the phase centre of the MeerKAT image, we restrict our reference source selection to a certain radius. If the FRB position is contained within the half-power beamwidth (HPBW) at the central frequency of the detection bandwidth, we use the sources contained within the HPBW (See M. S. Villiers & W. D. Cotton 2022; M. S. Villiers 2023, for HPBW measurements). If the source is outside the HPBW, we use sources contained within $4/3$ of the angular separation between the FRB and the phase centre.

Once we have selected our reference and MeerKAT sources, we match them using a 5 arcsec tolerance. We then compute an affine transformation matrix by performing a least squares minimization of the positional offsets between the matched MeerKAT and reference sources. We weight the minimization by the inverse of the positional variance, where we compute the variance as the quadrature sum of the positional errors from both the reference and the MeerKAT catalogues. Finally, we apply the resulting transformation matrix to the MeerKAT sources and images using the the python module `scikit-image` (S. v. d. Walt et al. 2014).

The resulting source position thus has several sources of uncertainty, and the final error we report is obtained by summing these in quadrature. The first source of error is the positional uncertainty obtained from `PyBDSF`, and the second one is from the astrometric correction. We estimate the astrometric error by computing the average separation in RA and Dec. between the

corrected and reference sources after each alignment, and add them in quadrature. For VLASS, we add an additional 0.2 arcsec systematic error in RA and Dec. to account for the typical positional source accuracy. This error budget explicitly accounts for potential residual systematics in the astrometry, ensuring that the reported uncertainties are conservative and that our host associations are robust. The final ‘on’ and ‘off’ images are shown in Fig. 1, while the 1σ localization regions overlaid on the optical background are shown in Fig. C1.

3.2 Localization of FRBs without TB data

Although the MeerTRAP search pipeline runs in near-real-time (K. Rajwade et al. 2021; K. M. Rajwade et al. 2022), only the bursts identified within 45 s of the time they occurred will trigger a TB dump. In some cases the identification can be delayed beyond this limit, most commonly when an excess of candidate transients slows the clustering process, although other factors may also contribute. A full investigation will be presented in future work. The bursts that did not produce a TB trigger thus have to be localized through different techniques. The python package `SeeKAT`² was developed to localize MeerTRAP and TRAPUM sources by combining the pointing direction of the beams where they were detected and the S/N values in each one through a maximum likelihood estimation approach. This method is described in detail in M. C. Bezuidenhout et al. (2023), and it has already been tested for previous MeerTRAP localizations (e.g. K. M. Rajwade et al. 2022; F. Jankowski et al. 2023). While a bright source detected in many CBs can be localized with (sub-)arcsecond accuracy, the position of a source detected in fewer than three CBs or in the IB, in most cases, will not be constrained enough to determine its most likely host galaxy.

Some FRBs do not produce enough CB detections to achieve a good localization with `SeeKAT`. For bright FRBs, an alternative method is to localize them using imaging data from the observations running commensally to MeerTRAP, where the shortest available integration times are either 8 or 2 s. This technique was successfully used to localize the MeerTRAP FRBs 202104051 (8 s integration time, L. N. Driessen et al. 2024) and 20210410D (2 s integration time, M. Caleb et al. 2023). However, since the typical millisecond duration of FRBs is orders of magnitude shorter than the available integration times, and the dispersion delay can smear the signal across multiple time intervals, which have fixed start times, only sufficiently bright FRBs can be localized through this method. To mitigate the signal loss, we select the snapshots containing the burst, and restrict the frequency range when appropriate to maximize S/N, though the reduced signal due to temporal averaging and dispersion cannot be fully corrected.

3.3 Host galaxy observations and analysis

Once we have localized an FRB with (sub-)arcsecond precision, we search for archival optical images at the burst location to identify any host galaxy candidates. The 10th data release (DR10) of the DESI Legacy Survey (DESI hereafter, A. Dey et al. 2019) covers the coordinates of most MeerKAT pointings. Since it can reach magnitude depths $\lesssim 25$ in the *g* and *r* bands, it is likely to contain the host galaxies of most of the FRBs presented in this paper. The DESI DR10 source catalogues provide source fluxes in the *g*, *r*, *i*,

¹RFC: <https://astrogéo.org/rfc/>

²SeeKAT: <https://github.com/BezuidenhoutMC/SeeKAT>

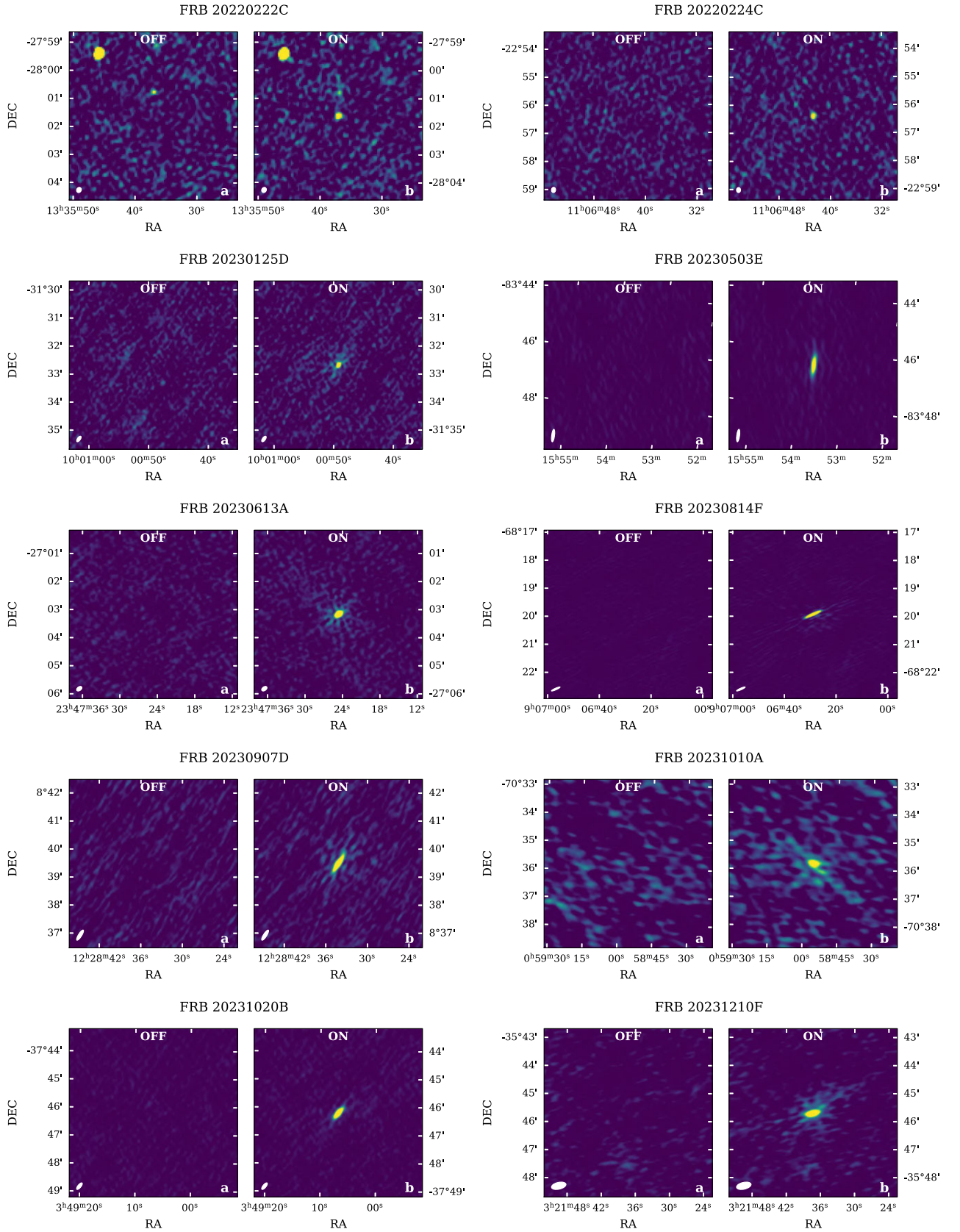


Figure 1. Localization images of the FRBs presented in this work. In each row, the left panel shows an image before the arrival time of the burst (off), and the right panel the image during the burst (on). The images are centred at the FRB location, and have an FoV of 6 arcmin \times 6 arcmin. The beam size of each observation is shown on the bottom left corner of each panel as an ellipse.

and z bands from the dark energy camera (DECam; B. Flaugher et al. 2015), as well as in the W1-4 infrared filters from the Near-Earth Object Wide-field Infrared Survey Explorer (NEOWISE; A. Mainzer et al. 2014) when available.

3.3.1 Optical data acquisition

Although archival optical images are generally available, for some FRB coordinates there were no observations deep enough to identify the most likely host galaxy. In those cases, we observed the FRB locations with ground-based optical telescopes, with observing time granted through various proposals to use Gemini/GMOS, Gemini/Flamingos2, Keck/DEIMOS, Keck/LRIS, and SOAR/Goodman; the program details of specific observations are listed in Table D2. After identifying the most likely FRB host galaxies using the Probabilistic Association of Transients to their Hosts (PATH; K. Aggarwal et al. 2021, see Section 3.3.2), we obtained spectra to identify their redshifts, using the same observatories as for the deep optical images.

The Gemini observatory consists of two optical/NIR telescopes with a diameter of 8.1 m, with Gemini-North located in Hawaii and Gemini-South in Chile, which together cover most of the Northern and Southern sky. Gemini-South is especially well suited to follow up any potential MeerTRAP source. Both telescopes are equipped with the Gemini Multi-Object Spectrograph (GMOS) instrument, which provides imaging over a $5.5 \text{ arcmin} \times 5.5 \text{ arcmin}$ FoV with a spatial sampling of $0.0807 \text{ arcsec pixel}^{-1}$ (I. M. Hook et al. 2004; G. Gimeno et al. 2016). Imaging can be performed using the Sloan Digital Sky Survey (SDSS) filter system (u' , g' , r' , i' , and z'). By default, we used the r -band filter to perform deep imaging observations and identify any potential host galaxy at the FRB location. GMOS further supplies 360 to 1030 nm spectroscopy, either in long-slit mode with a 5.5 arcmin long slit, or in multi-object mode with up to 60 custom-made slits.

Additionally, Gemini-South is currently equipped with the FLoridA Multi-object Imaging Near-infrared Grism Observational Spectrometer 2 (FLAMINGOS-2), a near-infrared imager and spectrograph (S. Eikenberry et al. 2006). In imaging mode, it covers a circular 6.1 arcsec FoV with a spatial sampling of $0.179''/\text{pixel}$, with the J, H, Ks, K-blue, and K-red filters available. For spectroscopic observations, it has both a long-slit mode and a multi-object spectroscopy mode, covering a wavelength range from 950 to 2500 nm with a spectral resolution $R = 1000 - 3000$.³

The W. M. Keck observatory (Keck hereafter) consists of two telescopes (I and II) with a 10 m diameter each, located in Hawaii. The Low Resolution Imaging Spectrometer (LRIS) is an instrument to perform optical imaging and spectroscopy installed on Keck I. The light received by the telescope is split into two and sent separately to a red and a blue camera (J. B. Oke et al. 1995; C. Rockosi et al. 2010; M. F. Kassis et al. 2022). The imaging mode on Keck I includes the UBVGRI filters; the FoV is $6 \text{ arcmin} \times 7.8 \text{ arcmin}$, and the pixels are 0.135 arcsec wide. We typically obtain images with the R -band filter. Spectra are obtained with a grating on the red side and a grism on the blue side, which together cover the wavelength range between 320 and 1000 nm. The different gratings and grisms that can be used result in different wavelength resolutions, roughly from 0.16 to 0.92 nm. The spectra are produced either with 175 arcsec long slits

of different widths, or with up to 100 slits in the multi-object mode using custom-made slit masks.⁴

The DEep Imaging Multi-Object Spectrograph (DEIMOS) instrument is a visible-wavelength, faint-object, multislit imaging spectrograph installed on the Keck II telescope. In imaging mode, DEIMOS has a FoV of $16.7 \text{ arcmin} \times 5.0 \text{ arcmin}$, with a pixel size of 0.1185 arcsec . It utilizes BVRIZ broad-band filters⁵ to obtain deep images from ~ 400 to ~ 1000 nm. To identify any potential host galaxy at the FRB location, we used the R -band filter by default (S. M. Faber et al. 2003). In spectroscopic mode, DEIMOS uses either custom-made slit masks to observe multiple objects simultaneously, with up to 130 slits per mask, or 20 arcsec long slits with various different widths. It covers a wavelength range between 410 and 1100 nm, and a spectral resolution between 0.11 and 0.35 nm depending on the selected grating.⁶

The Southern Astrophysical Research (SOAR) Telescope with a 4.1 m mirror is located in Cerro Pachón, Chile. The Goodman high throughput spectrograph, mounted on SOAR, can be used both for imaging and spectroscopy (J. C. Clemens, J. A. Crain & R. Anderson 2004). In imaging mode, it is equipped with the SDSS filters u , g , r , i , and z , as well as the Bessell filters U, B, V, R, and I, and an FoV of $7.2 \text{ arcmin} \times 7.2 \text{ arcmin}$ with a spatial sampling of $0.15 \text{ arcsec pixel}^{-1}$. In spectroscopic mode, both a long slit mode and multislit object masks over a field of $3 \text{ arcmin} \times 5 \text{ arcmin}$ are available, and it covers a wavelength range from 320 to 900 nm with a spectral resolution $R = 1400 - 14000$.⁷

All instruments thus possess similar characteristics, and the selection of the instruments will mainly be dependent on the declination of the target and on the available observing time.

We reduced the imaging data acquired with Gemini and Keck using the Pipeline for Optical/infrared Telescopes in Python for Reducing Images (POTPyRI).⁸ The pipeline first creates master bias, dark, and flat frames from the available calibration files, which are then applied to the science frames before aligning and stacking them. Next, it applies an astrometric correction using the Gaia DR3 catalogue (Gaia Collaboration 2023) for optical observations, or the Two Micron All Sky Survey (2MASS M. F. Skrutskie et al. 2006) for near-infrared observations. In cases where the positional discrepancy between the sources in the image and the reference catalogue is $\geq 0.5 \text{ arcsec}$, we first apply a manual WCS alignment. Once the astrometric solution is successful, we calculate the zero-point magnitude of the image and perform aperture photometry. This is done with the Python packages `photutils` (L. Bradley et al. 2024) and `petrofit` (R. Geda et al. 2022) to detect sources, perform aperture photometry, and determine their angular radii. We then cross-match these sources with reference catalogues that provide calibrated magnitudes in the FoV, such as Pan-STARRS (K. C. Chambers et al. 2019) or SkyMapper (C. A. Onken et al. 2024). These matches are used to compute the zero-point magnitude for each matched source, from which we derive the magnitudes of all detected objects using the median zero-point. The magnitudes reported in these two surveys

³Gemini/Flamingos2 specifications: <https://noirlab.edu/public/programs/gemini-observatory/gemini-north/flamingos-2/>

⁴Keck/LRIS specifications: https://www2.keck.hawaii.edu/inst/lris/pre_observing.html

⁵Keck/DEIMOS filters: www2.keck.hawaii.edu/realpublic/observing/public_instrument_info/deimos/filter_list.html

⁶Keck/DEIMOS specifications: https://www2.keck.hawaii.edu/inst/deimos/pre_observing.html

⁷SOAR/Goodman specifications: <https://noirlab.edu/public/programs/ctio/soar-telescope/goodman/>

⁸POTPyRI: <https://github.com/CIERA-Transients/POTPyRI>

are consistent to within $\lesssim 0.02$ mag in overlapping filters, with Pan-STARRS showing systematics of only 7–12 mmag across the sky.⁹ Any systematic offsets between the calibrations are hence negligible compared to our measurement uncertainties.

To reduce the SOAR imaging observations, we used the Phot-pipe pipeline¹⁰ instead (A. Rest et al. 2014). Photpipe performs image calibration, including bias subtraction and flat-fielding, followed by astrometric calibration using Gaia DR3 as a reference. The fits images are next sky-subtracted, stacked, and re-gridded to the same pixel scale and field centre using SWarp¹¹ (E. Bertin et al. 2002). Finally, it performs a photometric calibration, computing the zero-point magnitude using again either Pan-STARRS or SkyMapper as a reference, which is applied to the aperture photometry.

The uncertainty in the calibrated magnitude for each source combines the photometric measurement error with the zero-point uncertainty. The measurement error is computed from the flux uncertainty reported by photutils, while the zero-point uncertainty is estimated from the scatter of matched reference catalogue sources. These two contributions are added in quadrature to obtain the total magnitude uncertainty. The contribution from the reference catalogues themselves is negligible compared to the other sources of uncertainty.

3.3.2 Host galaxy identification

In order to identify the FRB hosts, we apply PATH (K. Aggarwal et al. 2021) in our optical/NIR images around the FRB positions. PATH is a Bayesian framework for identifying the most likely host galaxy of the FRB by computing priors using the magnitudes and angular separations of all candidate host galaxies, along with the probability that the host galaxy is not visible (unseen) in the image. These priors are then combined with the FRB localization region to calculate the posterior probabilities, identifying the most probable host.

If the FRB localization is covered by DESI, the survey’s optical depth is typically enough to identify a sample of the nearest galaxies. In cases where we obtained deeper optical imaging, we apply a source fitting technique to extract the properties of sources near the FRB location. Although the DESI photometric catalogues provide the magnitudes and angular radii in up to eight filters (griz, W1–4), along with photometric redshifts in regions outside of the Galactic plane (A. Dey et al. 2019), we process the r -band DESI images with the same pipeline used for the targeted FRB fields to ensure homogeneity across the whole sample. Whether we use the DESI catalogue or the sources identified in a targeted optical observation, we select the positions, half-light radii and magnitudes of the extended sources within a 10 arcsec separation from the FRB localization to generate the input source catalogue for the PATH analysis. We exclude from the candidate list the sources identified as stars in the DESI catalogue, or identified as point-like by the fitting technique.

