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Article

X-ray Properties of Two Complementary Samples of Intermediate Seyfert Galaxies

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Abstract: We present the X-ray spectral analysis of two complementary sets of intermediate Seyfert galaxies (ISs). Analyzing X-ray data, we estimate the hydrogen abundance N_H and test its connection with the [O III] luminosity acquired from optical observations. The results confirm the conclusions drawn in a previous study concerning the lack of a direct correlation between the obscuration measure (N_H) and the intrinsic characteristics of the active nuclei ([O III] luminosity). Instead, we validate the existence of a correlation between the Seyfert type and the N_H parameter, employing a separation threshold of approximately 10^{22} atoms cm^{-2} . Simultaneously, our findings align with prior research, corroborating the relationship between X-ray luminosity and the [O III] luminosity.

Keywords: Seyfert galaxies; intermediate Seyfert; AGN; X-ray spectroscopy

1. Introduction

Intermediate Seyfert galaxies (ISs) [1,2] represent a category within active galactic nuclei (AGN) that, according to the Unified Model (UM) [3–5], are viewed at intermediate angles. As a result, they exhibit a partial level of obscuration originating from dust surrounding the nucleus. Their optical spectra display a composite profile for permitted lines (e.g., $H\alpha$ and $H\beta$), comprising both a narrow and a broad component. These components stem from the narrow-line region (NLR) and the broad-line region (BLR), respectively. The NLR tends to be more extended, whereas the BLR is concentrated within a dusty torus. Consequently, the spectra of Seyferts at greater inclinations, such as type 2 (Sy2), exhibit solely the narrow component, whereas those at smaller inclinations, such as type 1 (Sy1), exclusively reveal the broad component. ISs can be further classified based on either an escalating obscuration sequence or an increasing prevalence of the narrow-line component in relation to the broad one: Sy1.2, Sy1.5, Sy1.8, and Sy1.9 (for a detailed discussion, refer to [6]).

As inclination angles decrease to their minimum, going from Sy2 to Sy1, the relevance of the γ -ray emission component becomes more pronounced. Among the various types of AGN, one—blazars—exhibits prominent γ -ray emissions. In this case, the relativistic jets are aligned, or almost aligned, with the observer's line of sight, intensifying the emission due to the beaming effect. Conversely, for the other AGN categories, as for the Seyferts, the presence of the γ -ray component is expected to be absent or nearly absent. However, there have been recent reclassifications, such as J2118 – 0732, identified as the first non-local IS capable of maintaining relativistic jets detectable at γ -ray energies [7]. This atypical behavior could be explained by introducing the concept of misaligned jets. Approximately 2.8% of γ -ray-detected AGN exhibit these misaligned jets [8]. Consequently, the scarcity of γ -ray-emitting ISs might be attributable to instrumental limitations linked to the faintness of high-energy emission or to intrinsic processes altering the line profile, which are unrelated to obscuration.



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To explore these contrasting scenarios, our focus centers on two complementary sets of ISs, chosen based on the selection criteria detailed in [6]. We categorize these groups as “jetted” and “non-jetted” objects. Our aim is to find some spectral quantities that separate between the two classes, independently from the obscuration level proposed by the UM. In [6], we concentrated on the optical spectra of these sources, particularly focusing on [O III] λ 5007, H α , and H β . The oxygen line serves as an indicator of the intrinsic properties of the AGN, in particular as a tool for the disk luminosity [9–11], even if its utility is subject to certain limitations [12,13]. The H α and H β lines enable the calculation of the Balmer decrement, offering an obscuration measure. However, the reliability of this quantity has been extensively debated by several authors [14–16]. These authors highlighted the limitations of assuming a theoretical value of 2.9 for the H α /H β ratio (case B recombination). In particular, ref. [16] found a range of 2.5–6.6 for the theoretical value, concluding that a single set of extinction properties cannot be assumed for all the BLRs. As an alternative, the X-ray spectrum provides a more dependable measure of obscuration through the hydrogen column density parameter (N_H).

In [6], we did not find a clear separation between the jetted and non-jetted sources in the [O III]/H β -L_[O III] plot which means that the central engine’s structure in Seyfert galaxies may be consistent across all types and the obscuration-dependent scenario remains the most probable solution. On the other hand, in the central panel of Figure 6 of [6], we compared the two IS classification methods introduced by [17,18], but a clear separation between the Seyfert sub-types is not evident. This can imply two scenarios; in the first, the jetted sources do not always correspond to Sy1.2/1.5, but can sometimes be related to Sy1.8/1.9. In the second, the selection methods based on hard X-rays and γ -rays only bias the classification. In this case, some hard-X-ray-selected objects can be Sy1.2/1.5 and γ -ray-selected AGN Sy1.8/1.9, which in turn can be explained through the misaligned jets introduced before. To confirm or reject these results, we need a more reliable measure of the intrinsic absorption that can be obtained using X-ray spectroscopy as stated in the previous paragraph.

The structure of this paper is as follows: Section 2 details the Data Reduction and Analysis, Section 3 presents the Results, and Section 4 outlines the Summary and Conclusions. We adopt a standard Λ CDM cosmology with the Hubble constant and the cosmological parameters, respectively, of: $H_0 = 73.3 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_{\text{matter}} = 0.3$, and $\Omega_{\text{vacuum}} = 0.7$ [19].

2. Data Reduction and Analysis

2.1. Sample and Data Reduction

The dataset comprises 38 sources from the *Swift*-BAT AGN Spectroscopic Survey (BASS) and 11 sources from the fourth Fermi Gamma-ray Large Area Telescope (4FGL). In the former, sources were selected using their emission in hard X-rays, suggesting a reflection of emitted radiation from the accretion flow towards the dusty region. This method tended to show a preference for higher Seyfert types (Sy2, Sy1.9, and Sy1.8). Conversely, the latter is γ -ray-selected with sources exhibiting a bias toward lower Seyfert types (Sy1, Sy1.2, and Sy1.5) due to the beaming effect. All the optical spectra come from the Sloan Digital Sky Survey (SDSS); for a comprehensive description, refer to Paper I.

For the X-ray counterpart, 42 sources presented *Swift* data, while for the remainder we employed higher-resolution observations from *Chandra* and *XMM-Newton*. Refer to Table A1 in the Appendix A for specific details. Our objective is to measure the N_H^{int} parameter (intrinsic component of N_H), without requiring a comprehensive understanding of the source’s physics during the fitting process.

Data retrieval was performed from the High Energy Astrophysics Science Archive Research Center (HEASARC), followed by the utilization of the *Swift*-XRT pipeline version 0.13.7 and CALDB files updated on 24 July 2023 to extract the spectra. Similarly, *Chandra* data were processed using dedicated reduction tools like ciao-4.15 and a calibration database version 4.10.7 (14 September 2023). For *XMM-Newton*, the same with SAS 21.0.0 (xmmsas_20230412_1735-21.0.0) for the extraction¹, and the respective calibration datasets

updated on 5 October 2023. For *Swift* and *XMM-Newton* data, the source extraction radius was 40", while the background spectra were extracted from source-free circular regions of 1' for each observation. For *Chandra* observations, the extraction area was of 6" for the sources and of 1' for the background.

Subsequently, for each source we summed the available spectra to increase the statistics and binned the data into groups of 20 counts/bin to employ the χ^2 statistics. This procedure was applied to all cases except J0958 + 3224 (3C 232). For this particular source, we applied likelihood statistics for Poisson-distributed data using the c-stat package within Xspec.

