



Dark matter capture by atomic nuclei

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ABSTRACT

We propose a new strategy to search for a particular type of dark matter via nuclear capture. If the dark matter particle carries baryon number, as motivated by a class of theoretical explanations of the matter-antimatter asymmetry of the universe, it can mix with the neutron and be captured by an atomic nucleus. The resulting state de-excites by emitting a single photon or a cascade of photons with a total energy of up to several MeV. The exact value of this energy depends on the dark matter mass. We investigate the prospects for detecting dark matter capture signals in current and future neutrino and dark matter direct detection experiments.

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1. Introduction

Two of the greatest mysteries of modern particle physics are the nature of dark matter and the origin of the matter-antimatter asymmetry of the universe. Among the plethora of existing theoretical models, minimal extensions of the Standard Model addressing both of those questions simultaneously are of particular interest.

One of the simplest interactions of a fermionic dark sector particle χ (a gauge singlet) with the baryonic sector can be effectively described by a dimension six operator

$$\mathcal{O} \sim \frac{qqq\bar{\chi}}{\Lambda^2}, \quad (1)$$

where Λ denotes the scale of new physics mediating this interaction, e.g., the mass of a heavy color triplet scalar mediator. A possible realization of this operator is $\epsilon^{ijk} \bar{u}_{iR}^c d_{jR} d_{kR} \bar{\chi} / \Lambda^2$. The dark sector particle χ carries baryon number $B = 1$.

The existence of such an interaction has been proposed as a possible solution to the neutron lifetime anomaly [1]. The anomaly arises from the fact that experiments measuring the free neutron lifetime by determining the change in the number of neutrons in time yield a lifetime shorter than experiments sensitive only to protons in the final state. The two results can be reconciled if the

neutron has a new *dark decay* channel with a $\sim 1\%$ branching fraction.

It has been argued [2–4] that such a dark decay channel, in the absence of additional interactions involving the dark particle, would allow neutron stars to reach masses only up to $\sim 0.8 M_\odot$, substantially below the observed cases with masses $2 M_\odot$. This problem, however, is solved either by introducing self-interactions in the dark sector [5,6] or an additional repulsive interaction between the dark matter and the neutron [7]. Interestingly, it has also been shown that models with neutron dark decay offer a successful framework for explaining the matter-antimatter asymmetry of the universe [8].

One of the possible neutron dark decay channels involves a dark sector particle χ and a photon in the final state, i.e., $n \rightarrow \chi \gamma$. This channel has been searched for directly in experiment and a 1% branching fraction is in tension with data [9]. Part of the available parameter space is further constrained by the results of the Borexino experiment [10] translated into a limit on the hydrogen lifetime [11].

Two particularly appealing scenarios are realized when the dark particle χ from the neutron dark decay $n \rightarrow \chi \gamma$ is the dark matter or the antiparticle of dark matter. In the latter case, which has been considered in Refs. [12,13], $\bar{\chi}$ is then the dark matter and carries baryon number $B = -1$. It can, therefore, annihilate with nucleons, leading to spectacular signatures in various experiments. Such dark matter-nucleon annihilation signatures were considered in a more general context in [14,15].

In this letter we focus on the first possibility, i.e., when the dark particle χ is the dark matter, and we propose a complementary

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method to search for χ . Since in this case χ carries baryon number $B = 1$, as the Earth moves through the dark matter halo in our galaxy, χ can be captured by atomic nuclei through its mixing with the neutron. We explore the prospects of using large volume neutrino experiments and dark matter detectors to look for such processes.

2. Preliminaries

The operator $\epsilon^{ijk} \bar{u}_{iR}^c d_{jR} d_{kR} \bar{\chi} / \Lambda^2$ induces mixing between the dark matter particle χ and the neutron. At the hadron level, such a theory, including also the neutron magnetic moment interaction, is described by the effective Lagrangian

$$\mathcal{L}_{\text{eff}} = \bar{n} \left(i \not{\partial} - m_n + \frac{g_n e}{8m_n} \sigma^{\mu\nu} F_{\mu\nu} \right) n + \bar{\chi} (i \not{\partial} - m_\chi) \chi + \epsilon (\bar{n} \chi + \bar{\chi} n), \quad (2)$$

where the model-dependent mixing parameter $\epsilon \sim b/\Lambda^2$ (with $b = 0.0144(3)(21) \text{ GeV}^3$ [16]) and g_n is the neutron g -factor. One of the possible particle physics realizations involves a heavy color triplet scalar mediating the mixing between the dark matter and the neutron [1]. The resulting dark matter mass eigenstate contains a small admixture of the neutron,

$$|\chi\rangle' = |\chi\rangle + \frac{\epsilon}{m_n - m_\chi} |n\rangle \quad (3)$$

and vice versa. In this scenario, if energetically allowed, the neutron decays to χ and a photon at a rate

$$\Delta\Gamma_{n \rightarrow \chi\gamma} = \frac{g_n^2 e^2}{128\pi} \left(1 - \frac{m_\chi^2}{m_n^2} \right)^3 \frac{m_n \epsilon^2}{(m_n - m_\chi)^2}. \quad (4)$$

A branching fraction for this dark decay channel at the level of 1% would explain the neutron lifetime anomaly. The allowed dark matter mass range is

$$937.993 \text{ MeV} < m_\chi < 938.783 \text{ MeV}, \quad (5)$$

where the lower bound assures that none of the stable nuclei undergo dark decays, whereas the upper bound is necessary for the stability of χ (with respect to β decay).