The PATH algorithm requires the unseen host prior, $P(U)$, as an input. We estimate this prior independently for each FRB, following an approach similar to that described by L. Marnoch et al. (2023, appendix B2). To define a plausible redshift range for each FRB, we begin by estimating the extragalactic DM. We assess the Galactic interstellar medium (ISM) contribution by averaging the expected

values from the NE2001 (J. M. Cordes & T. J. W. Lazio 2003) and YMW16 (J. M. Yao, R. N. Manchester & N. Wang 2017) models, and then add the Galactic halo contribution from the S. Yamasaki & T. Totani (2020) model. Subtracting the combined Galactic contribution from the observed DM yields the extragalactic DM. We then apply the Macquart relation (J.-P. Macquart et al. 2020) using the FRBs package (J. X. Prochaska et al. 2025), to obtain the 95 percent confidence redshift range expected for each FRB. We assume a host galaxy contribution to the DM in its rest frame of 100 pc cm^{-3} , which is consistent with empirical estimates and simulations (J.-F. Mo et al. 2023; T. O. Kovacs et al. 2024; L. Bernales-Cortes et al. 2025). Using a reference sample of 56 securely localized FRBs with known r -band magnitudes and spectroscopic redshifts, we compute how their apparent magnitudes would evolve if they were located at different redshifts. For each trial redshift, we construct a probability density function using a Gaussian kernel density estimate (KDE) based on the projected magnitudes. Then, for each FRB, we determine the limiting magnitude of the optical observation used to identify host galaxy candidates. Using the KDE corresponding to the Macquart redshift upper limit of each FRB, we compute the fraction of reference galaxies that would fall below the detection threshold. We adopt this fraction as the unseen prior $P(U)$. After running PATH, we identify any galaxy with a posterior probability of association $P(O_i|x) > 0.9$ as the putative host.

The estimate of $P(U)$ based on a DM-derived redshift upper limit, magnitude completeness, and the luminosity of known FRB hosts, provides a conservative prior compared to adopting a fixed 5–10 percent value, reducing arbitrary assumptions in cosmological applications. Our association probabilities are largely insensitive to moderate changes in the assumed host DM, because of DM dilution with redshift, simulations suggesting that host DMs tend to be larger at higher redshift (T. O. Kovacs et al. 2024), and our conservative unseen prior. However, $P(U)$ might be biased if FRB hosts are systematically fainter than the known sample, or if the Macquart relation underestimated redshifts along certain lines of sight.

Certain FRB sightlines, especially for FRBs with high DMs, may intersect foreground galaxy clusters that contribute significantly to the observed DM. In order to estimate whether these could affect the inferred redshift, we searched for clusters within 2° of each FRB location. For our search, we selected the galaxy cluster catalogue from the DESI Legacy Imaging Surveys (H. Zou et al. 2021), which offers the most extensive optical coverage available, and the second release of the meta-catalogue of X-ray-detected clusters of galaxies (MCXC-II; T. Sadibekova et al. 2024), which provides the most complete compilation of X-ray clusters. In Section 4, these searches are mentioned wherever we find clusters intersecting the FRB sightlines.

For each cluster intersecting the FRB sightline within twice its characteristic radius, R_{500} , we estimated the DM contribution by modelling their gas density profiles with a β -model (e.g. M. Arnaud 2009). We assume typical values of $\beta \sim 2/3$, a core radius $r_c \sim 0.15R_{500}$, and a gas fraction $f_{\text{gas}} \sim 0.1$ of the total mass. Next, we integrate the electron density at the projected impact parameter between the cluster centre and the FRB, up to $2R_{500}$. Finally, we apply a redshift correction to the resulting DM. While the values above are typical for X-ray detected clusters, their large uncertainties result in errors of up to 30 percent of the reported values. This simple calculation is only intended to verify that the observed host redshift falls within the expected Macquart range; a detailed analysis of the foreground cluster contributions will be reported in an upcoming publication.

⁹Pan-STARRS photometric quality: https://skymapper.anu.edu.au/data-release/dr1/?utm_source=chatgpt.com#dr1p1-photometric-quality

¹⁰Photpipe: <https://photpipe-docs.readthedocs.io/en/latest/>

¹¹SWarp: <https://www.astromatic.net/software/swarp/>

Table 1. FRB localization general properties.

ID	DM (pc cm ⁻³)	DM _{NE2001} (pc cm ⁻³)	RA (hh:mm:ss)	Dec. (dd:mm:ss)	z_{spec}	z_{phot}	Localization method
20220222C	1071.2(0.8)	56	13:35:37.08±0.54''	−28:01:36.93±0.55''	0.853	–	8-s image
20220224C	1140.2(1.8)	52	11:06:42.61±0.42''	−22:56:23.48±0.64''	0.6271	0.57 ^{+0.06} _{−0.06}	8-s image
20230125D	640.08(0.03)	88	10:00:49.21±0.24''	−31:32:40.77±0.26''	0.3265	–	TB
20230306F	689.5(0.9)	23	12:24:01±60''	+14:54:17±50''	–	–	SeeKAT
20230413C	1532.2(0.05)	45	05:13:00±63''	−39:50:00±57''	–	–	SeeKAT
20230503E	483.74(0.04)	88	15:53:43.19±0.69''	−83:46:30.93±0.87''	–	0.32 ^{+0.15} _{−0.15}	TB
20230613A	483.51(0.01)	30	23:47:24.65±0.4''	−27:03:10.01±0.52''	0.3923	0.42 ^{+0.03} _{−0.03}	TB
20230814F	471.44(0.03)	134	09:06:28.88±0.35''	−68:19:55.70±0.31''	–	–	TB
20230827E	1433.7(0.1)	38	04:08:28.242±0.75''	−18:16:58.55±1.5''	–	–	SeeKAT
20230907D	1030.79(0.04)	29	12:28:34.20±0.4''	+ 08:39:29.13±0.57''	0.4638	0.44 ^{+0.03} _{−0.02}	TB
20231007C	2660.4(1.9)	42	23:38:18±50''	+ 21:52:34±55''	–	–	SeeKAT
20231010A	442.59(0.02)	41	00:58:55.67±0.52''	−70:35:46.93±0.3''	–	0.61 ^{+0.18} _{−0.18}	TB
20231020B	952.2(0.3)	34	03:49:06.77±0.37''	−37:46:11.56±0.4''	0.4775	0.46 ^{+0.07} _{−0.08}	TB
20231204B	1772.1(0.3)	41	01:05:07±64''	−70:37:16.5±68''	–	–	SeeKAT
20231210F	720.6(0.2)	32	03:21:37.28±0.25''	−35:45:41.13±0.25''	–	0.50 ^{+0.08} _{−0.08}	TB

Notes. Additional details about the FRB host galaxies and their physical properties are listed in the appendix Tables D1 and D3. DM_{NE2001} is the dispersion measure from the NE2001 model evaluated at 30 kpc, z_{spec} is the spectroscopic redshift, and z_{phot} the photometric redshift.

3.3.3 Analysing spectroscopic and photometric data

The spectra are obtained once the putative host has been identified. If more than one host galaxy candidate remains after running PATH, we use either the long-slit or the multi-object spectroscopy slit masks available with Gemini-S/GMOS, Keck/LRIS, or Keck/DEIMOS to obtain simultaneous spectra of those sources. Once the spectra have been obtained, we reduce them with PyPeIt using standard practices (J. Prochaska et al. 2020), and determine the redshift.

For some FRBs, the observation constraints have not allowed us to obtain a spectrum yet. In those cases, we obtain a photometric redshift if enough photometry is available in at least four bands. The DESI-DR10 catalogue provides photometric redshifts computed through a random forest algorithm using DECam and WISE fluxes for all galaxies with more than one exposure in the required filters, and contained within unmasked regions.¹²

To infer the physical properties of the most likely host galaxies of our FRB sample, we use the Code Investigating GALaxy Emission (CIGALE; M. Boquien et al. 2019). Using the redshift and the photometric data as input, CIGALE fits the spectral energy distributions (SEDs) to estimate the stellar mass, SFR, metallicity, the age of the oldest stars, and other properties that characterize the host galaxies. We initialize CIGALE adopting the following models: delayed-exponential star formation history without burst population, a synthetic stellar population described by G. Bruzual & S. Charlot (2003), the initial mass function given by G. Chabrier (2003), dust attenuation models from D. Calzetti (2001), and dust emission template from D. A. Dale et al. (2014). This upper limit follows the prescription used in previous FRB host studies (I. S. Khrykin et al. 2024) and is consistent with the range expected from the D. A. Dale et al. (2014) templates for typical star-forming galaxies. In cases where SED residuals or mid-IR photometry suggest a different active galactic nucleus (AGN) contribution, we adjust this fraction to obtain more reliable estimates on the stellar population parameters.

CIGALE provides Bayesian estimates of the physical parameters and their associated uncertainties based on the likelihood-weighted average and standard deviation of their probability distribution function derived from a model grid. However, these estimates rely on a

good photometric coverage, with 10 to 25 photometric measurements across a broad wavelength range typically required to get robust constraints on the galactic properties. When spectroscopic redshifts are unavailable, CIGALE can also infer photometric redshifts. Since we verified that the spectroscopic redshifts in our sample are consistent with the photometric redshifts reported on the DESI DR10 catalogue, we limited the redshift range that CIGALE explores to within those uncertainties when possible (see Appendix B). In cases where no photometric redshift is available, we fit for a photometric redshift between 0 and 1. We note however that the uncertainties on the derived physical parameters are significantly larger for galaxies lacking spectroscopic redshifts. While other tools, such as Prospector (B. D. Johnson et al. 2021), can perform joint photometric and spectroscopic SED fits, such analysis is beyond the scope of this work. Moreover, recent comparisons between the properties derived by CIGALE and Prospector for FRB host galaxies (A. C. Gordon et al. 2023) show that the resulting SFR estimates are in good agreement (private communication), justifying the reliability of our CIGALE results. The photometric information and CIGALE results are summarized in Table D3.

4 RESULTS

During all of the MeerTRAP observations that were carried out throughout the year 2023, we discovered 14 new, so far non-repeating, FRBs. One of the bursts, FRB 20230808F, has already been presented in K. Y. Hanmer et al. (2025). In this work, we present the localization of the remaining thirteen 2023 FRBs, as well as the localization of two FRBs discovered in February 2022 before the TB data became available. Of the total 15 localized FRBs, two lie too close to the Galactic plane to be covered by DESI-DR10 imaging. Among the 11 FRBs with arcsecond localizations and DESI coverage, two have no visible host galaxy candidates in the DESI data, consistent with the survey’s imaging depth. The remaining four FRBs have arcminute localizations; although these regions are included in DESI-DR10, their large sizes prevent the identification of a single host galaxy. General FRB and localization properties are listed in Table 1, while detailed properties of the host galaxy candidates are shown in Table D1.

¹²Legacy survey files: <https://www.legacysurvey.org/dr10/files/>

4.1 Arcsecond localized FRBs

4.1.1 FRB 20220222C

FRB 20220222C was discovered while MeerTRAP was observing commensally in the L band with the MHONGOOSE (MeerKAT HI Observations of Nearby Galactic Objects – Observing Southern Emitters, W. J. G. Blok et al. 2024) large survey project (Proposal ID: SCI-20180516-EB-01), with a DM of $1071.2 \pm 0.8 \text{ pc cm}^{-3}$. The dynamic spectrum showed the burst to be present only in the lower half of the frequency band. It was detected prior to the successful implementation of the TB system, and there is thus no available TB trigger. However, it was detected in seven CBs and the IB with a S/N of 28 in the IB, and was hence reasonably well localized. The burst was identified in the 8-s images generated from the MHONGOOSE observation, by searching for new sources appearing at the burst time of arrival (ToA) within the coordinates the detection CBs were pointing at.¹³ Since the burst was only visible at low frequencies, we created images of the bottom half of the band. For these data, we found a new source in the difference images at the coordinates $13^{\text{h}}35^{\text{m}}37.08^{\text{s}} \pm 0.54'' -28^{\circ}01'36.93'' \pm 0.55''$, where the errors include the PyBDSF and the astrometric error on the source position, for which we used VLASS as a reference. The off/on images are shown in Fig. 1.

A search in the DESI images at the source position revealed no host galaxy candidate within the error region, and we thus acquired deep images in the r , z , J , and K_s bands with Gemini-South, with total exposure times of 2500, 2500, 600, and 1800 s respectively. The observation details are summarized in Table D2.

In the resulting images, we identified a galaxy (G1) centred 1.43 arcsec away from the FRB coordinates with a magnitude $r = 23.86 \pm 0.04$. This galaxy has a PATH probability of 94.4 per cent of being associated with the FRB after assuming an unseen prior of 5 per cent, and we therefore identify it as the putative host. A second source with magnitude $r = 25.01 \pm 0.07$ is located 3.08 arcsec away from the FRB coordinates (G2), and a third brighter galaxy with $r = 23.32 \pm 0.04$ is 3.81 arcsec away (G3). Both those galaxies have association probabilities < 0.01 per cent. A composite RGB image from the r , z , and J images, the FRB error region, and candidate host galaxies are shown in Fig. 2.

To measure the redshift of the two brightest galaxies, G1 and G3, we obtained a spectrum with Keck/LRIS in long-slit mode on 2023 April 17 (Program ID: U173; PI: X. Prochaska). The red-side grating was centred at 808.2 nm with the 600/10000 grating, with a total exposure time of 6450 s, while the blue-side grism used the 300/5000 configuration and the exposure time was 6600 s. No spectral lines were identified in the blue grism for the putative host. On the red grating, the $H\beta$ line and the [O III] doublet were identified in the spectrum, which allowed us to determine a redshift of 0.853. This is in agreement with the redshift range expected from the Macquart relation, $z_{\text{Macquart}} = 1.17^{+0.26}_{-0.58}$. The projected distance between the galaxy centre and the FRB location would thus be $11.3 \pm 6.1 \text{ kpc}$. For G3, we identified a feature in the blue-side spectrum which might be the [O II] doublet, and would correspond to a redshift of 0.795, but since no lines were identified in the red-side spectrum, the measurement is not convincing.

Based on the magnitudes of G1, that was detected in the r , z , and J filters, but not in the K_s filter, we performed an SED fit. We find the value of the stellar mass to be $\log(M_*/M_{\odot}) = 10.1 \pm 0.2$.

¹³This search was undertaken under the auspices of the ThunderKAT project R. Fender et al. (2016).

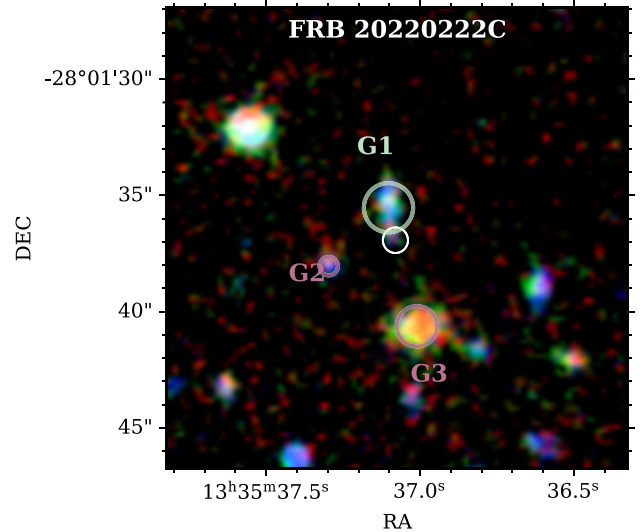


Figure 2. Localization of FRB 20220222C. The error region is shown as a white ellipse, and the background image shows a composite image with the r -, z -, and J -band observations. The candidate hosts are indicated by circles of half-light radii labelled G1-3.

The best SFR value is $7.2 M_{\odot} \text{ yr}^{-1}$, with a 0.3 dex error, and the metallicity $\log(Z/Z_{\odot}) < -0.13$. The oldest stars in the galaxy have an age of $1.8 \pm 0.7 \text{ Gyr}$. The CIGALE fit is shown in Fig. D2.

4.1.2 FRB 20220224C

This FRB was detected with a DM of $1140.2 \pm 1.8 \text{ pc cm}^{-3}$ in another MHONGOOSE observation (Proposal ID: SCI-20180516-EB-01), prior to the successful implementation of the TB system. Since it was found only in the IB, SeeKAT could not constrain the position well. However, the burst signal was strong enough to be localized in the 8-s images of the commensal MHONGOOSE observation¹³. Due to the dispersion delay, the burst signal was divided across two consecutive 8-s images. The on image, which adds together the two 8-s images where it was detected, is shown together with an off image in Fig. 1.

We performed an astrometric correction of the final on image using VLASS as a reference, and obtained the final coordinates $11^{\text{h}}06^{\text{m}}42.61^{\text{s}} \pm 0.42'' -22^{\circ}56'23.48'' \pm 0.64''$. We identified a host galaxy candidate in DESI DR10 located $0.52''$ away from the position and with an r -band magnitude of 21.8 ± 0.3 . Furthermore, we acquired deep r and i -band images with GMOS, with total exposure times of 2700 s each, on 2023 January 29 and 30 (Proposal ID: GS-2022B-Q-123; PI: A. Gordon). A composite colour image is shown in Fig. 3.