2.2. Data Analysis

The analysis was conducted utilizing Xspec version 12.13.0c. As mentioned earlier, our primary objective was to estimate N_H^{int} in units of atoms cm^{-2} . Initially, we employed a simple model comprising galactic absorption (tbabs) combined with either a single redshifted power-law (zpo) or a broken power-law (bknpo) to represent AGN emission, denoted as tbabs*zpo or tbabs*bknpo. The galactic absorption data (N_H^{gal}) were acquired through HEASARC's specialized tool² [20]. This initial model was suitable for Seyferts where internal obscuration plays a negligible role, such as Sy1. For these objects, the value reported in the plots refers to the sole external galactic absorption component (N_H^{gal}). However, in Sy2 and certain ISs, internal obscuration from the dust surrounding the nucleus (N_H^{int}) could be significant. Observationally, these sources indicate a deficiency in the residuals at low energies (<1 keV) when fitted solely with galactic absorption and a power-law or broken power-law model. Consequently, we modified the model by incorporating an internal obscuration component (ztbabs), adopting the following structure: tbabs*ztbabs*zpo or tbabs*ztbabs*bknpo. When this representation proved to be unsatisfactory, we resorted to a more complex model, composed of tbabs*(zpo+ztbabs*zpo), adapted from [21]. In this model, the first power-law is obscured through internal processes and represents the emission from the central regions, such as the accretion disk. The second power-law emission is dimmed only by the galactic dust and reflects the scattered emission. Furthermore, whenever fluctuations surpassing $>3\sigma$ were detected around 6.4 KeV, we incorporated a zgauss component to reproduce the Fe K α line.

However, in the case of 3C 234.0, fitting complications arose, in particular at energies <1 keV, see Figure 1. The residuals showed a significant deficit that it was not possible to eliminate introducing an intrinsic absorption. Furthermore, these residuals showed a peculiar trend made by several bumps. To address these complexities, we utilized the following model: tbabs*(zpo+apec+apec+ztbabs*(zpo+zga)) similar to those introduced by [22]. This model comprises two components representing collisionally ionized diffuse gas (apec) to reproduce the main bumps (see the region around 0.5 keV in Figure 1), a model for obscured-type emissions (tbabs*(zpo+ztbabs*zpo)), and a Gaussian line to account for the iron emission line (zgauss). The second part of the model is analogous to the one described earlier and represents the absorbed and scattered emission of the AGN, while the various bumps at <1 keV present a fitting challenge. In [22], the authors propose two models to produce a satisfactory result: two components of ionized diffuse gas (mekal, similar to apec), or a series of narrow-Gaussian lines. For our purposes, a complete physical description of the sources is not needed, so we decided to use the simplest of the two and obtain a reliable measure of the N_H^{int} parameter.

For a complete description of the models used, see Tables A2 and A3 and the HEASARC manual³.

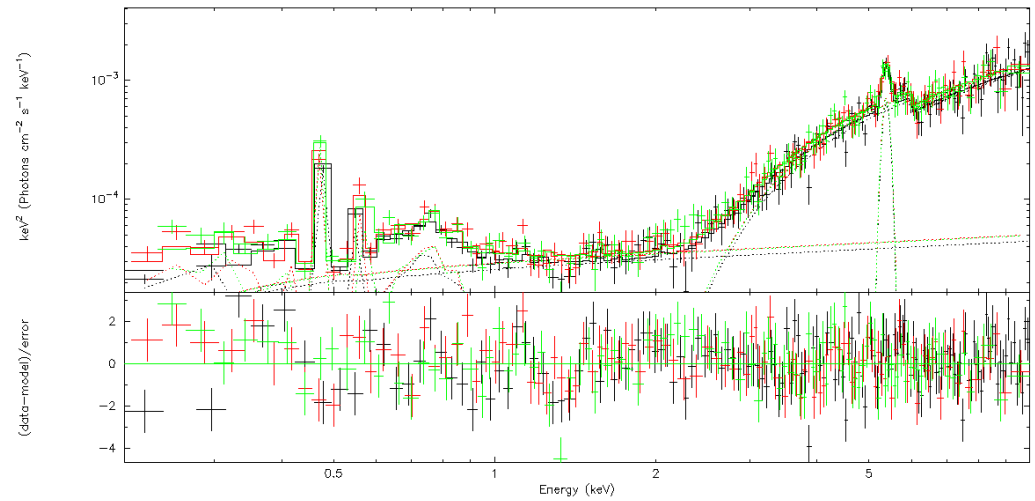


Figure 1. Upper: The result for the fitting of 3C 234.0. The datapoints from the PN camera are in black, while the points from the two MOS cameras are in green and red. **Lower:** The residuals of the three fitting procedures as the difference between the model and the datapoints.

3. Results

In this section, we present our findings in terms of various optical parameters plotted against the X-ray counterparts, notably the total column density as a sum of the galactic one and the intrinsic one $N_H^{\text{tot}} = N_H^{\text{int}} + N_H^{\text{gal}}$. This choice is related to the possibility of comparing our results with other studies and the intent to not pile up all the points with $N_H^{\text{int}} = 0$ along the x-axis. Nevertheless, we also generated identical plots excluding the sources with $N_H^{\text{int}} = 0$, see Figures A1 and A2 in Appendix B. This was done to assess whether the relationships with optical estimations of obscuration show improvement. This analysis corresponds to the one conducted in [6]; for reference, refer to Figure 6 and Section 5 of [6]. of the aforementioned work.

Figures 2–4 depict analogous plots to those in Figure 6 of [6], where we related the Seyfert type (calculated through the methods presented in [17,18]) with the Balmer decrement and the [O III] luminosity. Our objective is to identify a correlation between the parameters that demonstrates a differentiation between the jetted and non-jetted sources. Similarly to our previous study, Figures 2–4 represent the relation between N_H^{tot} and the extinction $A(V)$, the Seyfert type according to [17], and the [O III] luminosity. All these plots suggests that a distinction between jetted and non-jetted sources is not evident. The existence of separation would imply a relationship between the nature of the source and the involved quantities.

Moving to Figure 5, it showcases the comparison between optical and X-ray luminosities. Although the relationship between these two quantities is present, it is not particularly strong (see the next sections for the details). These results align with the findings from [23,24].

3.1. $A(V)$ versus N_H^{tot}

The parameter $A(V)$ quantifies the extinction in magnitudes and directly relies on the Balmer decrement, serving as another means to estimate it. These two parameters are interrelated through the equation provided in [25]:

$$A(V) = 7.215 \cdot \log \left(\frac{2.86}{H\alpha/H\beta} \right) \quad (1)$$

Here, $H\alpha$ and $H\beta$ represent the flux of the narrow component of the respective lines. The green area in Figure 2, sourced from [26,27], reflects the Galactic standard ratio and helps in comparing similar plots, such as the one found in Figure 3 of [28]. Notably, in the

latter study, the authors present a broader extinction range between 0 and 30 mag, whereas in our dataset, the objects fall within $0 \leq A(V) < 4.5$ mag. In this case, there is no need to produce the same plots by discarding objects with $N_H^{int} = 0$, as the Galactic contribution is emphasized through the green area.

In our study and in [28], the sources seem to scatter around the Galactic ratio within the initial extinction range (e.g., $A(V) < 5$ mag). A similar trend is also noted in another study (Figure 4 of [16]), where $A(V)$ ranges from 0 to 10 mag. However, the limited range of extinction in our dataset limits the possibility of completely comparing our results with previous works. From the UM and the selection criteria cited in Section 2.1, we will expect a separation between jetted and non-jetted objects, where the red points concentrate in the low-extinction region (right part of the plot), and the blue points in the high-extinction one (left side). Though the sources do not exhibit a clear separation according to this classification, this discrepancy with the UM might partially stem from the restricted range of extinction values available for the analysis.