Based on the DM of this FRB, after removing the MW contribution and 100 pc cm^{-3} from the host, the expected Macquart redshift would be $z_{\text{Macquart}} = 1.27^{+0.28}_{-0.62}$. Assuming an unseen prior of 67 per cent, we find its PATH association probability to be 93.8 per cent, with the probability of association of all surrounding sources being negligible. In order to measure the redshift, we obtained LRIS spectra of the putative host as well as two other nearby sources. The spectra were acquired on 2023 January 24 (Program ID: U055; PI: J. Hennawi) with the 300/5000 blue grism and an exposure of 2780 s, and with the 600/10000 red grating centred at 892.7 nm and an exposure time of 2740 s. In the spectra, we identified the [O II] and [O III]

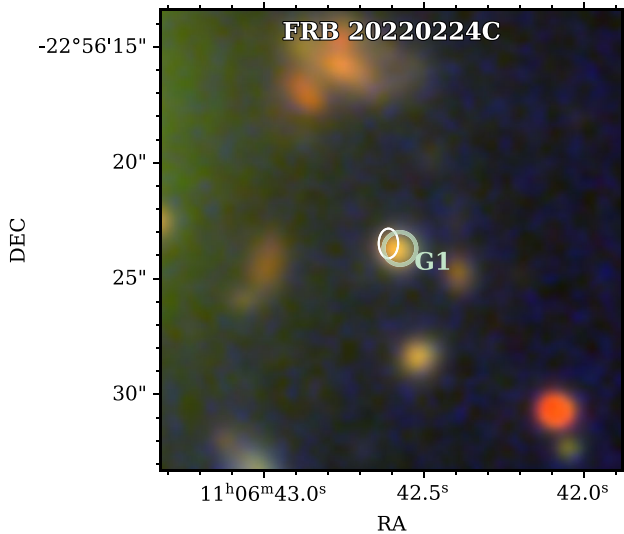


Figure 3. Localization of FRB 20220224C. The error region is shown by a white ellipse, and the background is a composite image from the r and i observations and the g image from DESI-DR10. The candidate host galaxy is indicated by a circle labelled G1.

doublets, as well as the $H\beta$ line, which we used to measure the FRB host redshift to be 0.6271. At this redshift, the projected physical separation between the FRB and the galaxy centre is 3.8 ± 5.4 kpc.

The spectroscopic redshift hence falls below the 95 per cent lower limit. To assess whether any structure in the foreground could be significantly contributing to the observed DM, we searched for galaxy clusters in the foreground. From the DESI Legacy Imaging Surveys, H. Zou et al. (2021) produced a galaxy cluster catalogue, and there we found the galaxy cluster ACO S 651 to be located 550 arcsec away from the FRB location, at a redshift of 0.0639. The cluster is also detected in X-rays (R. Piffaretti et al. 2011), and it has a characteristic radius $R_{500} = 830$ kpc. The impact parameter of the FRB, ~ 700 kpc, falls within its radius. From a cluster with a typical density profile, as described in Section 3.3.2, the expected DM contribution could be up to 800 ± 100 pc cm^{-3} in the cluster frame, or ~ 750 pc cm^{-3} after correcting for redshift. This contribution likely explains the large DM observed.

Through the DESI-DR10 source catalogues,¹⁴ we had access to photometric measurements of the host galaxy candidate in a broad wavelength range. These include the *griz* filters from DECam and the NEOWISE filters W1-4. This extensive wavelength coverage enables us to perform SED fitting of the putative host galaxy with CIGALE. The SED parameters derived for the putative host are as follows: a stellar mass $\log(M_*/M_\odot) = 10.3 \pm 0.1$, a star formation rate $\text{SFR} = 18.9 M_\odot \text{ yr}^{-1}$, with a 0.2 dex error, a metallicity $\log(Z/Z_\odot) < -0.7$, and an age of 1.4 ± 0.5 Gyr. The CIGALE fit is shown in Fig. D2.

4.1.3 FRB 20230125D

This FRB, with a DM of 640.08 ± 0.03 pc cm^{-3} , was detected in the L band during an open time observation (Proposal ID: SCI-20220822-IP-01). It was detected in the IB and the CB 261, and its detection triggered the storage of the TB data. We localized the burst in the TB data as described in Section 3.2, and performed an astrometric correction with VLASS as the reference catalogue.

¹⁴DESI-DR10: <https://www.legacysurvey.org/dr10/description/>

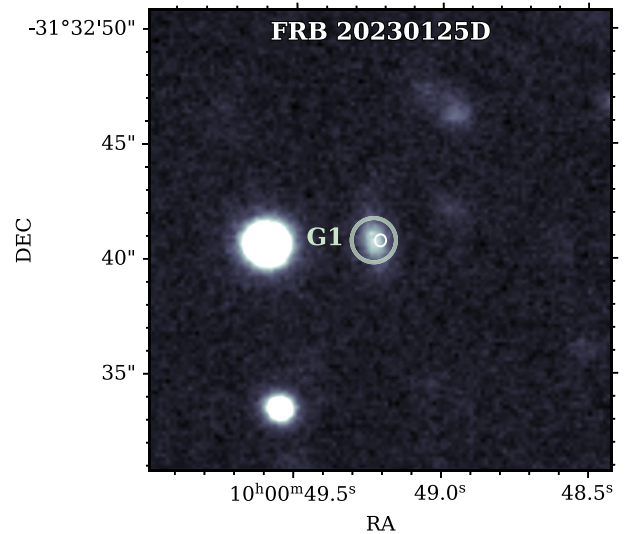


Figure 4. Localization of FRB 20230125D. The error region is shown by a white ellipse, and the background image is a Gemini-S observation. The candidate host galaxy is indicated by a coloured circle labelled G1.

The on and off images where the burst is found are shown in Fig. 1. We localized the burst to the coordinates $10^{\text{h}}00^{\text{m}}49.21^{\text{s}} \pm 0.24'' - 31^{\circ}32'40.77'' \pm 0.26''$. The FRB coordinates were not covered by DESI, and it only had been observed in the i band with Pan-STARRS. We identified a potential source at the location of the FRB in the i band, and followed up with a deep r -band image using the GMOS with a total 2700 s exposure time (see observation details in Table D2). In the deep image, we identified a galaxy 0.29 arcsec away from the FRB localization centre with a magnitude $r = 22.16 \pm 0.02$, as shown in Fig. 4. Based on the image magnitude limit and the Macquart redshift upper limit, we assume an unseen prior of 36 per cent and obtain a posterior \mathcal{P} ATH probability of 97.81 per cent, and we hence identify this as the host galaxy. The deep r -band image is shown in Fig. C1. Subsequently, we obtained a spectrum of the putative host using Keck/DEIMOS (see Table D2). In the spectrum, we identified the $H\alpha$ and $[\text{N II}]$ spectral lines (see Fig. D1), from which we measured a redshift of $z = 0.3265$, which falls on the lower end of what is expected from the Macquart relation ($z_{\text{Macquart}} = 0.62^{+0.14}_{-0.34}$). To explain its apparent DM excess, we searched for known galaxy clusters in the foreground, but found no matches in MCXC-II. For the same reason why no DESI images are available at this location, we do not have information about galaxy clusters identified in the optical. Future observations might explain whether the excess DM can be attributed to foreground clusters or to the host galaxy instead. At the spectroscopic redshift, the projected distance between the galaxy centre and the FRB location is 1.4 ± 1.7 kpc. Because we only have the host galaxy magnitude in one filter, we are not able to perform a CIGALE SED fitting at this stage.

4.1.4 FRB 20230503E

FRB 20230503E was discovered during a gain calibration observation in the L band of the calibrator J1619-8418 (Proposal ID: SCI-20180516-PW-03), and it has a DM of 483.74 ± 0.04 pc cm^{-3} . The burst was only detected in the IB, and it triggered the storage of the TB data, which we used to localize it. After obtaining the original coordinates, we performed an astrometric correction using

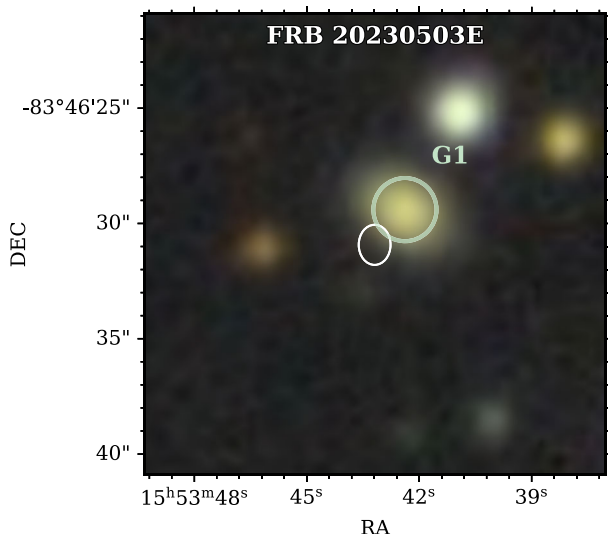


Figure 5. Localization of FRB 20230503E. The error region is shown by a white ellipse, and the background image is from DESI-DR10. The candidate host galaxies are indicated by coloured circles labelled G1-2. The source inside the FRB error region is a star.

the ATPMN sources in a 6° radius to correct the RACS-mid source coordinates.

We determined the final localization region to be $15^{\text{h}}53^{\text{m}}43.19^{\text{s}} \pm 0.69''$ $-83^\circ46'30.93'' \pm 0.87''$. We identify a galaxy located $2.02''$ away from the FRB localization, with an r-band magnitude of 20.1 ± 0.2 , as can be seen in Fig. 5. Assuming an unseen prior of 5 per cent, we obtain a PATH probability of association of 97.3 per cent, and we thus identify this galaxy as the putative host.

Although the DESI images of these coordinates exist, this region was masked on the DESI source detection pipeline, and thus the host galaxy candidates are not listed in the source catalogue. We used the computed DECam *griz* magnitudes and the WISE W1 and W2 magnitudes to perform an SED fit of the host galaxy with CIGALE. Because we do not have a spectroscopic or photometric redshift constraints for this galaxy, its physical properties are poorly constrained. CIGALE finds a loosely constrained redshift $z_{\text{phot}} = 0.29 \pm 0.26$, and subsequent parameters $\log(M_*/M_\odot) < 10.6$, $\text{SFR} < 26 M_\odot \text{yr}^{-1}$, $\log(Z/Z_\odot) < 0.2$, and age 2 ± 1 Gyr. Once a spectroscopic redshift is obtained for this FRB, refined SED fit results will be reported in a future publication.

4.1.5 FRB 20230613A

This FRB was detected with a DM of $483.51 \pm 0.01 \text{ pc cm}^{-3}$ in CB 33 (observation with Proposal ID: SCI-20220822-MC-01), which provides us with an initial position estimate. Its detection triggered the storage of the TB data, which allowed us to localize the FRB with difference images. At the ToA of the FRB, one of the *Galileo* satellites, part of the global navigation satellite system (GNSS)¹⁵ was passing approximately 1.6° away from the phase centre of the observation. Despite this relatively large distance, it appeared as an extremely bright source, producing strong artefacts that were difficult to clean in images with such short integration times. Given the

¹⁵*Galileo*: https://defence-industry-space.ec.europa.eu/eu-space/galileo-satellite-navigation_en

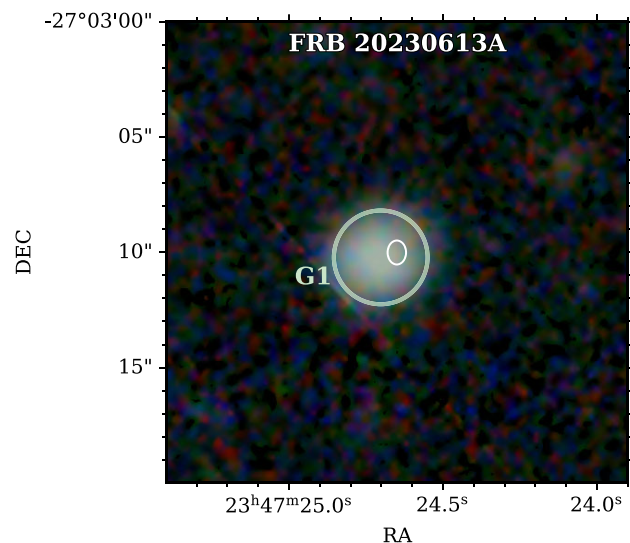


Figure 6. Localization of FRB 20230613A. The error region is shown by a white ellipse, and the background image is from DESI-DR10. The candidate host galaxy is indicated by the coloured circle labelled G1.

significant distance from the phase centre, the satellite position was located outside of the main lobe of the observation. However, because the power profile is frequency-dependent, the satellite's position fell within the sidelobes in some frequency subbands, while it coincided with the nulls between them in other subbands (M. S. Villiers & W. D. Cotton 2022). As a result, we had to flag out 36 of the 64 frequency subbands from the TB data to localize the FRB.

The FRB was fortunately bright enough to be localized in this reduced subband number. After locating it in the difference images (see Fig. 1), we performed an astrometric correction using VLASS, and determined its final localization region to be $23^{\text{h}}47^{\text{m}}24.65^{\text{s}} \pm 0.40''$ $-27^\circ03'10.01'' \pm 0.52''$. On DESI, we identified a galaxy centred $0.73''$ away from the FRB location, with a magnitude $r = 20.132 \pm 0.004$, resulting in an association probability 99.9 per cent after assuming an unseen prior of 7 per cent. Hence we determine this galaxy to be the putative host, which can be seen in Fig. 6.

To determine the redshift of the FRB, we obtained a spectrum with Keck/DEIMOS on 2023 August 11 (Program ID: U051; PI: X. Prochaska). We used the 600ZD grating at 700 nm central wavelength for a total exposure time of 2835 s. In the spectrum, we identified $\text{H}\alpha$, $[\text{N II}]$, $\text{H}\beta$, $[\text{O II}]$ and $[\text{O III}]$ lines, which we used to determine the redshift to be $z = 0.3923$ (see Fig. D1), consistent with the photometric redshift of the source, $z_{\text{phot}} = 0.42 \pm 0.03$, as well as the expected redshift from the Macquart relation, $z_{\text{Macquart}} = 0.48^{+0.14}_{-0.26}$. Given the redshift, the projected physical offset between the galaxy centre and the FRB location is $3.4 \pm 3.2 \text{ kpc}$.

We noticed that the FRB was located $\sim 5.5'$ away from the galaxy cluster Abell 4038, located in the foreground of the FRB host galaxy with a redshift of $z \sim 0.03$ (G. O. Abell, H. G. Corwin & R. P. Olowin 1989; P. A. A. Lopes et al. 2018). The cluster has a characteristic radius $R_{500} = 1.25 \text{ Mpc}$, while the impact parameter of the FRB is roughly twice that, 2.43 Mpc . With its mass $M_{500} = 5.78 \times 10^{10} M_\odot$, and following the method described in Section 3.3.2, we find that Abell 4038 could contribute $100\text{--}200 \text{ pc cm}^{-3}$ to the DM of the FRB.

The host galaxy of FRB 20230613A has broad photometric coverage from optical to infrared, including the *griz* DECam filters and the WISE W1 and W2 filters, although no W3 and W4 magnitudes are available. We used the measured magnitudes in each

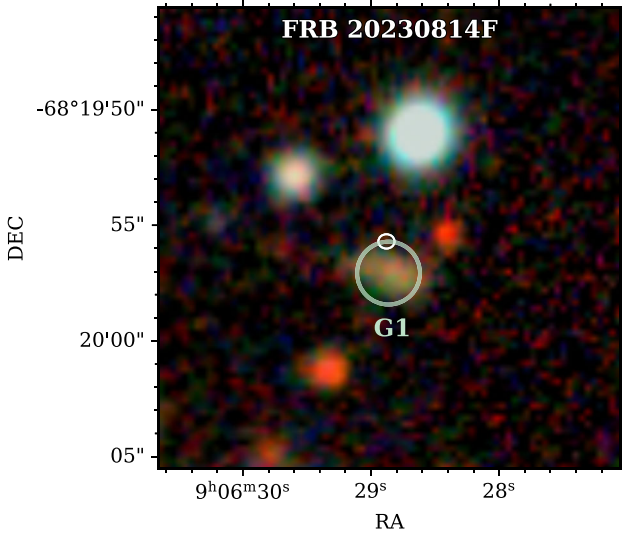


Figure 7. Localization of FRB 20230814F. The error region is shown by a white ellipse, and the background image is a composite RGB image using DECam data. The candidate host galaxies are indicated by coloured circles labelled G1-2.

filter to fit the SED with CIGALE. We find the stellar mass of the galaxy to be $\log(M_*/M_\odot) = 10.1 \pm 0.2$, the star formation rate $\text{SFR} = 4.5 M_\odot \text{ yr}^{-1}$ with 0.3 dex error, the metallicity $\log(Z/Z_\odot) < -0.13$, and the age of the galaxy 2.0 ± 0.8 Gyr.

4.1.6 FRB 20230814F

FRB 20230814F is the first ever one-off FRB discovered at *S* band. It has a DM of $471.44 \pm 0.03 \text{ pc cm}^{-3}$, and further details will be published in an accompanying paper (Pastor-Marazuela et al. in preparation). It was found during a gain calibration observation of the calibrator J0906–6829¹⁶ (Proposal ID: SCI-20200703-MK-03). The burst was detected in many CBs as well as the IB, and its detection triggered the storage of the TB data, which we used for the localization. Given the high S/N of the burst as well as the calibrator, we used deeper cleaning parameters with WSClean to remove artefacts and accurately determine the position of the burst and other sources in the field. The off/on images are shown in Fig. 1. We used other sources in the FoV to perform the astrometric correction with the RFC and RACS-mid catalogues, and determined its final error region to be $09^{\text{h}}06^{\text{m}}28.88^{\text{s}} \pm 0.35'' -68^\circ 19' 55.70'' \pm 0.31''$. Given this corresponds to a Galactic latitude $\sim -13^\circ 95$, these coordinates are not covered by DESI, whose sky coverage is bounded by $|b| > 18^\circ$. However, optical imaging already existed in the NOIRLab archive from DECam for the *g*, *r*, and *i* filters¹⁷ (M. J. Fitzpatrick et al. 2014; R. Nikutta et al. 2020). We thus ran the source finding method described in Section 3.3.1 and built an RGB composite image from the *irg* data, as shown in Fig. C1. We identify an extended source at the FRB position with a magnitude $r = 22.11 \pm 0.09$, and a separation of $1.37''$. By assuming an unseen prior of 1 per cent, we find the *PATH* association probability to be 92.0 per cent, and we thus identify it as the putative host. The FRB localization and optical background are shown in Fig. 7. A spectrum of the host galaxy,

¹⁶MeerKAT gain calibrators: <https://skaafrica.atlassian.net/wiki/spaces/ESDKB/pages/1452146701/L-band+gain+calibrators>

¹⁷Astro Data Lab image cutout access: <https://datalab.noirlab.edu/sia.php>

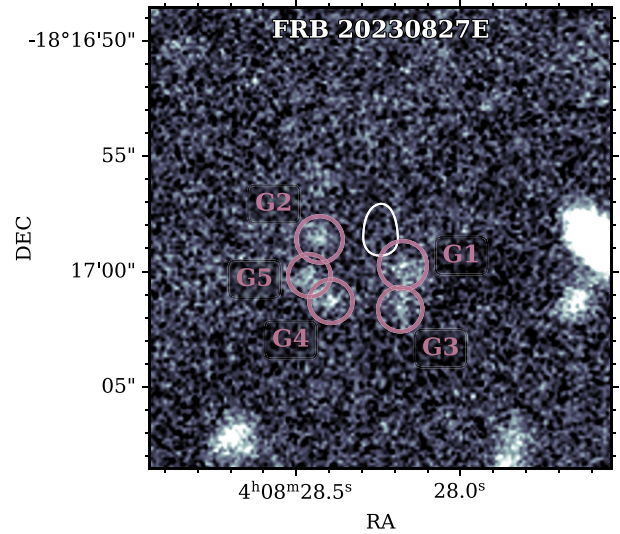


Figure 8. Localization of FRB 20230827E. The error region is shown by a white ellipse, and the background image is from a Keck/DEIMOS *R*-band observation. The candidate host galaxies are indicated by coloured circles labelled G1-5.

required to determine its redshift, has not been obtained yet. Because the existing photometric coverage is limited, we do not attempt to perform an SED fit.

4.1.7 FRB 20230827E

This bright FRB was detected with a DM of $1433.7 \pm 0.1 \text{ pc cm}^{-3}$, in a gain calibration observation of J0409-1757 (Proposal ID: EXT-20210823-RT-02). It was detected in the IB and many CBs, but unfortunately the clustering of radio transient candidates in the MeerTRAP pipeline took longer than the allowed 45 s after the burst ToA (see Section 2.2), and hence the TB data were not triggered. The large number of beams where it was detected, however, was sufficient to achieve a sub-arcsecond localization with SeeKAT. The coordinates of the FRB we determined are $04^{\text{h}}08^{\text{m}}28.24^{\text{s}} \pm 0.75'' -18^\circ 16' 58.55'' \pm 0.75''$.

The DESI DR10 archival images did not reveal any potential host galaxy coincident with the FRB location at the magnitude limit of 25.02. This is unsurprising given the high DM of the burst. Because of the high redshift expected for this source from the Macquart relation, the emission from the host galaxy is expected to be redshifted to the near-infrared (NIR) band. We thus obtained two 45 min Keck/DEIMOS observations, one in the *R* band on 2023 September 6, and the other in the *Z* band on 2023 October 7 (Program ID: O438; PI: A. Gordon). These observations revealed three to five faint sources coincident with the FRB location in the *R* band, as indicated in Fig. 8 by G1 to G5, but they were not detected in the *Z* band. The five main host galaxy candidates, located from 1.81 to 3.78 arcsec away from the best FRB position, have magnitudes ranging from 25.6 to 26.4. Given these magnitudes, obtaining sufficient signal to determine the redshift of the potential host galaxy is challenging with ground-based telescopes.

In the direction of this FRB, the MW ISM contribution to the DM ranges from 134 to 210 pc cm^{-3} according to the NE2001 and YMW16 models, respectively. The MW halo is expected to contribute ~ 50 additional DM units, resulting in an excess DM $\sim 1212 \text{ pc cm}^{-3}$. This corresponds to an expected redshift $z_{\text{Macquart}} = 1.45_{-0.68}^{+0.32}$, which could make this FRB one of the most distant

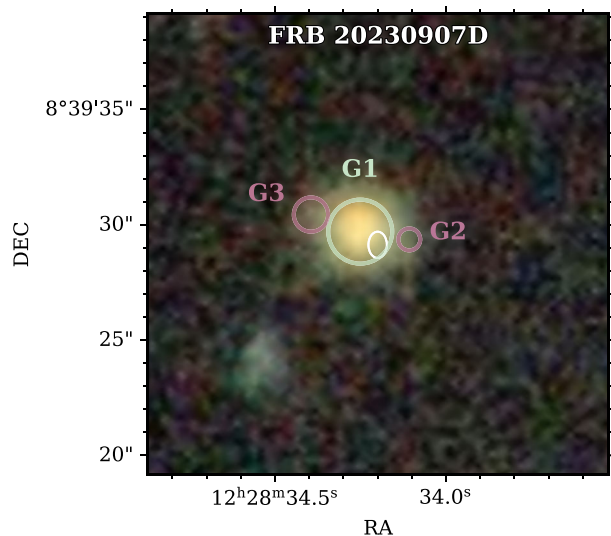


Figure 9. Localization of FRB 20230907D. The error region is shown by a white ellipse, and the background image is from DESI-DR10. The candidate host galaxy is indicated by a coloured circle labelled G1.

localized yet (S. D. Ryder et al. 2023; L. Connor et al. 2024; M. Caleb et al. 2025). We have not identified any galaxy cluster overlapping with this FRB’s sightline within twice its critical radius in the MCXC-II catalogue or the DESI galaxy cluster catalogue that could significantly contribute to the observed DM.

After performing a PATH analysis assuming an unseen prior of 10 per cent, we could not identify any of the candidate host galaxies as the host, with their posterior probabilities of association ranging from 0.8 per cent for G1, 0.7 per cent for G2, and <0.01 per cent for G3 to G5. The resulting unseen posterior is 98.4 per cent.

The configuration of the galaxy candidates is reminiscent of FRB 20220610A (S. D. Ryder et al. 2023), which was localized to a compact galaxy group at $z \sim 1$ (A. C. Gordon et al. 2024). If the galaxies in this field are physically associated, the FRB may also originate in a dense or interacting group that could enhance both recent star formation and local dispersion. Spectroscopic or integral field unit (IFU) observations will be crucial to determine the redshifts of the candidates and to confirm whether they form a physically bound group. These deeper observations could allow us to identify any faint galaxy co-located with the FRB position.

4.1.8 FRB 20230907D

FRB 20230907D was discovered during an *L*-band ThunderKAT observation (Proposal ID: SCI-20180516-PW-04). The burst, with a DM of $1030.79 \pm 0.04 \text{ pc cm}^{-3}$ in the IB, triggered the storage of the TB data. After locating the FRB in the images generated from the TB data (see off/on images in Fig. 1), we performed an astrometric correction with VLASS. We determined the final coordinates of the burst to be $12^{\text{h}}28^{\text{m}}34.20^{\text{s}} \pm 0.40'' + 08^{\circ}39'29.13'' \pm 0.57''$.

At the FRB location, our source finding pipeline identified a single galaxy with a magnitude $r = 19.83 \pm 0.09$, centred 0.92 arcsec away from the FRB position, as shown in Fig. 9. However, DESI-DR10 identifies three separate sources, a brighter one G1 with an *r*-band magnitude of 19.50 ± 0.01 , and two dimmer ones, G2 and G3, with magnitudes of 23.1 ± 0.1 and 23.0 ± 0.1 , located, respectively, 0.92 , 1.39 , and 3.18 arcsec away from the FRB localization centre.

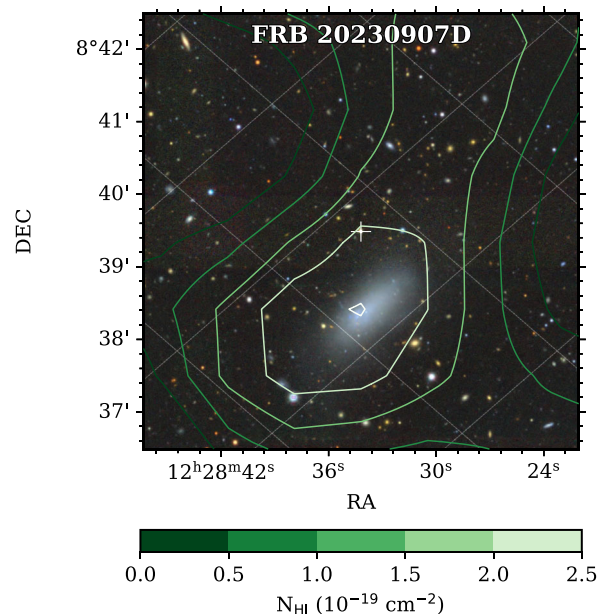


Figure 10. Localization of FRB 20230907D near the low surface brightness galaxy UGC 7596. The cross indicates the FRB and host galaxy position, and the contours the H I column density from A. Sorgho et al. (2017).

The propagation path of the FRB traversed 2.7 kpc away from the centre of the low surface brightness galaxy UGC 7596¹⁸ (P. Nilson 1973), with a redshift $z = 0.001886$ or a distance of 4.6 Mpc (I. D. Karachentsev, D. I. Makarov & E. I. Kaisina 2013), belonging to the Southern filament of the Virgo Cluster (A. Sorgho et al. 2017, see Fig. 10). Given the large angular extent of the Virgo Cluster in the sky, although the FRB location has an angular separation of $\sim 4^{\circ}$ with respect to the cluster centre, its contribution to the DM can be significant. To compute the expected DM contribution from the Virgo cluster to the DM of this FRB, we followed the method described in (D. Agarwal et al. 2019, equation 4), where they use the electron density distribution of the cluster derived from the *Planck* data (Planck Collaboration XL 2016), instead of the typical cluster parameters described in Section 3.3.2, and integrate at the FRB impact parameter (1.17 Mpc) up to twice the Virial radius of the cluster (2.4 Mpc). Using this method, we estimate the contribution of the Virgo cluster to the FRB DM to be $\sim 230 \text{ pc cm}^{-3}$. Additionally, UGC 7596 could contribute $\sim 50 \text{ pc cm}^{-3}$ to the FRB DM based on its stellar mass and optical radius (S. Díaz-García et al. 2016).

By subtracting these two contributions from the extragalactic DM of the FRB, we expect its redshift to be $z_{\text{Macquart}} = 0.84^{+0.18}_{-0.44}$. Based on the Macquart redshift upper limit, we apply an unseen prior of 40 per cent to the PATH analysis. If we consider the three galaxies identified in the DESI catalogue, we find the association probability to be $P(O|x) = 94.2$ per cent for G1, which we hence identify as the putative host, 2.7 per cent for G2, and 0.3 per cent for G3. If instead we only consider G1 as identified by our source finding algorithm, we find the association probability to be 98.5 per cent.

To determine the redshift of the host, we observed the galaxy with Keck/LRIS on 2024 January 5 (Program ID: O438; PI: A. Gordon). We observed with the 400/3400 blue grism and the 400/8500 red grating at 783 nm for a total exposure time of 2400 s . The spectrum revealed several bright emission lines, including $\text{H}\alpha$, $\text{H}\beta$, and the

¹⁸UGC 7596: SIMBAD

[N II], [S II], and [O III] doublets, which allowed us to determine a spectroscopic redshift $z = 0.464 \pm 0.015$. This redshift is well in agreement with the DESI photometric redshift, $z_{\text{phot}} = 0.44^{+0.03}_{-0.02}$, and it is at the lower end of what we expect from the Macquart relation after removing the Virgo cluster and UGC 7596, $z_{\text{Macquart}} = 0.84^{+0.18}_{-0.44}$. Motivated by this large apparent DM excess, we searched for additional intervening structures and identified three galaxy clusters listed in H. Zou et al. (2021) located in the foreground of the FRB. These clusters correspond to the catalogue identifiers 2074100051, 2074100062, and 2125300112. If we assume typical galaxy cluster profiles (Section 3.3.2), their combined DM contribution would exceed the observed value. Since these clusters are not detected in X-rays, our adopted cluster parameters are likely uncertain, and the actual DM contribution may be substantially lower. A detailed analysis of the DM budget for this and other FRBs in the sample will be presented in an upcoming publication.

The putative host galaxy of FRB 20230907D has a broad photometric coverage in the optical and IR (DECam *griz* and WISE W1-4), and we thus fitted the SED with CIGALE. We find the galaxy stellar mass to be $\log(M_*/M_\odot) = 10.9 \pm 0.2$, the $\text{SFR} = 14.9 M_\odot \text{ yr}^{-1}$ with a 0.3 dex uncertainty, the metallicity $\log(Z/Z_\odot) < -0.11$, and the age 2 ± 1 Gyr. The projected physical offset between the galaxy centre and the FRB location is 5.7 ± 4.2 kpc.

4.1.9 FRB 20231010A

This FRB was detected during an observation of the MeerTime project for pulsar timing (M. Bailes et al. 2020) (Proposal ID: SCI-20180516-MB-04). The FRB had a DM of $442.59 \pm 0.02 \text{ pc cm}^{-3}$, and it was found in CB 437. The burst detection triggered the storage of the TB data, which was used to localize it in the difference images, shown in Fig. 1. After the astrometric correction with RFC and RACS, we obtained the final position $00^{\text{h}}58^{\text{m}}55.67^{\text{s}} \pm 0.52''$ $-70^\circ 35' 46.93'' \pm 0.30''$. In the DESI DR10, images, we find this position to be close to an edge-on galaxy as well as a fainter galaxy, as can be seen in Fig. 11. The edge-on galaxy G1 has a magnitude of $r = 21.2 \pm 0.6$ and a half-light radius of 1.95 arcsec, and its centre has a separation of 1.9 arcsec from the FRB localization centre. In contrast, the fainter galaxy G2 has an *r*-band magnitude of 22.5 ± 0.8 , a half-light radius of 0.66 arcsec, and is located 0.47 arcsec away from the FRB best coordinates. Assuming an unseen prior of 5 per cent, PATH assigns a marginally higher probability of association to the brighter galaxy G1 (52.3 per cent) compared to the dimmer galaxy G2 (46.6 per cent). However, the photometric redshift from G1, $z_{\text{phot}, G1} = 0.61 \pm 0.18$, is more in agreement with the expected redshift from the Macquart relation, $z_{\text{Macquart}} = 0.42^{+0.12}_{-0.24}$, compared to G2 where $z_{\text{phot}, G2} = 0.84^{+0.27}_{-0.30}$. If the host galaxy was G2, there would be additional contributions to the DM from G1 located in the foreground, which could lead to a disagreement with the expected Macquart redshift.

The sky location where we localized this FRB has DECam *griz* coverage, but no WISE photometry is available. We use CIGALE to both estimate a photometric redshift within the DESI uncertainties, and to fit the SED, but with results poorly constrained. For G1 we find a $z_{\text{phot}} = 0.61 \pm 0.09$, and the physical parameters $\log(M_*/M_\odot) < 10.5$, $\text{SFR} = 13.8 M_\odot \text{ yr}^{-1}$ with a 0.5 dex uncertainty, $\log(Z/Z_\odot) < 0.2$, and age 1.4 ± 0.4 Gyr. For G2, the initial fit obtained with the default AGN fraction of < 20 per cent produced an excess of infrared emission and a poorly constrained stellar mass. To investigate this, we tested AGN fractions of 5, 10, 20, and 30 per cent, and found that the best-fit solution was obtained for an AGN fraction of 5 per cent.

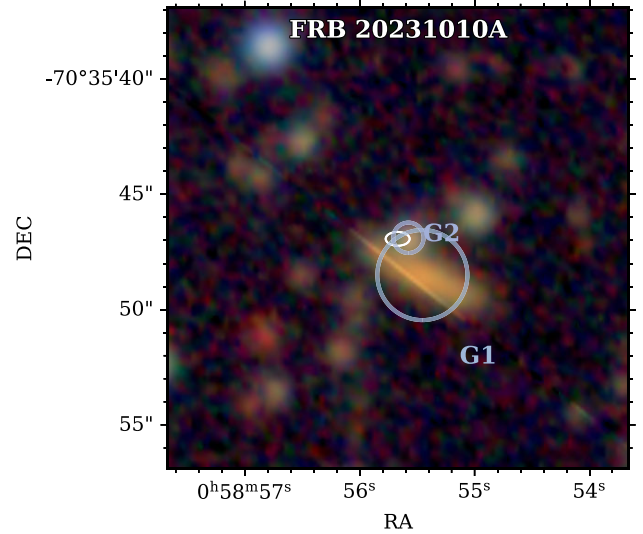


Figure 11. Localization of FRB 20231010A. The error region is shown by a white ellipse, and the background image is from DESI-DR10. The candidate host galaxies are indicated by coloured circles labelled G1-2. The diagonal line that crosses the FRB location is a product of the overlap between two adjacent exposures. The line going through G1 is an artefact from the CCD detector edges of the DESI observations.

For this model, we derive $z_{\text{phot}} = 0.54 \pm 0.1$, $\log(M_*/M_\odot) = 9.1 \pm 0.3$, $\text{SFR} = 2.5 M_\odot \text{ yr}^{-1}$ with a 0.2 dex error, $\log(Z/Z_\odot) < -0.4$, and age $= 1.1 \pm 0.1$ Gyr. The projected physical offset would be 13.2 ± 4.2 kpc for G1 and 3.1 ± 3.9 kpc for G2.