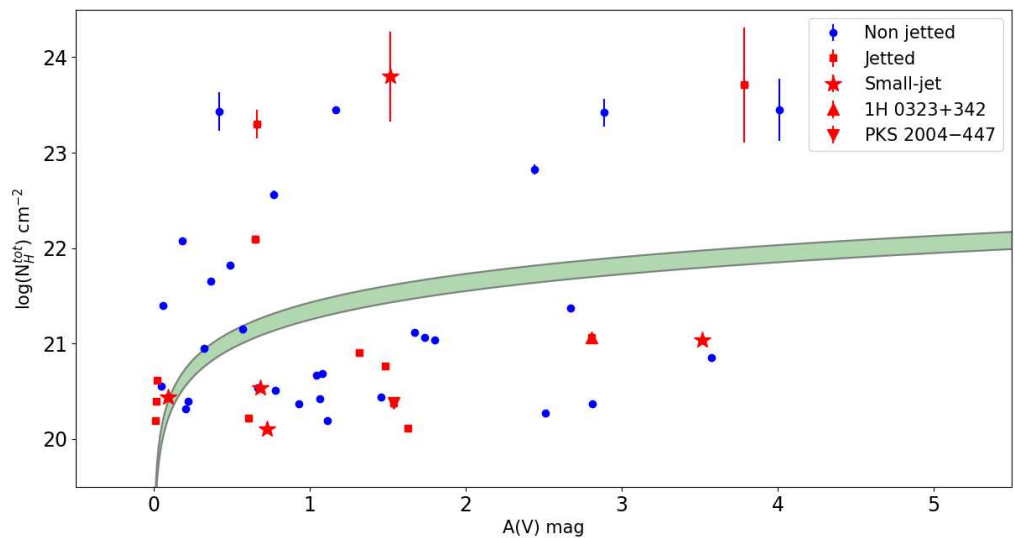


Figure 2. The relationship between the hydrogen abundance N_H^{tot} and the extinction parameter $A(V)$. The green area marked, sourced from [26,27], refers to the Galactic standard ratio.

3.2. $[\text{O III}]/\text{H}\beta$ versus N_H^{tot}

The parameter $R = [\text{O III}]/\text{H}\beta$ emerges as a better indicator for determining Seyfert types compared to the $\text{H}\alpha$ -based method, even at lower types like Sy1.2/Sy1.0, as highlighted in Figure 6 of [6]. In that plot, the relationship between the two approaches holds for Seyfert types $>\text{Sy}1.5$, while the scheme proposed by [18] tends to over-classify Sy1.0/1.2.

In the UM, the separation between obscured and unobscured AGN is typically indicated by an N_H^{tot} value around 10^{22} atoms cm^{-2} , as shown by the horizontal solid line in Figure 3. However, discriminant parameters can vary across different studies, for instance, using the dashed line from [29] or the dashed-dotted line from [28]. The vertical lines in the plot indicate Seyfert sub-classes based on the methodology presented in [17].

As anticipated, the objects concentrate within two yellow regions: one with $\log(N_H^{\text{tot}}) < 22$, $R < 1$ and another with $\log(N_H^{\text{tot}}) > 22$, $R > 4$. The faint yellow region corresponds to Sy1.5 ($1 < R < 4$), and it is interpreted as a transition zone between obscured and unobscured sources. A similar analysis was conducted by the authors of [28] in their Figure 2, although their horizontal axis is more refined due to a grouping process into the main sub-types. Nevertheless, our findings align with those reported in [28,30], confirming an ascending trend between the Seyfert type and the column density. In the latter study, the authors demonstrated that Sy1.8/1.9 typically exhibit $\log(N_H^{\text{tot}}) \sim 23\text{--}24$, while the distribution for Sy2.0 is centered at higher absorption columns (see Figure 5 of [30]).

Figure 3 reveals that the majority of jetted sources reside in the lower-left part while non-jetted sources occupy both the upper-right and the lower-left part of the figure. This result partially aligns with the UM. However, three jetted objects lie in the high-obscuration region. Their estimation of the N_H^{tot} parameter is not as reliable as in other cases, possibly due to low-resolution X-ray spectra. These objects include 4C +29.30 and 3C 234.0, both radio sources, the latter involving a complicated fitting process (refer to Section 2.2), and 2MASX J04234080 + 0408017 with small radio jets [31]. For radio-emitting objects, significant obscuration is not contradictory to the UM. Indeed, radio emission can be divided into two components: core emission, which is enhanced along the observer's line of sight, akin to γ -rays; and extended emission, often called lobe emission, which is not beamed. Furthermore, some authors have [32] stressed the limits in classifying the AGN adopting a radio-loud/radio-quiet scheme. In this view, there is a continuum between the two classes which can explain the ambiguity of some objects in Figure 3. The non-jetted sources in the lower-left region do not pose a significant challenge to the UM, as not all poorly-obscured sources necessarily exhibit jets. This ambiguity stems from the selection method adopted, biased toward lower Seyfert types for the jetted group and higher Seyfert types for the non-jetted group, but this does not result in a distinct separation.

Even if it is not possible to find a clear division between the two identified classes in this study, our findings broadly concur with previous analyses. Furthermore, eliminating sources with zero intrinsic absorption yields analogous results. The only noticeable difference is the absence of objects in the Sy1 class. This suggests that the Galactic contribution to the obscuration does not significantly alter the conclusions. See Figure A2 in Appendix B for the comparison plots.

The statistical tests performed result in the following: Spearman ($r_s = 0.31$, p -value = 3.56×10^{-2}), Pearson ($r_p = 0.55$, p -value = 5.25×10^{-5}), and Kendall ($r_k = 0.20$, p -value = 4.36×10^{-2}).

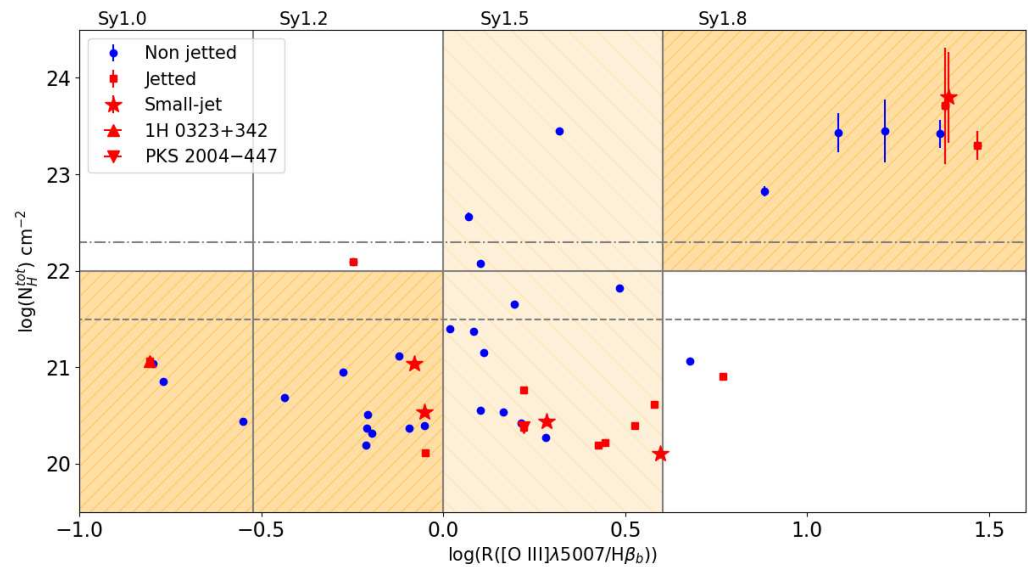


Figure 3. The Seyfert type based on the $[\text{O III}]/\text{H}\beta$ ratio against the hydrogen abundance. A solid horizontal grey line distinguishes between unobscured and obscured objects, using a threshold of 10^{22} atoms cm^{-2} . Additional lines—dashed [29] and dashed-dotted [28]—serve similar purposes. Vertical lines separate Seyfert sub-types according to Equation (1) and Table 1 in [6]. Yellow areas enhance the separation.

3.3. $L_{[\text{O III}]}$ versus N_H^{tot}

The relationship between the $[\text{O III}]$ luminosity ($L_{[\text{O III}]}$) and the column density aims to compare the intrinsic properties of AGN, represented by $L_{[\text{O III}]}$ (as detailed in [6]), and obscuration, N_H^{tot} . Comparing Figure 4 with the similar panel in the previous study, it

is evident that the results obtained from optical-based methods align closely with those derived from X-ray-based approaches. In both cases, the objects are distributed without a distinct separation between the classes. Similar to the previous section, the exclusion of objects with $N_H^{int} = 0$ does not impact the results; in this case as well, the sources tend to mix and do not exhibit a clear separation between the classes. See Figure A2 in Appendix B for the details.

Consequently, the orientation-dependent explanation proposed by the UM remains the most viable solution, given the lack of substantial correlation between the intrinsic properties of AGN and obscuration levels across different classes.