4.1.10 FRB 20231020B

FRB 20231020B was detected with a DM of $952.2 \pm 0.3 \text{ pc cm}^{-3}$ during a MHONGOOSE observation (Proposal ID: SCI-20180516-EB-01), a low-mass, late-type, gas-rich dwarf galaxy (W. J. G. d. Blok et al. 2020). The burst was detected in the IB, and it triggered the storage of the TB data. We used the TB data to generate images where we localized the burst, as shown in Fig. 1. After performing an astrometric correction with VLASS, we determined the burst coordinates to be $03^{\text{h}}49^{\text{m}}06.77^{\text{s}} \pm 0.37''$ $-37^\circ 46' 11.56'' \pm 0.40''$. In the archival DESI DR10 images, we find a single galaxy coincident with the FRB location, with a magnitude of 21.79 ± 0.01 , and a photometric redshift $z_{\text{phot}} = 0.46^{+0.07}_{-0.08}$, shown in Fig. 12. The PATH analysis determines the probability of association to be 99.8 per cent, assuming an unseen prior of 35 per cent. In order to determine the spectroscopic redshift of the putative host galaxy, we obtained a Keck/DEIMOS observation (Program ID: U051; PI: X. Prochaska) on 2023 December 14. We used the 600ZD grating at 650 nm central wavelength for a total exposure time of 730 s. From the [O III], H β and [O II] spectral lines, we determined the redshift of the host to be 0.477. The redshift is in good agreement with the photometric redshift, but it is lower than what we expect from the Macquart relation, $z_{\text{Macquart}} = 1.04^{+0.24}_{-0.52}$.

We searched for galaxy clusters in the catalogue from H. Zou et al. (2021), and identified the cluster with ID 3 959 300 018 and photometric redshift ~ 0.327 to be centred $\sim 127''$ away from the FRB location. The cluster has a characteristic radius $R_{500} = 526$ kpc, while the impact parameter of the FRB is 620 kpc. Its mass is $M_{500} = 5.75 \times 10^{13} M_\odot$. Following the method described in Section 3.3.2, we compute the expected DM contribution from the foreground cluster to be up to $\sim 400 \text{ pc cm}^{-3}$ in the cluster frame, or $\sim 300 \text{ pc cm}^{-3}$

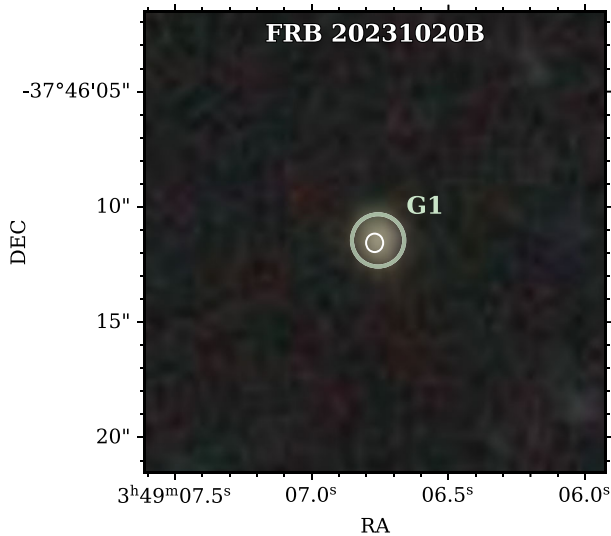


Figure 12. Localization of FRB 20231020B. The error region is shown by a white ellipse, and the background image is from DESI-DR10. The candidate host galaxy is indicated by a coloured circle labelled G1.

after correcting for redshift. This contribution from the foreground cluster could explain the observed DM.

This sky location has photometric coverage in the DECam *griz* filters and the WISE W1-3 filters, which we used to perform an SED fitting with CIGALE. We obtain a mass $\log(M_*/M_\odot) = 10.7 \pm 0.3$, an SFR = $14 M_\odot \text{ yr}^{-1}$ with a 1 dex uncertainty, a metallicity $\log(Z/Z_\odot) < 0.1$, and an age of 2.2 ± 1.4 Gyr. The projected physical offset between the FRB location and the host galaxy centre is 1.2 ± 3.3 kpc.

4.1.11 FRB 20231210F

This burst was detected in CB 306 during an observation of the MeerKAT Fornax Survey (MFS, P. Serra et al. 2023, Proposal ID: SCI-20180516-PS-01), with a DM of $720.6 \pm 0.2 \text{ pc cm}^{-3}$. The detection triggered the TB data, from which we generated images to localize the burst, as shown in Fig. 1, and then we performed an astrometric correction of the position with VLASS. The resulting FRB coordinates are $03^{\text{h}}21^{\text{m}}37.28^{\text{s}} \pm 0.25'' -35^\circ 45' 41.13'' \pm 0.25''$. In the DESI DR10 images, we find this position to coincide with an edge-on galaxy G1, and a fainter galaxy G2, apparently in the background based on the photometric redshift, in a similar configuration to those of FRB 20231010A (see Fig. 13). The edge-on galaxy has a magnitude of 21.1 ± 0.01 and it is located at 2.58 arcsec from the FRB location, while the second galaxy has a magnitude of $r = 26.4 \pm 1.0$ and a separation of 0.78 arcsec. Two further galaxies, G3 and G4, fainter than the first one but brighter than the second one, are located, respectively, 4.12 and 2.27 arcsec away from the FRB location. Assuming an unseen prior of 20 per cent, the PATH analysis finds a probability of association of 81.0 per cent with G1, 14.8 per cent with G2, 1.6 per cent with G3, and 0.03 per cent with G4. Although the probability does not reach a 90 per cent confidence for G1, we still identify it as the putative host. A spectrum of the two brightest potential host galaxies is pending to be obtained, and although we cannot yet determine the spectroscopic redshift, DESI-DR10 reports a photometric redshift of 0.50 ± 0.08 for the G1, and 1.37 ± 0.43 for the G2. The Macquart relation predicts

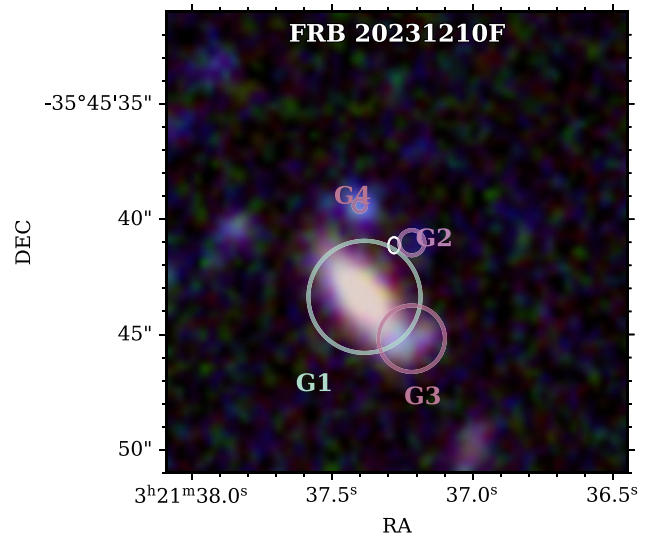


Figure 13. Localization of FRB 20231210F. The error region is shown by a white ellipse, and the background image is a composite RGB image from DESI-DR10. The candidate host galaxies are indicated by coloured circles labelled G1-4.

a redshift $z_{\text{Macquart}} = 0.78^{+0.20}_{-0.40}$, which also favours G1 as the host galaxy.

The coordinates of this FRB have photometric coverage in all DECam *griz* filters and in the WISE W1, W2, and W4 filters, which we used to fit the SED for G1. We find an optimal photometric redshift of $z_{\text{phot}} = 0.43 \pm 0.01$, a mass $\log(M_*/M_\odot) = 9.7 \pm 0.1$, which is the lowest in our host galaxy sample, SFR $< 0.2 M_\odot \text{ yr}^{-1}$, metallicity $\log(Z/Z_\odot) < -2.2$, and age 1.0 ± 0.1 Gyr. The projected physical offset would be 19.6 ± 3.7 kpc.

4.2 Arcminute localized FRBs

Some of the FRBs that were detected in 2022 and 2023 did not trigger the storage of the TB data and were too faint to be detected in several CBs or were localized by imaging the 2-s/8-s data that were taken simultaneously to the MeerTRAP observations. In this section, we present the four FRBs detected during the 2023 observations that we could not accurately localize. We show their localization areas and the sources identified as galaxies within those in Fig. 14.

4.2.1 FRB 20230306F

FRB 20230306F, with a DM of $689.5 \pm 0.9 \text{ pc cm}^{-3}$, was detected during an observation of the Virgo Cluster in CB 149 (Proposal ID: SCI-20220822-FD-02). The TB was triggered < 2 s after the maximum possible time delay, and thus the voltage data could not be saved. Given the low detection S/N ~ 10 , a localization through imaging of the 8-s data was not possible. Since it was detected only in a single CB, we could only constrain its localization to the hexagon shown in Fig. 14 centred at the coordinates $12^{\text{h}}24^{\text{m}}01^{\text{s}} \pm 60'' + 14^\circ 54' 17'' \pm 50''$. The DESI-DR10 catalogue contains 56 sources identified as galaxies within this error region and the photometric redshift upper limit of 0.94 (within errors) expected from the Macquart relation. The DESI DR10 *r*-band limiting magnitude in the direction of this FRB is 24.9, so the resulting unseen prior is ~ 23 per cent at this limiting redshift. We ran a PATH analysis on these galaxies, and the highest probability of association, given to the

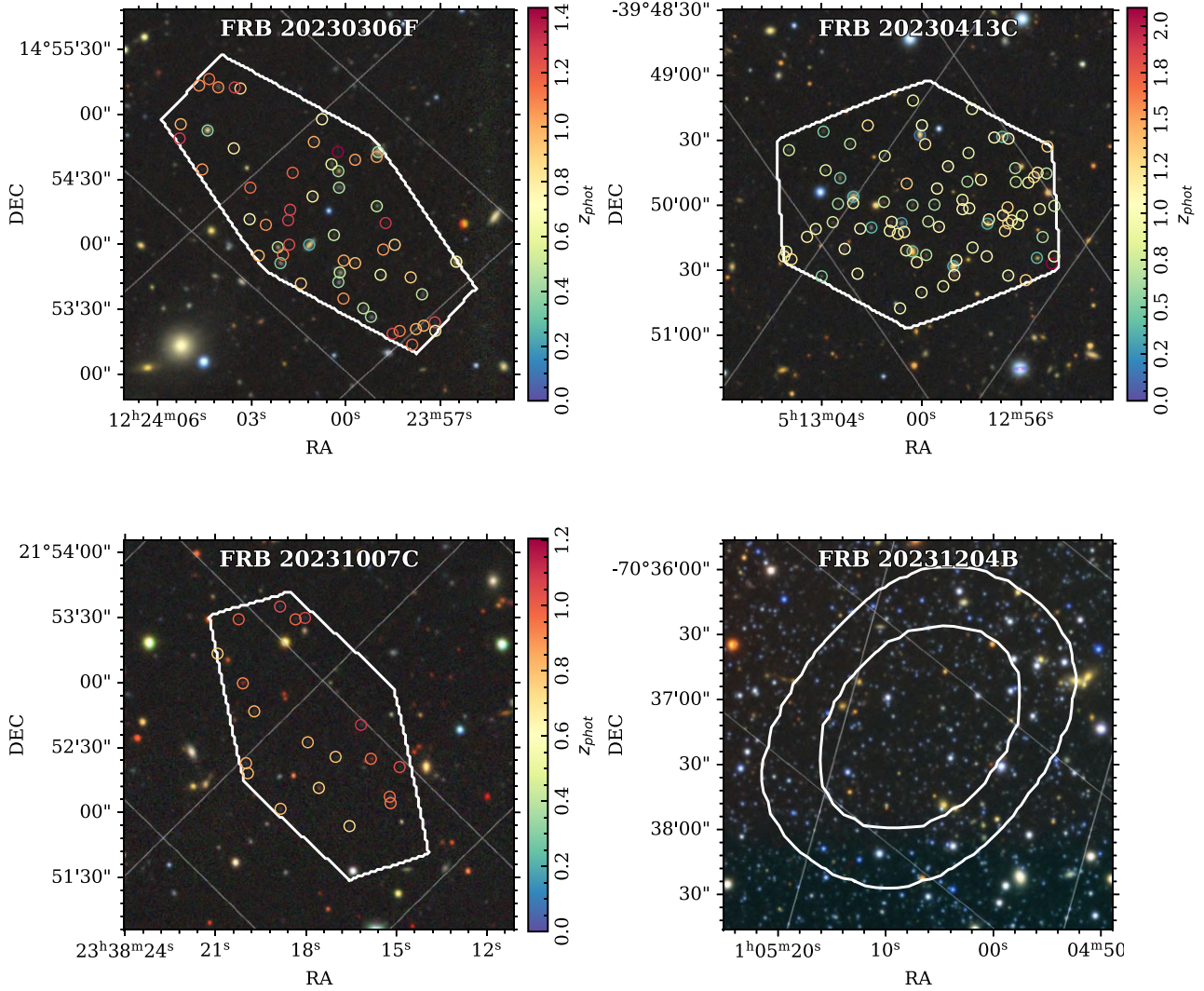


Figure 14. SeeKAT FRB localizations. The white contours show the FRB localization regions, the background shows the DESI-Legacy image, and the coloured markers show the position of the known galaxies within the error regions, going from blue to red with increasing photometric redshift.

brightest galaxy with an r -band magnitude of 19.81, is 3 per cent. The next most likely host has an association probability of 1.3 per cent, and all other galaxies are lower than 1 per cent. Meanwhile, the probability of an unseen host is around 90 per cent. This confirms that with such large localization errors it becomes unfeasible to identify the host galaxy.

4.2.2 FRB 20230413C

This FRB was detected during a UHF MeerTime observation (Proposal ID: SCI-20180516-MB-04) with a DM of $1532.2 \pm 0.5 \text{ pc cm}^{-3}$. Unfortunately, the burst was detected when the observation had just started, and since the buffers were still being cleared, the TB data were not triggered. The burst was not detected in any other beams, and although its S/N of ~ 15 was higher than that of FRB 20230306F, localization was still not possible through the imaging of the simultaneous 8-s resolution data. The SeeKAT localization thus resulted in the hexagon shown in Fig. 14, centred around the coordinates $05^{\text{h}}13^{\text{m}}00^{\text{s}} \pm 63'' - 39^{\circ}50'00'' \pm 57''$. This region contains 84 sources identified as galaxies in the DESI-DR10 catalogue with

photometric redshifts below the maximum of 2.11 expected from the Macquart relation, or 60 between a minimum of 0.92 and the maximum of 2.11 (including 1σ photometric errors). In this field, the DESI DR10 limiting r -band magnitude is 25.1, which results in an unseen prior of ~ 68 per cent. All galaxies in this sample have PATH association probabilities < 1 per cent, while the probability of an unseen host is 98 per cent, similarly to the previous FRB. It is thus not feasible to confidently identify the host galaxy of this source.

4.2.3 FRB 20231007C

This burst was detected in CB 324 during an Open Time observation (J. Healy et al. 2021, Proposal ID: SCI-20230907-TD-01). It has a DM of $2660.4 \pm 1.9 \text{ pc cm}^{-3}$, the largest of the current MeerTRAP FRB sample, but unfortunately the trigger arrived 0.5 s too late and the TB data were not stored. Since it was a single CB detection with a S/N ~ 10 , we could only localize it with SeeKAT to the hexagon shown in Fig. 14, centred around the coordinates $23^{\text{h}}38^{\text{m}}18^{\text{s}} \pm 50'' + 21^{\circ}52'34'' \pm 55''$. The DESI-DR10 catalogue contains 19 sources

identified as galaxies between photometric redshifts of 0.9 –half of the expected minimum redshift from the Macquart relation– and 3.74, the maximum expected redshift. The source number increases to 25 if we make no redshift selection. Additionally, many more galaxies probably exist within that volume but have magnitudes above the DESI depth limit. The r -band limiting magnitude in this field from the DESI DR10 is 24.3, so we would expect ~ 99 per cent of the FRB hosts to not be seen at the redshift upper limit. The PATH analysis finds the probability of an unseen host to be 99.95 per cent; it is thus not possible to confidently identify the host galaxy candidate.

4.2.4 FRB 20231204B

FRB 20231204B was found during a UHF MeerTime observation (Proposal ID: SCI-20180516-MB-04) with a DM of $1772.1 \pm 0.3 \text{ pc cm}^{-3}$. At the time of detection, this was the highest DM ever found for a MeerTRAP UHF FRB. It was detected in a CB (335) at the edge of the tiling and in the IB. Unfortunately, the trigger arrived 2 s after the maximum allowed delay, and the TB data were not stored. Although the $S/N \sim 25$ in the CB is higher than other FRBs in this section, it is still not enough to localize the FRB in the 8-s images. By combining the CB and IB detections, we localized the FRB with SeeKAT to an annular region centred around the coordinates $01^{\text{h}}05^{\text{m}}07^{\text{s}} \pm 64'' - 70^{\circ}37'16.5'' \pm 68''$. Although some galaxies are visible by eye in the DESI-DR10 image shown in Fig. 14, the DESI source finding algorithm has not been run in that region due to the proximity to the Small Magellanic Cloud. Similarly to all other FRBs presented in this section, the large localization area is expected to contain a considerable amount of galaxies, and hence determining the most likely host is not possible. At the r -band limiting magnitude of 25.5 from DESI DR10, we expect ~ 68 per cent would not be visible at the redshift upper limit.

4.3 Host galaxy sample

For eight of the FRBs with arcsecond-scale localization presented in this work, we had access to multiwavelength photometric data, which we used to perform an SED fitting with CIGALE. This allowed us to estimate key parameters of the host galaxy candidates, including stellar mass, SFR, metallicity, and age. Fig. 15 provides a visual representation of some of these properties compared to previous samples of localized FRBs, including several repeaters, and those from A. C. Gordon et al. (2023), and K. Sharma et al. (2024). The top panel shows the stellar masses and SFRs, while the bottom one shows r -band magnitudes as a function of redshift.