The statistical tests give the following results: Spearman ($r_s = 0.08$, p -value = 5.91×10^{-1}), Pearson ($r_p = 0.11$, p -value = 4.78×10^{-1}), and Kendall ($r_k = 0.06$, p -value = 5.57×10^{-1}).

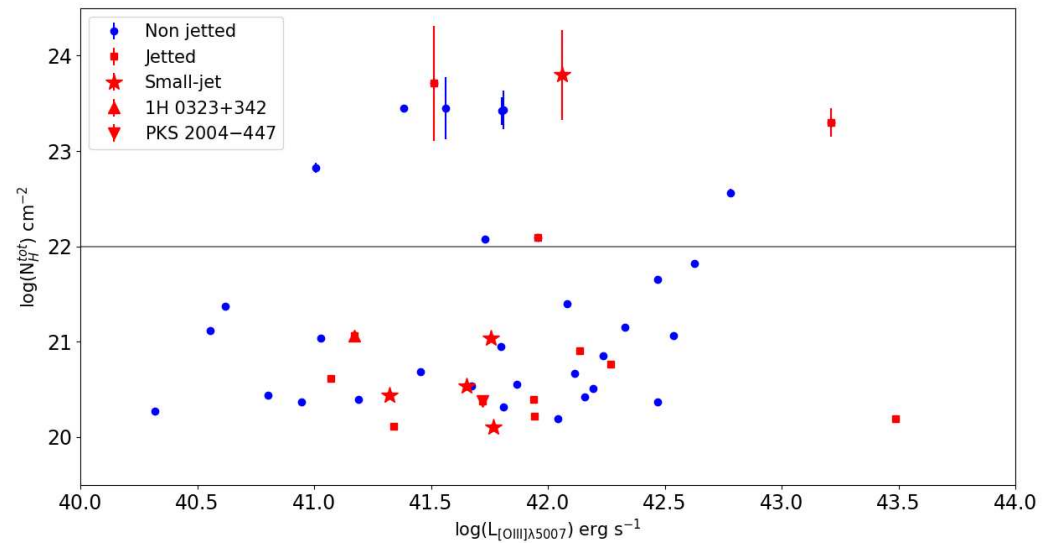


Figure 4. The comparison between the [O III]λ5007 luminosity and the column density. Similar to Figure 3, a horizontal line denotes the separation highlighted in the central panel, likely using the same threshold criteria.

3.4. Luminosities

The plot reported in Figure 5 compares the X-ray luminosity between 2 and 10 keV ($L_{X[2-10\text{keV}]}$) with $L_{[\text{O III}]}$. Within the same figure, three reference lines are presented: our dataset's linear regression with the dotted black line, the relation from [23] indicated by the dashed grey line, and the dashed-dotted line from [24].

Previous studies [9] explored the L_X - $L_{[\text{O III}]}$ comparison, attempting to distinguish between Sy1 and Sy2. However, even in their analysis, the objects showed a mix within the two classes. These authors emphasized the selection-dependent nature of this relation, noting its robustness for hard-X-ray-selected objects but its weaker correlation for optically selected AGN. In our study, where the objects are hard-X-ray-selected, we support the initial findings of [9].

Focusing on the steepness of the correlations observed, we find different relationships: $L_{[\text{O III}]} \propto 7.34 L_X^{0.82}$ for [24], $L_{[\text{O III}]} \propto -1.30 L_X^{0.96}$ for [23], and $L_{[\text{O III}]} \propto 15.74 L_X^{0.60}$ in our case. The discrepancies between our findings and those of [23] might arise from differences in the nature of the samples studied. In our dataset, there is a mixture of Seyfert types, whereas [23] focused solely on type 1 AGN. Additionally, their analysis encompassed 3579 AGN, a considerably larger sample size compared to our roughly 50 objects. However, the sample size difference between [24] (60 AGN) and our study is not significant. Neither [24] nor [9] observed a clear separation between Sy1 and Sy2, akin to the jetted/non-jetted classification in our case. This lack of distinction might stem from the limited number of sources considered in both studies.

The two outliers in the bottom part of the plot correspond to 2MASX J01492228 – 5015073 and GB6 J0937 + 5008. The optical spectrum of the former exhibits irregularities at different wavelengths, which might influence the process. However, it is important to note that this source is not far from the linear regression identified in this study. Regarding the second case, we have a limited number of bins in the X-ray analysis. This situation can potentially lead to a displacement of the datapoints towards either lower or higher values of $L_{X[2-10\text{keV}]}$, due to the challenge in accurately constraining the parameter within the analysis process.

Nevertheless, despite these variations, there is a consensus among these studies about the presence of a relationship between the two luminosities involved, emphasizing the correlation between X-ray and [O III] luminosities across different samples of AGN.

The results for the statistical tests performed are as follows: Spearman ($r_s = 0.52$, $p\text{-value} = 1.52 \times 10^{-4}$), Pearson ($r_p = 0.54$, $p\text{-value} = 8.15 \times 10^{-5}$), and Kendall ($r_k = 0.40$, $p\text{-value} = 6.14 \times 10^{-5}$).

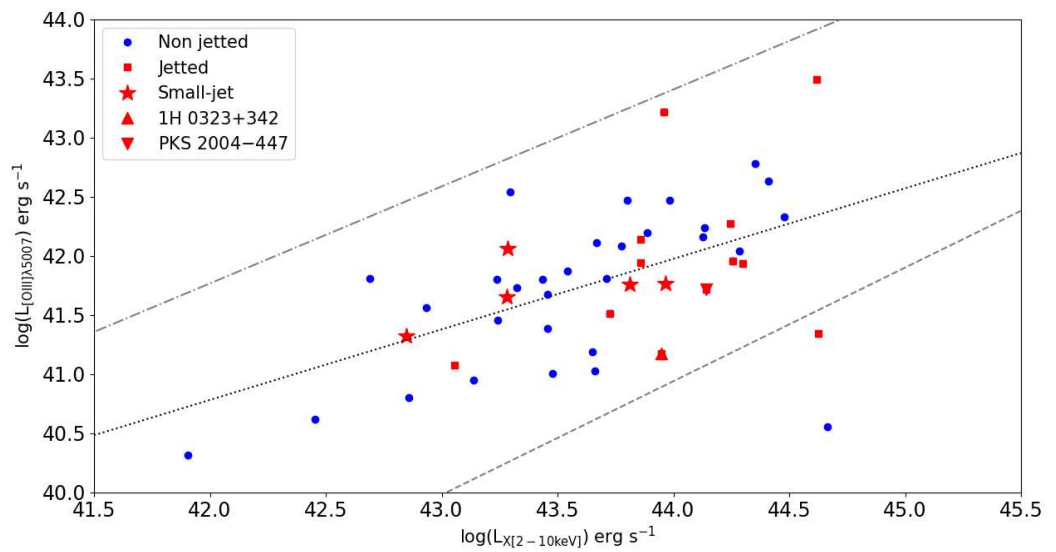


Figure 5. The plot compares the X-ray luminosity with the oxygen luminosity. The dotted black line represents the linear regression line based on our dataset, while the dashed line corresponds to the regression line calculated by [23]. Additionally, the dashed–dotted line reflects the regression analysis conducted by [24].

4. Summary and Conclusions

This work serves as the continuation of [6], focusing on the X-ray spectral analysis after analyzing the optical spectra of a combined sample of hard-X-ray- and γ -ray-selected sources. Initially, our goal was to establish a connection between the intrinsic properties of IS and the inclination-dependent view proposed by the UM. In [6], we utilized the Balmer decrement and the Seyfert type derived from the $H\beta/[O\text{ III}]$ ratio introduced by [17] to estimate the obscuration. However, these attempts did not reveal a significant trend or distinct separation between the two classes of sources.