Our host galaxy sample has stellar masses that are consistent with the samples of repeaters, A. C. Gordon et al. (2023), and K. Sharma et al. (2024). We assessed this by performing a Kolmogorov–Smirnov (KS) two-sample test to compare the underlying distributions of these galaxy properties, considering only the most likely hosts. Since all p -values exceed 0.05, we find no significant differences in the sample distributions in terms of stellar mass given the current data set size. However, the SFRs of the MeerTRAP FRB hosts appear to be larger than the sample of repeaters, with a p -value of 0.03. Our SFRs are consistent with those from A. C. Gordon et al. (2023) and K. Sharma et al. (2024), where we find p -values > 0.05 . We note however that the galaxy SEDs were modelled with a different code using different assumptions from A. C. Gordon et al. (2023) and K. Sharma et al. (2024), and thus our results might be subject to systematic differences. While our characterization of FRB host

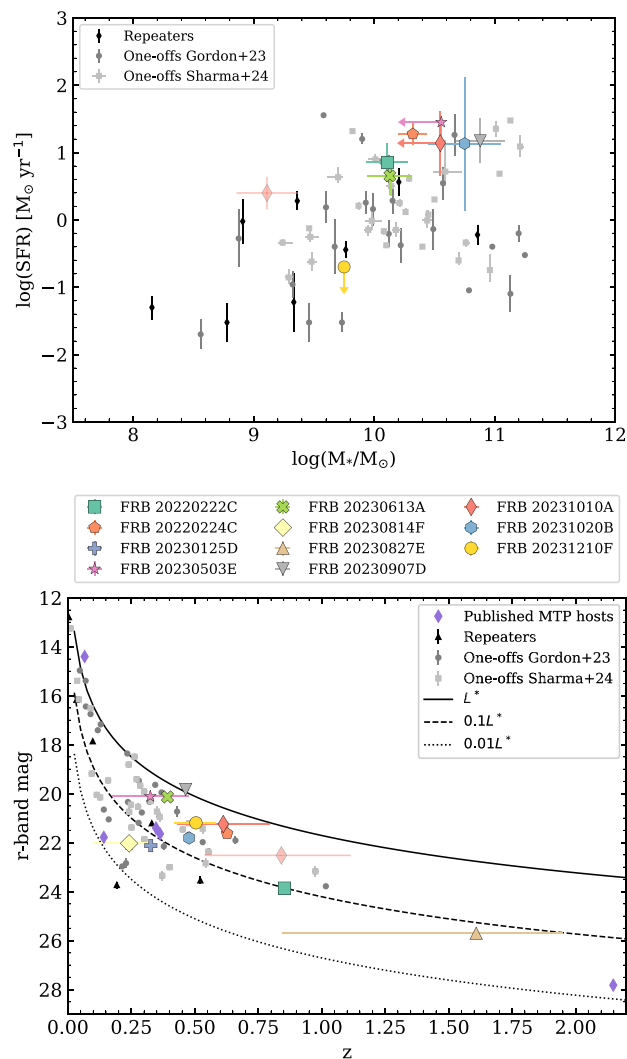


Figure 15. FRB host galaxy properties. The top panel shows the SFR as a function of stellar mass, and the bottom panel the r -band magnitude as a function of redshift. In each panel, the coloured markers show the FRBs presented in this work, the black triangles the host galaxy properties of repeaters, medium grey the galaxy properties of one-off FRBs presented in A. C. Gordon et al. (2023), and light grey the galaxies presented in K. Sharma et al. (2024). The purple diamonds are previously published MeerTRAP FRB hosts (M. Caleb et al. 2023, 2025; L. N. Driessen et al. 2024; K. Y. Hanmer et al. 2025). Note that on the top panel, the error bars on the MeerKAT FRBs are too large to be shown.

galaxies remains incomplete, in future work we will incorporate the missing hosts and host galaxy optical spectra into the SED fitting to refine our analysis.

Despite the limited sample size, MeerTRAP is producing a population of high-redshift FRBs. A KS test comparing the redshifts (spectral or photometric) of MeerTRAP FRB hosts with those of all previously localized FRBs returns p -values < 0.01 , suggesting a statistically significant difference in their redshift distributions. Given the high sensitivity of MeerKAT and the high DM values found in the pre-existing MeerTRAP FRB sample (F. Jankowski et al. 2023; I. Pastor-Marazuela et al. 2025), which naturally predispose it to detecting higher redshift sources, this result is consistent with our expectations.

5 CONCLUSIONS

In this work, we have presented the localization of 15 newly discovered FRBs within the MeerTRAP project in 2022 and 2023. While MeerTRAP has previously localized FRBs with (sub-)arcsecond accuracy (M. Caleb et al. 2023; L. N. Driessen et al. 2024; K. M. Rajwade et al. 2024; J. Tian et al. 2024, 2025; K. Y. Hanmer et al. 2025), this work presents the first large sample of MeerTRAP localized FRBs. The two FRBs from 2022 were localized in the 8-s commensal imaging observations, since the TB system was not yet operational, achieving arcsecond precision. In 2023, eight FRBs were localized by imaging the 300 ms of TB data stored upon their detection. While five additional FRBs were localized using SeeKAT, only one of these had sufficient beam detections to achieve sub-arcsecond accuracy. The remaining four were constrained to larger regions, making host galaxy identification unfeasible.

Among the 11 FRBs with (sub-)arcsecond localization accuracy, we identified host galaxies with >90 per cent confidence for eight of them. One additional FRB has a likely host galaxy candidate with >80 per cent confidence. While the host of FRB 20231010A remains ambiguous based on the PATH analysis, with two nearby galaxies showing comparable association probabilities, the photometric redshift of the brighter, edge-on galaxy G1 is consistent with the range we expect from the Macquart relation. Obtaining a spectroscopic redshift for both host galaxy candidates could hence help us confirm the host of this FRB. FRB 20230827E has five host candidates and a large unseen posterior. Its host galaxy could either be too faint to be detected in ground-based optical observations given its large excess DM of $\sim 1212 \text{ pc cm}^{-3}$. Alternatively, it could be associated to a high-redshift compact galaxy group, akin to FRB 20220610A (A. C. Gordon et al. 2024). For four FRBs in our sample, FRB 20220224C, FRB 20230613A, FRB 20230907D, and FRB 20231020B, we identify foreground galaxy clusters which significantly contribute to the observed DM, demonstrating how FRBs can be used to probe foreground structures.

Spectroscopic redshifts were obtained for six of these hosts using Keck and Gemini observations, ranging from 0.32 to 0.85. Determining redshifts for potential host galaxies remains challenging, especially for FRBs with faint or distant hosts, such as those we typically localize with MeerTRAP. Timing constraints and securing optical follow-up observations add further difficulties. As our sample grows, photometric redshift catalogues, such as DESI-DR10, are becoming essential, and they will play an even greater role when large FRB localization projects, such as the CHIME/FRB Outriggers (A. E. Lanman et al. 2024), begin detecting hundreds of FRBs yearly.

We estimated the galaxy properties of eight FRB host galaxies (or candidate hosts) with sufficient photometric coverage using the CIGALE code. The inferred stellar masses span from $10^{9.75}$ to $10^{10.94} M_{\odot}$, while the SFRs range from 0.1 to $28.5 M_{\odot} \text{ yr}^{-1}$, although some of these values represent upper limits. These stellar masses and SFR estimates are broadly consistent with those reported for other FRB samples (A. C. Gordon et al. 2023; K. Sharma et al. 2024), despite potential systematic differences arising from our different SED fitting techniques. We emphasize that the availability of spectroscopic redshifts, broader photometric coverage, and the inclusion of spectroscopic data in the SED fitting process can substantially improve constraints on the host galaxy properties of our sample. These improvements will be explored in future work.

With an increasing number of localized FRBs, MeerTRAP is uniquely assembling one of the largest uniform samples of high-redshift FRBs to date, following ASKAP (R. M. Shannon et al. 2025) and DSA-110 (C. J. Law et al. 2024). This data set is crucial to use FRBs as cosmological tools, particularly for measuring the

baryonic content of the IGM, a topic that will be explored in an upcoming publication (Caleb et al. in preparation). The localization of these distant FRBs brings us closer to answering fundamental questions about the distribution of cosmic baryons, the nature of their host galaxies, and the environments that give rise to these enigmatic bursts.

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DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Abell G. O., Corwin H. G., Jr, Olowin R. P., 1989, *ApJS*, 70, 1
 Adámek K., Armour W., 2020, *ApJS*, 247, 56
 Agarwal D. et al., 2019, *MNRAS*, 490, 1
 Aggarwal K., Budavári T., Deller A. T., Eftekhari T., James C. W., Prochaska J. X., Tendulkar S. P., 2021, *ApJ*, 911, 95
 Armour W. et al., 2020, *AstroAccelerate*, Zenodo, <https://zenodo.org/records/4282748>
 Arnaud M., 2009, *A&A*, 500, 103
 Bailes M. et al., 2020, *PASA*, 37, e028
 Barr E. D., 2017, in Weltevrede P., Perera B. B. P., Preston L. L., Sanidas S., eds, Proc. IAU Symp. 13, An S-band Receiver and Backend System for MeerKAT. Kluwer, Dordrecht, p. 175
 Bernales-Cortes L., Tejos N., Prochaska J. X., Khrykin I. S., Marnoch L., Ryder S. D., Shannon R. M., 2025, *A&A*, 696, A81
 Beroiz M., Cabral J. B., Sanchez B., 2020, *Astron. Comput.*, 32, 100384
 Bertin E., Mellier Y., Radovich M., Missonnier G., Didelon P., Morin B., 2002, in Bohlender D. A., Durand D., Handley T. H., eds, ASP Conf. Ser. Vol. 281, Astronomical Data Analysis Software and Systems XI. Astron. Soc. Pac., San Francisco, p. 228
 Bezuidenhout M. C. et al., 2023, *RASTI*, 2, 114
 Bhandari S., Flynn C., 2021, *Universe*, 7, 85
 Bhandari S. et al., 2022, *AJ*, 163, 69
 Bhardwaj M., Lee J., Ji K., 2024, *Nature*, 634, 1065
 Blok W. J. G. d. et al., 2020, *A&A*, 643, A147

- Bochenek C. D., Ravi V., Belov K. V., Hallinan G., Kocz J., Kulkarni S. R., McKenna D. L., 2020, *Nature*, 587, 59
 Boquien M., Burgarella D., Roehilly Y., Buat V., Ciesla L., Corre D., Inoue A. K., Salas H., 2019, *A&A*, 622, A103
 Bradley L. et al., 2024, *astropy/photutils: 2.0.2*, Zenodo, <https://zenodo.org/records/13989456>
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
 Caleb M. et al., 2023, *MNRAS*, 524, 2064
 Caleb M. et al., 2025, preprint (arXiv:2508.01648)
 Calzetti D., 2001, *PASP*, 113, 1449
 CASA Team, 2022, *PASP*, 134, 114501
 Chabrier G., 2003, *PASP*, 115, 763
 Chambers K. C. et al., 2019, preprint (arXiv:1612.05560)
 Chatterjee S. et al., 2017, *Nature*, 541, 58
 Chen W., Barr E., Karuppusamy R., Kramer M., Stappers B., 2021, *J. Astron. Instrum.*, 10, 2150013
 CHIME/FRB Collaboration, 2020, *Nature*, 587, 54
 CHIME/FRB Collaboration, 2021, *ApJS*, 257, 59
 CHIME/FRB Collaboration, 2023, *ApJ*, 947, 83
 CHIME/FRB Collaboration, 2025a, *ApJS*, 280, 6
 CHIME/FRB Collaboration, 2025b, *ApJ*, 989, L48
 Clark M. A., La Plante P. C., Greenhill L. J., 2011, preprint (arXiv:1107.4264)
 Clemens J. C., Crain J. A., Anderson R., 2004, in Moorwood A. F. M., Iye M., eds, Proc. SPIE Conf. Ser. Vol. 5492, The Goodman Spectrograph. SPIE, Bellingham, p. 331
 Connor L. et al., 2024, *Nat. Astro.*, 9, 1226
 Cordes J. M., Chatterjee S., 2019, *ARA&A*, 57, 417
 Cordes J. M., Lazio T. J. W., 2003, preprint (arXiv:astro-ph/0207156)
 Dale D. A., Helou G., Magdis G. E., Armus L., Díaz-Santos T., Shi Y., 2014, *ApJ*, 784, 83
 de Blok W. J. G. et al., 2024, *A&A*, 688, A109
 de Villiers M. S., 2023, *AJ*, 165, 78
 de Villiers M. S., Cotton W. D., 2022, *AJ*, 163, 135
 Dey A. et al., 2019, *AJ*, 157, 168
 Díaz-García S., Salo H., Laurikainen E., Herrera-Endoqui M., 2016, *A&A*, 587, A160
 Driessen L. N. et al., 2022, *MNRAS*, 512, 5037
 Driessen L. N. et al., 2024, *MNRAS*, 527, 3659
 Duchesne S. W. et al., 2023, *PASA*, 40, e034
 Eftekhari T. et al., 2025, *ApJ*, 979, L22
 Eikenberry S. et al., 2006, in McLean I. S., Iye M., eds, Proc. SPIE Conf. Ser. Vol. 6269, Ground-based and Airborne Instrumentation for Astronomy. SPIE, Bellingham, p. 626917
 Faber S. M. et al., 2003, in Iye M., Moorwood A.F.M., eds, Proc. SPIE Conf. Ser. Vol. 4841, Instrument Design and Performance for Optical/Infrared Ground-based Telescopes. SPIE, Bellingham, p. 1657
 Fender R. et al., 2016, in Proc. Sci., MeerKAT Science: On the Pathway to the SKA. SISSA, Trieste, PoS#13
 Fitzpatrick M. J. et al., 2014, in Peck A. B., Benn C. R., Seaman R. L., eds, Proc. SPIE Conf. Ser. Vol. 9149, Observatory Operations: Strategies, Processes, and Systems V. SPIE, Bellingham, p. 664
 Flaughner B. et al., 2015, *AJ*, 150, 150
 Gaia Collaboration, 2023, *A&A*, 674, A1
 Geda R., Crawford S. M., Hunt L., Bershady M., Tollerud E., Randriamampandry S., 2022, *AJ*, 163, 202
 Gimeno G. et al., 2016, in Evans C. J., Simard L., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 9908, Ground-Based and Airborne Instrumentation for Astronomy VI. SPIE, Bellingham, p. 872
 Gordon A. C. et al., 2023, *ApJ*, 954, 80
 Gordon A. C. et al., 2024, *ApJ*, 963, L34
 Hagstotz S., Reischke R., Lilow R., 2022, *MNRAS*, 511, 662
 Hale C. L. et al., 2021, *PASA*, 38, e058
 Hanmer K. Y. et al., 2025, *MNRAS*, 538, 1800
 Healy J., Willner S. P., Verheijen M. A. W., Blyth S. L., 2021, *AJ*, 162, 193
 Heintz K. E. et al., 2020, *ApJ*, 903, 152
 Hook I. M., Jørgensen I., Allington-Smith J. R., Davies R. L., Metcalfe N., Murowinski R. G., Crampton D., 2004, *PASP*, 116, 425
 Horowitz A., Margalit B., 2025, preprint (arXiv:2504.08038)