To deepen this analysis, we turned to the X-ray spectra to derive a more reliable measure of obscuration represented by N_H^{tot} . The results, illustrated in Figure 4, mirror the optical-based analysis from Figure 6 of [6]. This further confirms the absence of a direct relation between the intrinsic properties of IS and obscuration. To further validate this result, we attempted to exclude objects with $N_H^{\text{int}} = 0$ from the sample, considering only those sources exhibiting an intrinsic component of obscuration. Nevertheless, all the conclusions drawn in the first case apply to this second scenario as well, emphasizing that the Galactic contribution is marginal.

In Figure 3, we observe a separation of IS types based on an N_H^{tot} threshold around 10^{22} atoms cm^{-2} with a transition region corresponding to Sy1.5. This supports the UM's perspective, wherein an increase in Seyfert type correlates with a gradual increase in obscuration, particularly when also considering ISs.

Moving to Figure 2, comparing optical extinction with column density, our analysis is constrained due to the limited horizontal axis. Consequently, our findings lack substantial data for objects with significant extinction. However, they concur with [28], where low-extinction points are not significantly distributed above the Galactic reference ratio in green.

In Figure 5, depicting X-ray and oxygen luminosity, our observations display a slightly different trend compared to [23,24]. This difference in trend might be attributed to variations in the number or nature of sources analyzed across the three studies. Nonetheless, it reinforces the interpretation of oxygen luminosity as a measure of disk luminosity and, by extension, intrinsic properties, supporting the notion that the X-ray-emitting region is central and closely linked to the accreting flow.

In summary, our comprehensive multi-wavelength analysis does not offer clear evidence linking IS type to accretion flow properties. However, we confirm the connection between L_X - $L_{[\text{O III}]}$ and the differentiation between Seyfert types and column density.

ISs might play a pivotal role in understanding other Seyfert families, but the analysis of large sample sizes can dilute interesting variations between individual objects. This suggests that a more detailed study of individual sources could provide a more robust investigation into the nature of ISs as detailed by [7,33].

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Data Availability Statement: A complete list of the BASS sources can be found in the survey archive (accessed on 28 January 2024) <http://www.bass-survey.com/data.html>. The *Swift*-BAT, *Chandra*, and *XMM-Newton* X-ray spectra are available in the HEASARC archive <https://heasarc.gsfc.nasa.gov/cgi-bin/W3Browse/w3browse.pl> (accessed on 28 January 2024).

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Conflicts of Interest: The authors declare no conflicts of interest.

Appendix A

In Table A1, we reported the list of sources and their main characteristics with the observation IDs. In Table A2, we reported the fitting parameters.

Table A1. The list of sources with the coordinates (RA, DEC), the redshift (z), the column density (N_H), the observation IDs, the satellite from which the observations were carried (for *XMM-Newton* we use XMM as abbreviation), the exposure time, and the starting observation date.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10^{-3}]	N_H [10^{20} cm^{-2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
BASS									
J0149 – 5015	2MASX J01492228 – 5015073	27.34	−50.25	2.99	1.75	00038018001	<i>Swift</i>	10.58	2008-09-11
J0157 + 4715	2MASX J01571097 + 4715588	29.30	47.27	4.79	11.0	00040895001	<i>Swift</i>	6.94	2010-10-22
						00040895002		1.61	2010-11-19
J0206 – 0017	Mrk 1018	31.57	−0.29	4.27	2.48	00035166001	<i>Swift</i>	5.25	2005-08-05
						00035776001		4.86	2008-06-11
J0214 – 0046	Mrk 590	33.64	−767.00	2.63	2.77	00095662033	<i>Swift</i>	9.93	2021-01-10
						00037590015		8.29	2016-12-13
J0234 – 0847	NGC 985	38.66	−8.79	4.30	3.48	00036530005	<i>Swift</i>	9.58	2008-06-06
						00096449002		7.99	2021-06-27
J0238 – 4038	2MASX J02384897 – 4038377	39.70	−40.64	6.13	2.06	00049419004	<i>Swift</i>	4.76	2016-03-27
						00037344002		4.13	2008-03-27
						00037344003		3.08	2008-04-04
						00037344001		2.71	2008-03-13
						00049419001		1.69	2013-03-29
						00093020001		1.12	2017-04-13
						00049419002		0.88	2013-12-27
						00049419003		0.58	2015-09-06
						00049419005		0.57	2016-04-01
J0312 + 5029	2MASX J03120291 + 5029147	48.01	50.49	6.15	3.45	00038026002	<i>Swift</i>	6.83	2008-12-07
						00038026001		1.62	2008-10-11
J0330 + 0538	2MASX J03305218 + 0538253	52.72	5.64	4.58	1.16	00039819001	<i>Swift</i>	8.86	2009-07-04
						00039819002		2.51	2009-07-15
						00047706001		0.56	2013-01-03
						00047706002		0.40	2013-01-09
J0333 + 3718	2MASX J03331873 + 3718107	53.33	37.30	5.47	1.48	00040693001	<i>Swift</i>	6.41	2011-04-03
						00031490001		5.02	2009-09-16
						00047806001		1.47	2013-01-06
J0423 + 0408	2MASX J04234080 + 0408017	65.92	4.13	4.62	12.6	13897	<i>Chandra</i>	20.06	2012-10-20

Table A1. Cont.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10^{-3}]	N_H [10^{20} cm^{-2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
J0503 + 2300	LEDA 097068	75.74	23.00	5.81	23.4	00037117001	Swift	9.26	2007-10-17
						00037117002		0.81	2008-02-17
J0605 – 2754	2MASX J06054896 – 2754398	91.45	–27.91	8.98	2.64	00090155001	Swift	3.89	2009-07-06
						00040696003		3.43	2011-05-22
						00040696001		1.95	2011-01-30
						00047747001		1.31	2013-06-24
J0733 + 4555	1RXS J073308.7 + 455511	113.29	45.92	14.15	7.19	00041758001	Swift	6.56	2010-12-19
						00041758002		3.23	2010-12-28
						00081466001		1.99	2015-12-26
						00047790002		1.08	2015-05-25
						00041758003		1.00	2010-12-29
						00047790001		0.39	2015-05-24
J0736 + 5846	Mrk 9	114.24	58.77	3.99	4.83	00080535001	Swift	6.49	2013-10-29
						00041759002		4.80	2010-12-17
						00041759005		2.69	2021-09-22
						00041759004		2.39	2010-12-23
						00041759003		1.81	2010-12-19
						00041759001		1.74	2010-12-07
						07021915001		0.18	2019-10-01
J0742 + 4948	Mrk 79	115.64	49.81	2.21	5.43	00037589002	Swift	2.92	2009-02-26
						00037589003		2.83	2021-09-22
						00037589001		2.34	2009-02-19
						00096121001		2.17	2021-09-01
J0752 + 1935	2MASX J07521780 + 1935423	118.07	19.60	11.70	4.06	00039551001	Swift	5.10	2010-05-30
						00038041001		4.20	2008-10-31
						00038041004		2.91	2010-05-30
						00038041003		1.74	2009-09-12
						00038041002		1.37	2009-02-24
J0803 + 0841	2MASX J08032736 + 0841523	120.86	8.70	4.68	2.87	00038042002	Swift	6.60	2010-09-15
J0804 + 0506	Mrk 1210	121.02	5.11	1.35	3.86	00037233002	Swift	9.61	2010-09-10
						00080387001		6.00	2012-10-05
						00035588001		2.33	2006-10-06
						00035588003		2.32	2007-02-25
						00035588002		1.52	2007-01-22
						00037233001		1.26	2008-04-29

Table A1. Cont.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10 ^{−3}]	N _H [10 ²⁰ cm ^{−2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
J0829 + 4154	2MASX J08294266 + 4154366	127.43	41.91	12.61	3.57	00040937004	Swift	4.41	2011-01-17
J0832 + 3707	FBQS J083225.3 + 370736	128.11	37.13	9.20	3.27	00035634002	Swift	4.76	2007-02-17
						00035634003		2.51	2007-02-24
						00035634004		1.82	2007-05-01
						00035483003		1.69	2021-02-14
						00035483001		1.55	2006-01-08
						00035483002		1.12	2011-09-09
						00035483004		1.10	2021-03-14
						00035634001		0.96	2007-01-08
J0840 + 2949	4C +29.30	130.01	29.82	6.47	4.56	11688	Chandra	125.09	2010-02-19
J0842 + 0759	2MASX J08420557 + 0759253	130.52	7.99	13.37	5.76	00040938004	Swift	7.13	2011-01-18
						00040938002		1.14	2010-12-31
						00040938003		0.98	2011-01-09
						00040938001		0.50	2010-12-18
J0843 + 3549	2MASX J08434495 + 3549421	130.94	35.83	5.39	2.98	18143	Chandra	25.07	2016-12-26
J0904 + 5536	2MASX J09043699 + 5536025	136.15	55.60	3.72	2.32	00035260002	Swift	7.76	2006-01-06
						00035260001		6.08	2005-12-15
						07000264001		0.06	2016-01-13
J0918 + 1619	Mrk 704	139.61	16.31	2.95	2.74	00080414001	Swift	6.60	2014-12-28
						00035590003		5.60	2006-09-28
						00035590002		2.26	2006-06-14
						00031965001		1.85	2011-04-23
						00031965003		1.85	2022-04-23
						00035590005		1.75	2007-01-21
						00015881001		1.60	2023-02-20
						00031965002		1.14	2011-04-23
						00080414002		1.05	2021-03-17
						00035590001		0.67	2006-01-06
J0923 + 2255	MCG +04 − 22 − 042	140.93	22.91	3.30	3.60	00035263001	Swift	9.20	2005-12-10
						00080416001		3.95	2012-12-26
						00089106003		1.93	2021-03-20
						00089106001		1.92	2020-06-05
						00089106004		1.68	2021-05-30
						00089106002		1.66	2020-10-20
						00035263002		1.02	2019-12-25

Table A1. Cont.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10^{-3}]	N_H [10^{20} cm^{-2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
J0925 + 5219	Mrk 110	141.30	52.29	3.52	1.27	00037561001 03111716005 00092396006 00092396007	Swift	10.42 4.68 4.12 4.10	2010-01-06 2022-09-15 2016-04-29 2016-04-30
J0926 + 1245	Mrk 705	141.51	12.73	2.86	3.43	00090998001 00085341001 00085341004 00085341002	Swift	1.77 1.01 0.91 0.90	2011-05-24 2015-01-02 2015-04-05 2015-01-17
J0935 + 2617	2MASX J09352707 + 2617093	143.86	26.29	12.21	1.56	00040947001 00040947003 00040947002	Swift	5.82 3.24 1.57	2011-01-02 2011-01-12 2011-01-11
J0942 + 2341	CGCG 122 – 055	145.52	23.69	2.13	2.42	00090166001	Swift	5.37	2009-11-25
J0945 + 0738	3C 227	146.94	7.42	8.60	2.00	00080695001 00080695002 00080695003	Swift	3.17 2.40 2.08	2014-02-20 2014-02-21 2014-02-26
J0959 + 1302	NGC 3080	149.98	13.04	3.54	2.25	00081059001 00037365004 00037365003 00037365001 00037365002	Swift	6.42 4.32 3.60 1.49 0.97	2017-10-21 2008-10-10 2008-10-09 2008-06-14 2008-06-30
J1001 + 2847	3C 234.0	150.46	28.79	18.48	1.62	0864430101	XMM	95.20	2021-05-22
J1023 + 1951	NGC 3227	155.88	19.87	0.33	1.86	00092212006 00092212011 00092212013 00092212016 00092212014	Swift	4.16 3.96 3.91 3.87 3.68	2015-05-09 2015-10-25 2015-11-08 2015-11-29 2015-11-15
J1043 + 1105	SDSS J104326.47 + 110524.2	160.86	11.09	4.77	2.38	00040954001 00081072001 00088253001	Swift	9.95 5.68 4.79	2010-10-29 2016-06-14 2019-05-06

Table A1. Cont.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10 ^{−3}]	N _H [10 ²⁰ cm ^{−2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
J1315 + 4424	UGC 08327	198.82	44.41	3.55	1.68	00037093001 00080109001 00037093003 00080109002 00080109003 00088636002 00037093002 00088636001	Swift	11.49 6.64 2.45 2.01 1.94 1.62 1.23 0.31	2007-06-03 2013-04-21 2007-09-20 2013-11-18 2013-11-23 2017-11-19 2007-09-19 2017-11-15
J1445 + 2702	CGCG 164 − 019	221.40	27.03	2.96	2.37	0865140901	XMM	27.00	2021-01-19
J1508 − 0011	Mrk 1393	227.22	−197.00	5.44	4.64	00030309001 00088254003 00088254001 00088254002	Swift	9.76 7.11 6.04 5.76	2005-09-10 2019-03-02 2018-03-27 2018-05-13
4FGL									
J0038 − 0207	3C 17	9.59	−2.13	22.04	2.48	9292	Chandra	8.04	2008-02-02
J0324 + 3410	1H 0323+342	51.17	34.18	6.29	11.7	00035630002 00035372001 00035630003	Swift	8.96 8.36 6.87	2006-07-09 2006-07-06 2006-07-14
J0937 + 5008	GB6 J0937 + 5008	144.30	50.15	27.55	1.31	00040543003 00040543002 00040543001	Swift	2.99 1.08 0.85	2012-09-06 2011-03-30 2011-03-30
J0958 + 3224	3C 232	149.59	32.40	53.06	1.57	00010070001	Swift	1.96	2017-04-29
J1443 + 5201	3C 303	220.76	52.03	14.12	1.66	1623	Chandra	15.10	2001-03-23
J1516 + 0015	PKS 1514+00	229.17	0.25	5.26	4.17	00088374001 00036340002 00036340003 00036340006 00036340004 00036340007	Swift	6.53 6.28 4.59 2.41 2.38 1.25	2020-03-15 2007-06-22 2008-01-06 2018-12-23 2010-09-25 2018-12-26
J2007 − 4434	PKS 2004 − 447	301.98	−44.58	24.00	2.97	00032492015 00032492006 00032492005 00032492017 00032492007	Swift	12.26 11.62 11.45 9.70 8.42	2013-11-20 2013-07-14 2013-07-07 2014-03-16 2013-09-27

Table A1. Cont.

IAU Name	Alias	RA [deg J2000]	DEC [deg J2000]	z [10^{-3}]	N_H [10^{20} cm^{-2}]	Obs. ID(s)	Satellite	Exposure [ks]	Date (YYYY-MM-DD)
J2118 + 0013	PMN J2118 + 0013	319.57	0.22	46.28	5.87	00091075002	Swift	3.58	2011-08-11
						00040665003		2.69	2012-08-26
						00091075006		2.14	2011-08-22
						00091075004		1.83	2011-08-16
						00040665004		1.56	2012-09-02
						00091075001		0.76	2011-05-29
						00040665001		0.57	2011-08-27
						00091075003		0.48	2011-08-15
						00091075005		0.47	2011-08-17
						00040665002		0.31	2011-08-30
J2118 – 0732	TXS 2116 – 077	319.72	−7.54	26.01	7.99	00011263004	Swift	3.79	2019-08-05
						00011263005		3.79	2019-09-05
						00011263003		3.37	2019-07-05
						00011263006		2.71	2019-10-11
						00011263001		2.55	2019-05-05
						00011263002		2.18	2019-06-05
						00011263008		1.27	2019-10-21
						00011263009		1.21	2019-10-22

Table A2. The fitting parameters are as follows: Γ_1 and Γ_2 for the power-law indexes, E_{break} for the break energy for the broken power-law cases, E_{line} for the Fe $K\alpha$ line energy (when present), kT for the two *apec* functions used in the fitting of 3C 234.0, χ^2/dof for the goodness-of-fit in terms of the χ^2 and the degrees of freedom (dof), the adopted model according to the abbreviations listed in Table A3, $F_{2-10\text{keV}}^{\text{obs}}$ for the observed flux in units of $10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$, $L_{2-10\text{keV}}^{\text{int}}$ for the intrinsic luminosity removing the absorption components coming from galactic and internal factors in units of $10^{43} \text{ erg s}^{-1}$, and N_H^{int} for the additional obscuration when present in units of $10^{20} \text{ atoms cm}^{-2}$. With (^a), we highlight the fixed parameters in the fit and with (^b) the results for the fitting of 3C 232 based on the rare-events statistics (*c-stat*).