- James C. W. et al., 2022, *MNRAS*, 516, 4862
- Jankowski F. et al., 2022, in Ruiz J. E., Pierfedereci F., Teuben P., eds, ASP Conf. Ser. Vol. 532, *Astronomical Data Analysis Software and Systems XXX*. Astron. Soc. Pac., San Francisco, p. 273
- Jankowski F. et al., 2023, *MNRAS*, 524, 4275
- Johnson B. D., Leja J., Conroy C., Speagle J. S., 2021, *ApJS*, 254, 22
- Jonas J., the MeerKAT Team, 2018, in Proc. Sci., MeerKAT Science: On the Pathway to the SKA – PoS (MeerKAT2016). SISSA, Trieste, PoS#001
- Karachentsev I. D., Makarov D. I., Kaisina E. I., 2013, *AJ*, 145, 101
- Kassis M. F. et al., 2022, in Evans C. J., Bryant J. J., Motohara K., eds, Proc. SPIE Conf. Ser. Vol. 12184, *Ground-based and Airborne Instrumentation for Astronomy IX*. SPIE, Bellingham, p. 1218405
- Khrykin I. S. et al., 2024, *ApJ*, 973, 151
- Kirsten F. et al., 2022, *Nature*, 602, 585
- Kovacs T. O., Mao S. A., Basu A., Ma Y. K., Spitler L. G., Walker C. R. H., 2024, *A&A*, 690, A47
- Lacy M. et al., 2020, *PASP*, 132, 035001
- Lanman A. E. et al., 2024, *ApJ*, 168, 87
- Law C. J. et al., 2024, *ApJ*, 967, 29
- Lee K.-G., Ata M., Khrykin I. S., Huang Y., Prochaska J. X., Cooke J., Zhang J., Batten A., 2022, *ApJ*, 928, 9
- Lopes P. A. A., Trevisan M., Laganá T. F., Durret F., Ribeiro A. L. B., Rembold S. B., 2018, *MNRAS*, 478, 5473
- Lorimer D. R., Bailes M., McLaughlin M. A., Narkevic D. J., Crawford F., 2007, *Science*, 318, 777
- Loudas N., Li D., Strauss M. A., Leja J., 2025, *ApJ*, 991, 85
- Macquart J.-P. et al., 2020, *Nature*, 581, 391
- Mainzer A. et al., 2014, *ApJ*, 792, 30
- Marcote B. et al., 2020, *Nature*, 577, 190
- Marnoch L. et al., 2023, *MNRAS*, 525, 994
- McConnell D., Sadler E. M., Murphy T., Ekers R. D., 2012, *MNRAS*, 422, 1527
- Mckinven R. et al., 2025, *Nature*, 635, 43
- Michilli D. et al., 2023, *ApJ*, 950, 134
- Mo J.-F., Zhu W., Wang Y., Tang L., Feng L.-L., 2023, *MNRAS*, 518, 539
- Moffat A. F. J., 1969, *A&A*, 3, 455
- Mohan N., Rafferty D., 2015, *Astrophysics Source Code Library*, record ascl:1502.007
- Morello V., Rajwade K. M., Stappers B. W., 2021, *MNRAS*, 510, 1393
- Nikutta R., Fitzpatrick M., Scott A., Weaver B. A., 2020, *Astron. Comput.*, 33, 100411
- Nilson P., 1973, *Acta Universitatis Upsaliensis. Nova Acta Regiae Societatis Scientiarum Upsaliensis. Astronomiska Observatorium, Uppsala*
- Offringa A. R. et al., 2014, *MNRAS*, 444, 606
- Oke J. B. et al., 1995, *PASP*, 107, 375
- Onken C. A., Wolf C., Bessell M. S., Chang S.-W., Luvaul L. C., Tonry J. L., White M. C., Da Costa G. S., 2024, *PASA*, 41, e061
- Pastor-Marazuela I. et al., 2025, *A&A*, 693, A279
- Pastor-Marazuela I., 2025, *Philos. Trans. R. Soc. A*, 383, 20240121
- Petroff E. et al., 2017, preprint ([arXiv:1710.08155](https://arxiv.org/abs/1710.08155))
- Petroff E., Hessels J. W. T., Lorimer D. R., 2019, *Astron. Astrophys. Rev.*, 27, 4
- Petroff E., Hessels J. W. T., Lorimer D. R., 2022, *Astron. Astrophys. Rev.*, 30, 2
- Petrov L., Phillips C., Bertarini A., Murphy T., Sadler E. M., 2011, *MNRAS*, 414, 2528
- Piffaretti R., Arnaud M., Pratt G. W., Pointecouteau E., Melin J. B., 2011, *A&A*, 534, A109
- Planck Collaboration XL, 2016, *A&A*, 596, A101
- Prochaska J. et al., 2020, *J. Open Source Softw.*, 5, 2308
- Prochaska J. X. et al., 2025, *FRBs/FRB: Release for Shannon + 2025 (ICS) paper*, Zenodo, <https://zenodo.org/records/14804392>
- Rajwade K. et al., 2021, in Evans C. J., Bryant J. J., Motohara K., eds, Proc. SPIE Conf. Ser. Vol. 11447, *Ground-based and Airborne Instrumentation for Astronomy VIII*. SPIE, Bellingham, p. 114470J
- Rajwade K. M. et al., 2022, *MNRAS*, 514, 1961
- Rajwade K. M. et al., 2024, *MNRAS*, 532, 3881
- Ravi V. et al., 2023, *ApJ*, 949, L3
- Rest A. et al., 2014, *ApJ*, 795, 44
- Rockosi C. et al., 2010, in McLean I. S., Ramsay S. K., Takami H., eds, Proc. SPIE Conf. Ser. Vol. 7735, *Ground-based and Airborne Instrumentation for Astronomy III*. SPIE, Bellingham, p. 355
- Ryder S. D. et al., 2023, *Science*, 382, 294
- Sadibekova T., Arnaud M., Pratt G. W., Tarrío P., Melin J.-B., 2024, *A&A*, 688, A187
- Sanidas S., Caleb M., Driessen L., Morello V., Rajwade K., Stappers B. W., 2017, in Weltevrede P., Perera B. B. P., Preston L. L., Sanidas S., eds, Proc. IAU Symp. 13, *MeerTRAP: A Pulsar and Fast Transients Survey with MeerKAT*. Kluwer, Dordrecht, p. 406
- Scott D. R. et al., 2023, *Astron. Comput.*, 44, 100724
- Serra P. et al., 2023, *A&A*, 673, A146
- Shah V. et al., 2025, *ApJ*, 979, L21
- Shannon R. M. et al., 2025, *PASA*, 42, e036
- Sharma K. et al., 2024, *Nature*, 635, 61
- Simha S. et al., 2020, *ApJ*, 901, 134
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Sorgho A., Hess K., Carignan C., Oosterloo T. A., 2017, *MNRAS*, 464, 530
- Spitler L. G. et al., 2014, *ApJ*, 790, 101
- Spitler L. G. et al., 2016, *Nature*, 531, 202
- Tendulkar S. P. et al., 2017, *ApJ*, 834, L7
- Thornton D. et al., 2013, *Science*, 341, 53
- Tian J. et al., 2024, *MNRAS*, 533, 3174
- Tian J. et al., 2025, *MNRAS*, 540, 1685
- Walt S. v. d., Schönberger J. L., Nunez-Iglesias J., Boulogne F., Warner J. D., Yager N., Gouillard E., Yu T., 2014, *PeerJ*, 2, e453
- Yamasaki S., Totani T., 2020, *ApJ*, 888, 105
- Yao J. M., Manchester R. N., Wang N., 2017, *ApJ*, 835, 29
- Zou H. et al., 2021, *ApJS*, 253, 56

APPENDIX A: DISPERSION MEASURE SEARCH LIMITS

The DM ranged searched by the real-time MeerTRAP pipeline have varied over time for the different available observing bands. While originally the searches went up to 5118 pc cm^{-3} , when the TB system was implemented, the limits were lowered in order to reduce the trigger generation time, and improve the chances of capturing the TB data for a larger fraction of bursts. The DM limits remain much larger than the largest DM FRB we have ever detected, and hence it should not significantly bias our observed population. Table A1 lists the DM upper limits we used for the searches at different frequencies and periods of time.

Table A1. DM upper limits (pc cm^{-3}) that were used for MeerTRAP transient searches throughout the observations when the FRBs in this work were detected.

Date	UHF	L	S
2022–2023-05-05	2664	5241	5241
2023-05-05–2023-06-01	2288	4241	5241
2023-06-02–2023-10-22	2162	4241	5241
2023-10-22 onwards	2162	3841	5241

APPENDIX B: PHOTOMETRIC REDSHIFTS

To evaluate the accuracy of the photometric redshifts reported for the FRB host galaxies, we compared the performance of *CIGALE* when the redshift was left free to vary versus when it was constrained by the DESI-DR10 photometric redshift uncertainties. For each host with both spectroscopic and DESI photometric redshift measurements available, we performed two tests with *CIGALE*: first allowing the

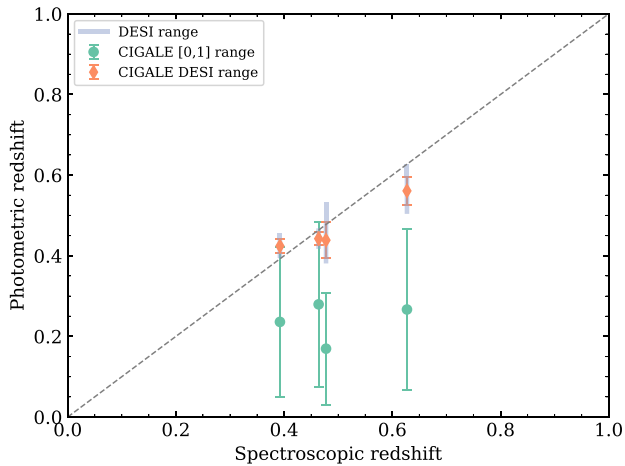


Figure B1. Comparison between spectroscopic and photometric redshifts for the FRB host galaxy sample. For each host with available spectroscopic and DESI photometric redshift (light grey bars), we show the photometric redshift derived from CIGALE when (i) the fitting redshift range is allowed to vary freely between 0 and 1 (green circles), and (ii) constrained to the DESI photometric redshift uncertainty range (orange diamonds). The dashed grey line indicates the one-to-one relation expected for perfect agreement.

redshift to vary freely between 0 and 1, and second restricting it to the range defined by the DESI photometric redshift 1σ uncertainties.

Fig. B1 shows the resulting comparison between spectroscopic and photometric redshifts. The CIGALE results with an unconstrained redshift (green circles) typically cluster around 0.2, and show larger deviations from the spectroscopic values than those constrained to the DESI range (orange diamonds), which are generally closer to the ideal relation (dashed grey line). This confirms that DESI photometric redshifts provide a more accurate prior than the unconstrained CIGALE results, and we subsequently adopt the DESI uncertainties when fitting their SEDs with CIGALE.

We acknowledge that we have performed this test on a small galaxy sample. While adopting this method to analyse a large galaxy sample would introduce biases, since the models assumed by DESI and CIGALE are different, this appears to be a good option for our small galaxy sample with limited photometric and spectroscopic information.

APPENDIX C: OPTICAL IMAGES

Fig. C1 shows a collection of the optical images matching the location of the FRBs presented in this paper with (sub-)arcsecond localization.

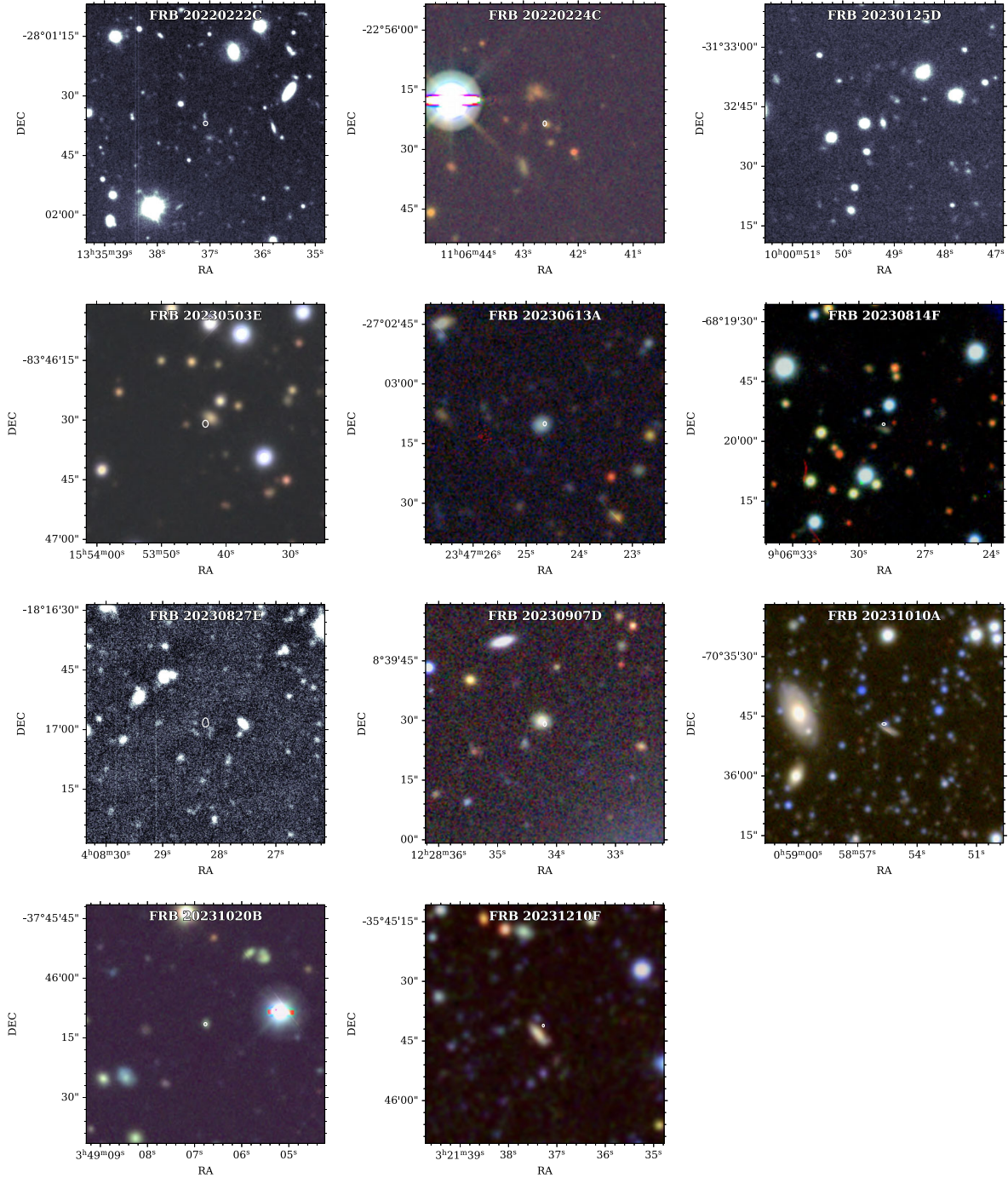


Figure C1. FRB localizations. The white ellipses at the centre of each image show the FRB localization error regions and the background shows the optical image from either the DESI-Legacy survey or the deep r -band images obtained through the observing programmes mentioned in the text. The FRB identifier is indicated at the top of each image. Each image has an FoV of $60 \text{ arcsec} \times 60 \text{ arcsec}$.

APPENDIX D: HOST GALAXY SPECTRA

In this section, we show the optical and NIR properties of the putative host galaxies of the FRBs presented in Section 4.1 for which we were able to obtain a spectrum and/or magnitudes in several bands. The

optical spectra and emission lines we identified are shown in Fig. D1. The SED fitting to the magnitudes from DECam and WISE (when available) is shown in Fig. D2. The configuration of each imaging and spectroscopic observations that were taken are detailed in Table D2.

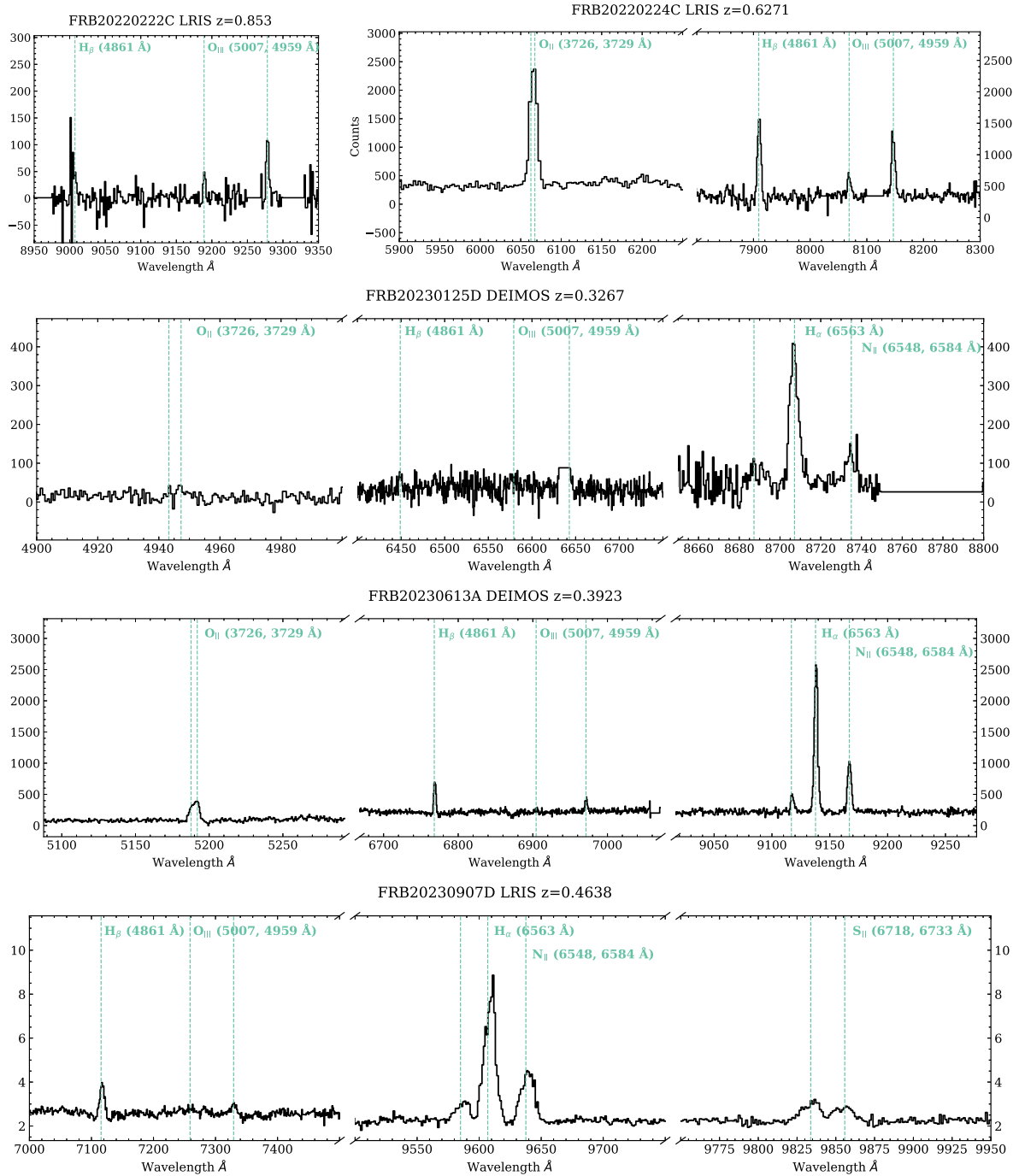


Figure D1. Spectra of FRB host galaxies. Each panel shows the counts as a function of wavelength in Å. The black line shows the spectrum, while the green vertical dashed lines show the spectral line position at the given redshift. The text next to each line indicates which spectral line it is and its wavelength at rest. In the case of doublets, two wavelengths are indicated. For each FRB host, one or more panels are shown zooming into the spectral lines that were used to determine the redshift. The title of each panel gives the FRB id, the instrument that was used to obtain the spectrum, and the redshift that was measured.