IAU Name	Γ_1	Γ_2	E_{break} [keV]	E_{line} [keV]	kT_1 [keV]	kT_2 [keV]	χ^2/dof	Model	$F_{2-10\text{keV}}^{\text{obs}}$ [$10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$]	$L_{2-10\text{keV}}^{\text{int}}$ [$10^{43} \text{ erg s}^{-1}$]	N_H^{int} [10^{20} cm^{-2}]
BASS											
J0149 – 5015	$2.00^{+0.19}_{-0.18}$	-	-	-	-	-	33/38	ABS+PL	394.05	0.99	$0.11^{+0.04}_{-0.04}$
J0157 + 4715	$2.87^{+2.86}_{-1.25}$	$1.98^{+0.06}_{-0.06}$	$1.24^{+0.16}_{-0.07}$	-	-	-	113/104	PL	14.57	9.24	0.0
J0206 – 0017	$1.90^{+0.04}_{-0.04}$	-	-	-	-	-	198/184	PL	21.07	8.88	0.0
J0214 – 0046	$1.61^{+0.05}_{-0.05}$	-	-	-	-	-	104/111	PL	7.48	1.13	0.0
J0234 – 0847	$3.53^{+0.13}_{-0.12}$	$1.60^{+0.04}_{-0.04}$	$0.94^{+0.03}_{-0.03}$	-	-	-	293/262	ABS+BPL	25.99	0.15	$7.43^{+0.03}_{-0.03}$
J0238 – 4038	$1.79^{+0.04}_{-0.04}$	-	-	-	-	-	181/180	PL	10.94	9.35	0.0
J0312 + 5029	$1.72^{+0.17}_{-0.16}$	-	-	-	-	-	52/48	ABS+PL	9.11	0.11	$0.11^{+0.08}_{-0.08}$
J0330 + 0538	$1.71^{+0.08}_{-0.08}$	-	-	-	-	-	62/59	PL	6.31	3.38	0.0
J0333 + 3718	$1.57^{+0.10}_{-0.10}$	-	-	-	-	-	87/90	ABS+PL	11.48	9.26	$0.10^{+0.05}_{-0.04}$
J0423 + 0408	$2.59^{+0.23}_{-0.22}$	1.90^a	-	-	-	-	61/33	OBS	1.51	8.28	$62.84^{+11.22}_{-9.72}$
J0503 + 2300	$1.23^{+0.04}_{-0.04}$	-	-	-	-	-	149/149	PL	19.71	0.18	0.0
J0605 – 2754	$1.47^{+0.05}_{-0.05}$	-	-	-	-	-	81/88	PL	10.85	0.20	0.0
J0733 + 4555	$2.07^{+0.07}_{-0.07}$	-	-	-	-	-	91/78	PL	5.19	0.32	0.0
J0736 + 5846	$1.99^{+0.04}_{-0.04}$	-	-	-	-	-	231/170	PL	10.06	3.77	0.0
J0742 + 4948	$2.14^{+0.12}_{-0.11}$	$1.31^{+0.33}_{-0.35}$	-	-	-	-	240/212	OBS	30.96	3.92	$1.14^{+0.80}_{-0.53}$
J0752 + 1935	1.90^a	$1.51^{+1.11}_{-0.51}$	-	-	-	-	32/34	OBS	6.02	0.33	$3.60^{+1.07}_{-0.89}$
J0803 + 0841	$1.51^{+0.14}_{-0.13}$	-	-	-	-	-	62/54	ABS+PL	7.85	4.31	$0.17^{+0.07}_{-0.06}$
J0804 + 0506	$2.29^{+0.18}_{-0.18}$	$1.86^{+0.74}_{-0.66}$	-	$6.25^{+0.19}_{-0.11}$	-	-	34/41	OBS+FE	7.54	1.50	$28.26^{+7.65}_{-6.51}$
J0829 + 4154	$1.55^{+0.11}_{-0.10}$	-	-	-	-	-	57/74	ABS+PL	11.09	0.46	$0.11^{+0.04}_{-0.04}$
J0832 + 3707	$1.76^{+0.07}_{-0.07}$	-	-	-	-	-	85/76	PL	6.71	0.14	0.0
J0840 + 2949	$2.98^{+0.43}_{-0.42}$	$1.72^{+0.93}_{-0.80}$	-	-	-	-	46/29	OBS	1.26	9.12	$51.44^{+14.30}_{-11.34}$
J0842 + 0759	$1.36^{+0.21}_{-0.19}$	-	-	-	-	-	24/36	ABS+PL	6.97	0.36	$0.60^{+0.21}_{-0.17}$
J0843 + 3549	$2.42^{+0.70}_{-0.76}$	1.90^a	-	-	-	-	44/36	OBS	1.52	5.38	$26.30^{+3.44}_{-2.90}$
J0904 + 5536	$1.49^{+0.06}_{-0.06}$	-	-	-	-	-	97/75	BPL	7.40	2.20	0.0

Table A2. Cont.

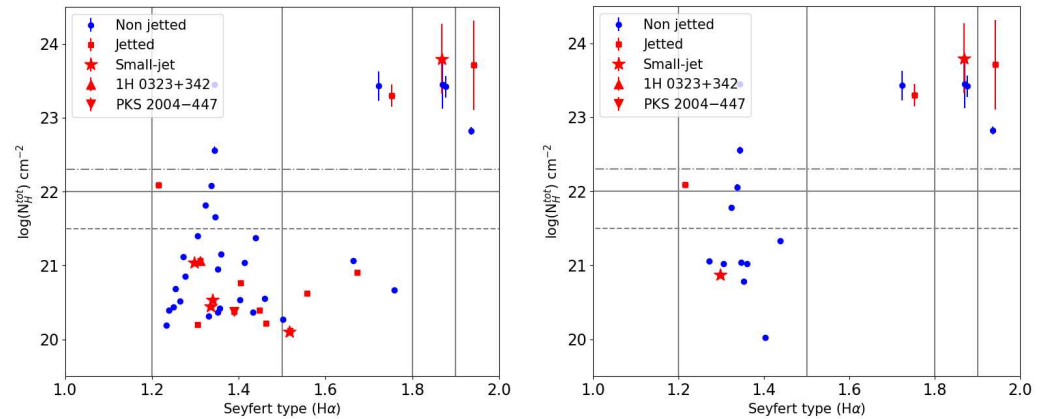
IAU Name	Γ_1	Γ_2	E_{break} [keV]	E_{line} [keV]	kT_1 [keV]	kT_2 [keV]	χ^2/dof	Model	$F_{2-10\text{keV}}^{\text{obs}}$ [10^{-12} erg cm $^{-2}$ s $^{-1}$]	$L_{2-10\text{keV}}^{\text{int}}$ [10^{43} erg s $^{-1}$]	N_H^{int} [10^{20} cm $^{-2}$]
J0918 + 1619	$3.29^{+0.38}_{-0.37}$	$1.34^{+0.06}_{-0.06}$	$0.95^{+0.05}_{-0.05}$	-	-	-	251/215	ABS+BPL	14.54	3.14	$6.10^{+0.03}_{-0.03}$
J0923 + 2255	$2.01^{+0.06}_{-0.05}$	$1.76^{+0.04}_{-0.04}$	$1.29^{+0.23}_{-0.20}$	-	-	-	363/297	BPL	26.24	6.59	0.0
J0925 + 5219	$2.09^{+0.05}_{-0.04}$	$1.60^{+0.04}_{-0.04}$	$1.29^{+0.13}_{-0.13}$	-	-	-	397/373	BPL	60.55	0.16	0.0
J0926 + 1245	$2.05^{+0.06}_{-0.06}$	-	-	-	-	-	103/82	PL	22.47	4.37	0.0
J0935 + 2617	1.79^a	-	-	-	-	-	84/98	PL	9.65	0.35	0.0
J0942 + 2341	$1.47^{+0.13}_{-0.13}$	-	-	-	-	-	30/19	ABS+BPL	4.23	1.17	$0.21^{+0.13}_{-0.13}$
J0945 + 0738	$1.08^{+0.25}_{-0.24}$	-	-	-	-	-	64/39	OBS	11.81	0.25	$1.22^{+0.31}_{-0.26}$
J0959 + 1302	$1.88^{+0.07}_{-0.07}$	-	-	-	-	-	74/64	PL	5.08	1.41	0.0
J1001 + 2847	$1.78^{+0.22}_{-0.27}$	$1.08^{+0.18}_{-0.17}$	-	$6.38^{+0.02}_{-0.02}$	$0.16^{+0.01}_{-0.01}$	$0.81^{+0.05}_{-0.05}$	483/369	2APEC+FE	1.37	0.22	$20.02^{+3.53}_{-3.00}$
J1023 + 1951	$1.01^{+0.05}_{-0.05}$	$1.56^{+0.07}_{-0.06}$	$1.92^{+0.24}_{-0.18}$	-	-	-	398/364	BPL	38.11	11.25	0.0
J1043 + 1105	$1.63^{+0.07}_{-0.07}$	-	-	-	-	-	118/137	ABS+PL	9.00	4.67	$1.05^{+0.02}_{-0.01}$
J1315 + 4424	$2.32^{+0.95}_{-0.91}$	$1.66^{+0.29}_{-0.27}$	-	-	-	-	54/66	OBS	7.77	5.11	$6.67^{+1.22}_{-1.02}$
J1445 + 2702	$2.68^{+0.10}_{-0.10}$	$1.74^{+0.34}_{-0.32}$	-	-	-	-	173/125	OBS	1.20	1.03	$26.96^{+4.65}_{-4.22}$
J1508 – 0011	$2.33^{+0.16}_{-0.15}$	$1.44^{+0.05}_{-0.05}$	$0.96^{+0.11}_{-0.10}$	-	-	-	167/168	BPL	10.81	7.48	0.0
4FGL											
J0038 – 0207	$1.86^{+0.42}_{-0.38}$	$1.27^{+0.09}_{-0.16}$	$0.98^{+0.80}_{-0.26}$	-	-	-	77/66	BPL	2.36	0.28	0.0
J0324 + 3410	$3.10^{+0.30}_{-0.53}$	$1.83^{+0.08}_{-0.11}$	$1.29^{+0.22}_{-0.07}$	-	-	-	261/231	PL	15.76	0.28	0.0
J0937 + 5008	$1.44^{+0.16}_{-0.16}$	-	-	-	-	-	8/11	PL	3.02	0.61	0.0
J0958 + 3224	$2.03^{+0.99}_{-0.80}$	-	-	-	-	-	13/22 ^b	PL	0.78	0.71	0.0
J1443 + 5201	$1.24^{+0.05}_{-0.05}$	-	-	-	-	-	148/121	PL	2.11	9.43	0.0
J1516 + 0015	$1.57^{+0.07}_{-0.07}$	-	-	-	-	-	43/65	PL	2.78	1.77	0.0
J2007 – 4434	$1.48^{+0.07}_{-0.07}$	-	-	-	-	-	46/54	PL	1.37	0.20	0.0
J2118 + 0013	$1.94^{+0.28}_{-0.27}$	-	-	-	-	-	5/3	PL	0.43	24.74	0.0
J2118 – 0732	$1.51^{+0.21}_{-0.20}$	-	-	-	-	-	11/8	PL	0.57	0.11	0.0

Table A3. List of adopted models and the abbreviations in Table A2.

Abbreviation	Extended Name	Model
PL	Power-law	tbabs*zpo
BPL	Broken power-law	tbabs*bknpo
ABS	Intrinsic absorption	tbabs*ztbabs*zpo
		tbabs*ztbabs*bknpo
OBS	Obscured-type emission	tbabs*(zpo+ztbabs*zpo)
FE	Iron line	(One of the previous)+zgauss
APEC	Collisionally ionized diffuse gas	(One of the previous)+apec

Appendix B

In Figure A1, we provide an alternative representation of Figure 3, utilizing the Seyfert type classification determined solely by the properties of the $H\alpha$ line [18]. However, as explained in [6], the method based on $H\beta/[O III]$ ratio proposed by [17] is more effective in classifying objects at lower Seyfert types. This is evident in Figure A1, where none of the objects fall within the Sy1 region. Notably, even 1H 0323+342 and PKS 2004–447, known as two narrow-line Seyfert 1, are not classified within the Sy1 region using the $H\alpha$ -based method. The same plot has been generated by excluding objects with $N_H^{int} = 0$. From this alternative approach, we can observe that the conclusions drawn in the first case also hold here.

**Figure A1.** Comparison between the Seyfert types using $H\alpha$ and N_H .

In Figure A2, we present alternative plots, considering only those objects with $N_H^{int} = 0$. This serves as a secondary validation of our results, focusing exclusively on the intrinsic characteristics of the selected sources. In Sections 3.2 and 3.3, we plotted the total N_H parameter, which includes the Galactic component. This may scatter the points as it is not solely dependent on intrinsic obscuration, as in the optical case. However, as depicted in the following figure, the results confirm the earlier conclusions and emphasize the absence of a clear separation between jetted and non-jetted sources. Statistical tests further support this secondary conclusion: Spearman ($r_s = 0.73$, p -value = 2.39×10^{-4}), Pearson ($r_p = 0.81$, p -value = 1.23×10^{-5}), and Kendall ($r_k = 0.54$, p -value = 8.27×10^{-4}) for the first panel, Spearman ($r_s = -0.07$, p -value = 7.67×10^{-1}), Pearson ($r_p = 0.09$, p -value = 7.16×10^{-1}), and Kendall ($r_k = 0.02$, p -value = 9.22×10^{-1}) for the second panel.

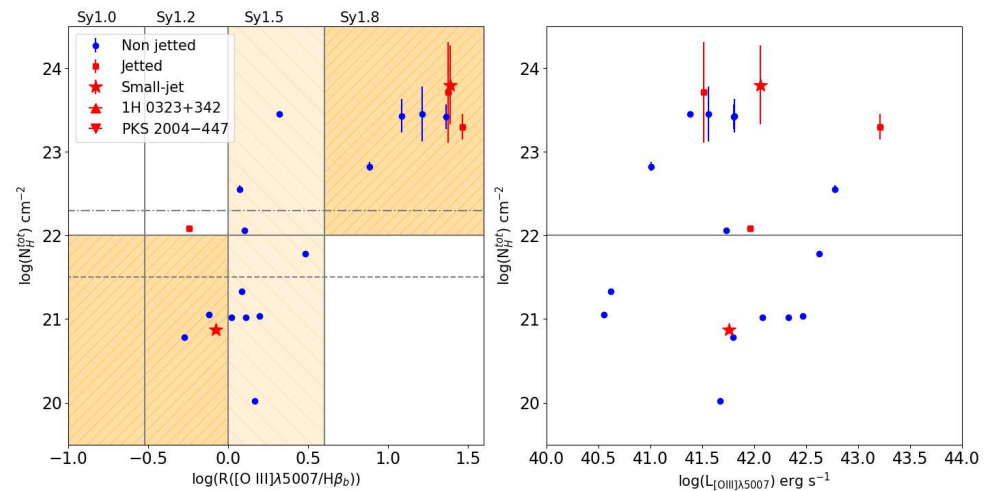


Figure A2. The same plots as in Figures 2 and 3 excluding the objects with $N_H^{int} = 0$.

Notes

- ¹ <https://www.cosmos.esa.int/web/xmm-newton/sas-threads> (accessed on 28 January 2024)
- ² <https://heasarc.gsfc.nasa.gov/cgi-bin/Tools/w3nh/w3nh.pl> (accessed on 28 January 2024)
- ³ <https://heasarc.gsfc.nasa.gov/xanadu/xspec/manual/Models.html> (accessed on 28 January 2024)

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