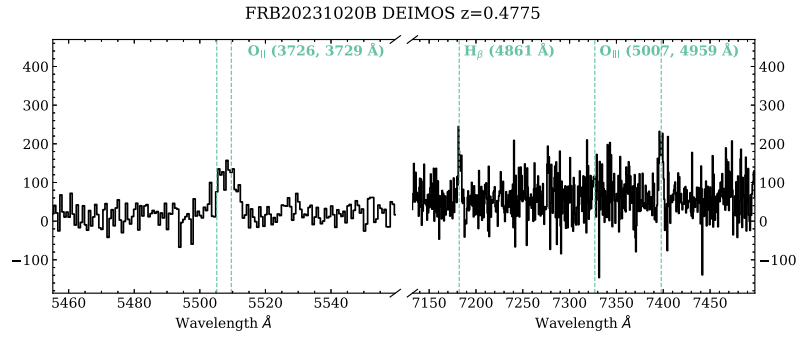


Figure D1. – continued

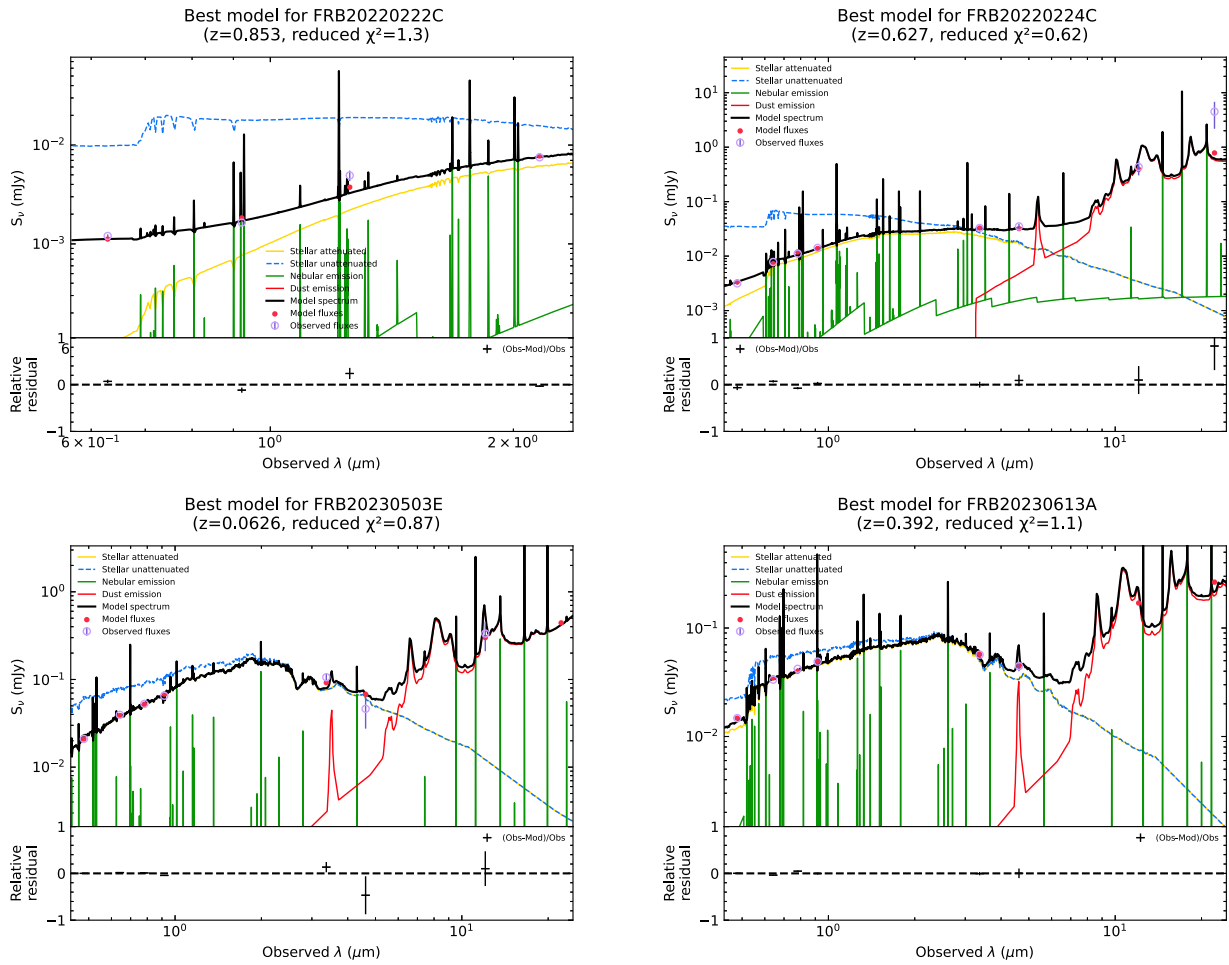


Figure D2. Spectral energy distribution fitting results with CIGALE.

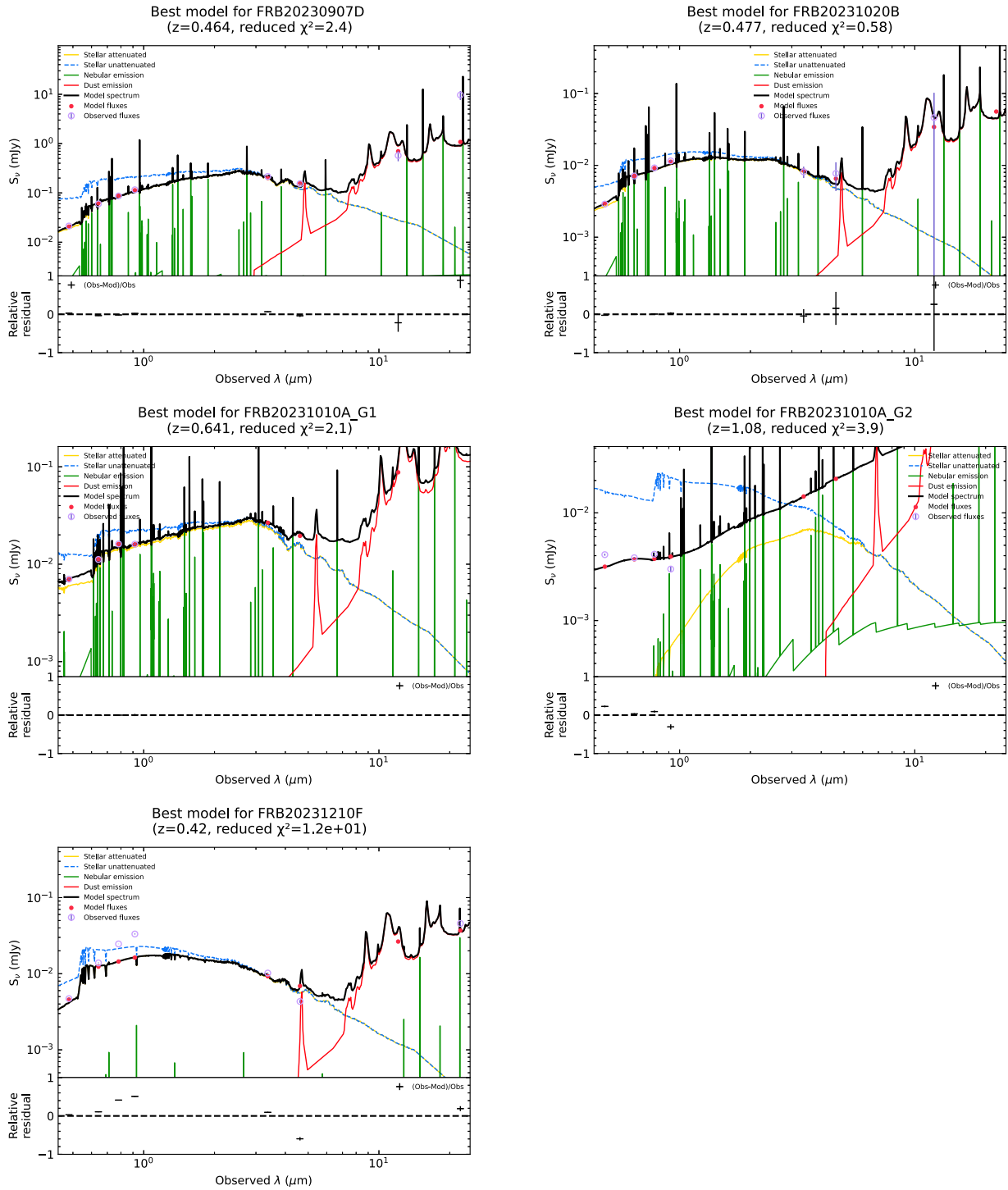


Figure D2. – continued

Table D1. FRB and host galaxy candidates. We have listed galaxy candidates within 5 arcsec of the FRB localization centre with a probability of association >1 per cent. The galaxy labels match those in the main text and figures.

FRB ID	DM	RA	Dec.	GID	Galaxy coord.	z_{spec}	z_{phot}	mag _r	R_{hl} (")	sep. (")	$P(O x)$
20220222C	1070.17	13:35:37.08 (0.54)	-28:01:36.93 (0.55)	G1	13:35:37.10 -28:01:35.53	0.853	-	23.86±0.04	1.06	1.43	0.9442
20220224C	1145.0	11:06:42.61 (0.42)	-22:56:23.48 (0.64)	G1	11:06:42.58 -22:56:23.70	0.6271	0.57 ^{+0.06} _{-0.06}	21.63±0.05	0.68	0.52	0.9964
20230125D	640.37	10:00:49.21 (0.24)	-31:32:40.77 (0.26)	G1	10:00:49.23 -31:32:40.78	0.3265	-	22.12±0.02	0.96	0.29	0.9781
20230503E	484.03	15:53:43.19 (0.69)	-83:46:30.93 (0.87)	G1	15:53:42.38 -83:46:29.40	-	0.32 ^{+0.15} _{-0.15}	20.11±0.01	1.37	2.02	0.9733
20230613A	483.82	23:47:24.65 (0.40)	-27:03:10.01 (0.52)	G1	23:47:24.70 -27:03:10.22	0.3923	0.42 ^{+0.03} _{-0.03}	20.132±0.004	1.39	0.73	0.9993
20230814F	471.73	09:06:28.88 (0.35)	-68:19:55.70 (0.31)	G1	09:06:28.86 -68:19:57.07	-	-	22.03±0.05	1.36	1.37	0.9199
20230827E	1434.0	04:08:28.242 (0.75)	-18:16:58.55 (1.50)	G1	04:08:28.17 -18:16:59.71	-	-	26.40±0.02	1.04	1.51	0.0082
				G2	04:08:28.43 -18:16:58.59	-	-	25.81±0.02	1.01	2.64	0.0074
				G3	04:08:28.18 -18:17:01.63	-	-	26.40±0.02	0.96	3.19	3 × 10 ⁻⁶
				G4	04:08:28.39 -18:17:01.28	-	-	25.56±0.01	0.94	3.48	7 × 10 ⁻⁶
				G5	04:08:28.46 -18:17:00.16	-	-	26.13±0.02	0.92	3.48	0.0001
20230907D	1031.11	12:28:34.20 (0.40)	+08:39:29.13 (0.57)	G1	12:28:34.25 +08:39:29.71	0.4638	0.44 ^{+0.03} _{-0.02}	19.83±0.09	1.39	0.92	0.9585
20231010A	442.81	00:58:55.67 (0.52)	-70:35:46.93 (0.30)	G1	00:58:55.45 -70:35:48.50	-	0.61 ^{+0.18} _{-0.18}	21.24±0.01	1.9	2.38	0.5232
				G2	00:58:55.58 -70:35:46.89	-	0.84 ^{+0.27} _{-0.30}	22.52±0.01	0.47	1.61	0.4662
20231020B	953.84	03:49:06.77 (0.37)	-37:46:11.56 (0.40)	G1	03:49:06.76 -37:46:11.45	0.4775	0.46 ^{+0.07} _{-0.08}	21.81±0.02	0.88	0.18	0.9984
20231210F	721.46	03:21:37.28 (0.25)	-35:45:41.13 (0.25)	G1	03:21:37.39 -35:45:43.37	-	0.50 ^{+0.08} _{-0.08}	21.19±0.02	1.97	2.58	0.8103
				G2	03:21:37.22 -35:45:41.02	-	1.37 ^{+0.43} _{-0.43}	26.4±1.0	0.41	0.78	0.1483
				G3	03:21:37.22 -35:45:45.18	-	0.99 ^{+0.70} _{-0.54}	24.1±0.1	1.08	4.12	0.0160
				G4	03:21:37.40 -35:45:39.41	-	0.86 ^{+0.27} _{-0.33}	23.9±0.1	0.56	2.27	0.0003

Table D2. Optical and near-infrared imaging and spectroscopic observation summary.

FRB ID	Facility	Instrument	Obs. date	Type	Configuration	Exp. time	Seeing*	Program ID	PI
20220222C	Gemini-S	GMOS	2022 May 25	Imaging	r band	28 × 100 s	0.6"	GS-2022A-Q-143	N. Tejos
	Gemini-S	GMOS	2023 June 16	Imaging	z band	25 × 100 s	0.6"	GS-2023A-Q-215	A. Gordon
	Gemini-S	Flamingos2	2024 June 16	Imaging	Ks band	60 × 10 s	0.5"	GS-2023A-Q-215	A. Gordon
	Gemini-S	Flamingos2	2025 June 16	Imaging	J band	30 × 60 s	0.5"	GS-2023A-Q-215	A. Gordon
	Keck	LRIS	2023 Apr 17	Spec	R600/10000, 8082 Å	6450 s	1.5"	U173	X. Prochaska
	Keck	LRIS	2023 Apr 17	Spec	B300/5000	6600 s	1.5"	U173	X. Prochaska
20220224C	Gemini-S	GMOS	2023 Jan 29	Imaging	r band	15 × 180 s	0.8"	GS-2022B-Q-123	A. Gordon
	Gemini-S	GMOS	2023 Jan 30	Imaging	i band	15 × 180 s	0.7"	GS-2022B-Q-123	A. Gordon
	Keck	LRIS	2023 Jan 24	Spec	R600/10000, 8927 Å	2740 s	1.1"	U055	J. Hennawi
	Keck	LRIS	2023 Jan 24	Spec	B300/5000	2780 s	1.1"	U055	J. Hennawi
20230125D	Gemini-S	GMOS	2023 May 16	Imaging	r band	15 × 180 s	0.7"	GS-2023A-Q-215	A. Gordon
	Keck	DEIMOS	2023 Dec 14	Spec	600ZD, 6500 Å	2700 s	2.1"	U051	X. Prochaska
20230613A	Keck	DEIMOS	2023 Aug 11	Spec	600ZD, 7000 Å	2835 s	1.3"	U051	X. Prochaska
20230827E	Keck	DEIMOS	2023 Sept 06	Imaging	R band	9 × 300 s	0.9"	O438	A. Gordon
	Keck	DEIMOS	2023 Oct 07	Imaging	Z band	9 × 300 s	0.9"	O439	A. Gordon
20230907D	Keck	LRIS	2024 Jan 05	Spec	R400/8500, 7830 Å	2400 s	1.0"	O438	A. Gordon
	Keck	LRIS	2025 Jan 05	Spec	B400/3400	2400 s	1.0"	O438	A. Gordon
20231020B	Keck	DEIMOS	2023 Dec 14	Spec	600ZD, 6500 Å	730 s	1.0"	U051	X. Prochaska
20231210F	SOAR	Goodman	2024 Feb 02	Imaging	r band	12 × 300 s	0.4"	SOAR2024A-002	A. Gordon

*Notes.** Seeing: for imaging observations, the seeing is estimated from the FWHM of the non-saturated stars contained in the image, fitted to a Moffat profile (A. F. J. Moffat 1969). For spectroscopic observations, the seeing is obtained as a `TYPEIT` output (J. Prochaska et al. 2020).

Table D3. Photometric data and CIGALE SED fitting results.

Property	220222 G1	220224 G1	230503 G1	230613 G1	230907 G1	231010 G1	231010 G2	231020 G1	231210 G1
z_{spec}	0.853	0.6271	–	0.39239	0.4637	–	–	0.4775	–
z_{phot}	–	–	0.29 ± 0.26	–	–	0.61 ± 0.09	0.84 ± 0.3	–	0.43 ± 0.01
g	–	22.65 ± 0.06	20.891 ± 0.007	20.97 ± 0.01	20.58 ± 0.01	21.87 ± 0.02	22.46 ± 0.02	22.75 ± 0.01	22.28 ± 0.01
r	23.86 ± 0.04	21.63 ± 0.03	20.113 ± 0.04	20.132 ± 0.004	19.83 ± 0.09	21.24 ± 0.01	22.52 ± 0.02	21.81 ± 0.02	21.19 ± 0.02
i	–	21.29 ± 0.02	19.749 ± 0.05	19.84 ± 0.01	19.06 ± 0.01	20.93 ± 0.01	22.41 ± 0.03	21.49 ± 0.01	20.45 ± 0.01
z	23.45 ± 0.06	21.03 ± 0.04	19.481 ± 0.008	19.67 ± 0.02	18.74 ± 0.01	20.92 ± 0.03	22.74 ± 0.07	21.25 ± 0.03	20.12 ± 0.02
J	22.22 ± 0.13	–	–	–	–	–	–	–	–
W1	–	22.81 ± 0.07	18.85 ± 0.12	22.21 ± 0.04	20.77 ± 0.01	–	–	24.33 ± 0.2	21.38 ± 0.02
W2	–	23.38 ± 0.13	19.75 ± 0.44	23.12 ± 0.11	21.77 ± 0.03	–	–	25.02 ± 0.47	22.31 ± 0.05
W3	–	22.49 ± 0.33	17.6 ± 0.4	–	22.18 ± 0.26	–	–	24.91 ± 1.32	–
W4	–	21.39 ± 0.56	–	–	20.56 ± 0.22	–	–	–	–
$\log(M_*/M_\odot)$	10.11 ± 0.17	10.32 ± 0.12	<10.56	10.13 ± 0.18	10.88 ± 0.2	<10.54	9.11 ± 0.25	10.75 ± 0.3	9.75 ± 0.04
SFR*	7.15 (0.29)	18.88 (0.17)	<25.8	4.5 (0.28)	14.91 (0.33)	11.05 (0.48)	2.51 (0.24)	13.58 (0.99)	<0.2
$\log(Z/Z_\odot)$	<-0.13	<0.70	<0.21	<-0.13	<-0.11	<0.18	<-0.43	<0.09	<-2.26
Age (Gyr)	1.8 ± 0.7	1.4 ± 0.5	2.0 ± 1.0	2.0 ± 0.8	2.3 ± 1.1	1.4 ± 0.4	1.1 ± 0.1	2.2 ± 1.4	1.0 ± 0.1

Notes. * SFR: star formation rate in $M_\odot \text{ yr}^{-1}$. The uncertainties between parentheses are in dex.

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