Fig. 1 shows the values of the dimensionless parameter $\epsilon/(m_n - m_\chi)$ that yield the neutron dark decay $n \rightarrow \chi\gamma$ branching fractions: 1% (red curve), 0.5% (green) and 0.1% (blue), for dark matter masses in the range specified in Eq. (5). The boundaries of the orange and gray-shaded regions correspond to the 90% confidence level upper limits on $\epsilon/(m_n - m_\chi)$ based on the analysis of the low-energy photon spectrum of the Borexino data conducted in Ref. [11] and the direct search for $n \rightarrow \chi\gamma$ [9], respectively.

3. Dark matter capture

Thus far, it has not been appreciated that in models with a mixing between the dark matter and the neutron, it is possible for the dark matter particle to be captured by atomic nuclei. This process can be interpreted as a nuclear capture of an off-shell neutron with a mass and kinetic energy equal to those of the dark matter particle.

The capture of χ on a nucleus (A, Z) leads to

$$\chi + (A, Z) \rightarrow (A+1, Z)^* \rightarrow (A+1, Z) + \gamma_c, \quad (6)$$

where γ_c denotes a single photon or a cascade of photons from the de-excitation of $(A+1, Z)^*$ to the ground state. For nonrelativistic χ ($v_\chi \ll c$), the total energy of the cascade is

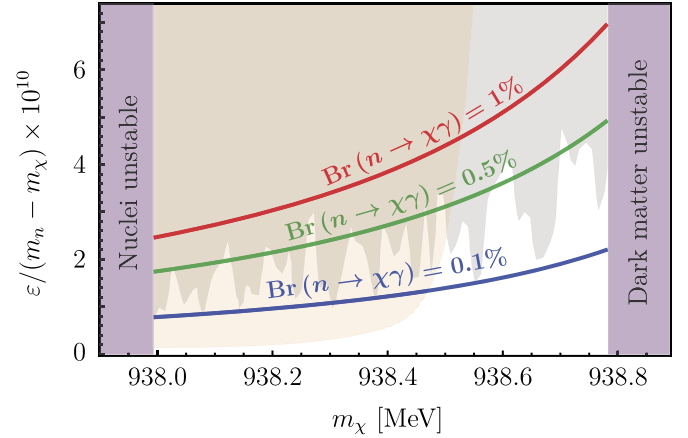


Fig. 1. The values of $\epsilon/(m_n - m_\chi)$ in units of 10^{-10} (as a function of the dark matter mass m_χ) that yield the neutron dark decay branching fractions: 1% (red curve), 0.5% (green) and 0.1% (blue). For dark matter masses $m_\chi < 937.993 \text{ MeV}$ stable nuclei would undergo dark decays, whereas for $m_\chi > 938.783 \text{ MeV}$ the dark matter χ would be unstable with respect to β decays. The boundaries of the orange and gray-shaded regions are the upper limits on $\epsilon/(m_n - m_\chi)$ derived from the Borexino data in Ref. [11] and from the direct search for $n \rightarrow \chi\gamma$ [9], respectively, both at the 90% confidence level.

$$E_c = S(n) - (m_n - m_\chi), \quad (7)$$

where $S(n)$ is the neutron separation energy in $(A+1, Z)$, and we neglected the kinetic energy of χ . Given Eq. (3), the cross section for dark matter capture can be written as

$$\sigma_\chi = \sigma_{n^*} \frac{\epsilon^2}{(m_n - m_\chi)^2}, \quad (8)$$

where σ_{n^*} is the cross section for the capture of an off-shell neutron, or, approximately, a particle with identical properties as the neutron, but with mass m_χ and kinetic energy E_k equal to the kinetic energy of χ .

In the energy region of the capture state far away from resonances, the only contribution to the cross section σ_{n^*} comes from the nonresonant (NR) channel capture. This contribution was calculated for standard neutron capture in Ref. [17] (see also [18]), and in our case it can be written as

$$\sigma_{n^*}^{\text{NR}} = \frac{0.062}{R\sqrt{E_k}} \frac{Z^2}{A^2} \sum_f \mu \frac{2J_f + 1}{6(2I + 1)} S_{dp} W(y_f), \quad (9)$$

where

$$W(x) = \frac{(x+3)^2}{(x+1)^2} x^2 \left[1 + \frac{(R - a_s)}{R} \frac{x(x+2)}{x+3} \right]^2. \quad (10)$$

The sum is over all the final states resulting from the emission of the first photon in each cascade. In the above equations I , J_f , S_{dp} , a_s and R are: the target spin, the final state spin, the spectroscopic factor, the coherent scattering length (roughly equal to the potential scattering radius) and the nuclear radius, respectively. The variable μ takes the following values: if $I = 0$ then $\mu = 1$; if $I \neq 0$ then $\mu = 1$ for $J_f = I \pm 3/2$ and $\mu = 2$ for $J_f = I \pm 1/2$. The parameter y_f is given by

$$y_f = \frac{\sqrt{2m_\chi E_f}}{\hbar} R, \quad (11)$$

where E_f is the energy of the first emitted photon.

The first term inside the brackets in Eq. (10) corresponds to hard sphere capture, whereas the second term represents the contribution from distant resonances. Since $m_\chi \approx m_n$, the only differ-

ence between our case and neutron capture is in the values of E_f , which are related via

$$E_f^{(\chi \text{ capture})} = E_f^{(n \text{ capture})} - (m_n - m_\chi). \quad (12)$$

For $E_f^{(n \text{ capture})}$ on the order of several MeV, the cross section σ_n^{NR} is not very different from the nonresonant channel cross section of standard neutron capture σ_n^{NR} .

If the capture state energy is within the resonance region, there are additional contributions to the cross section that depend on the resonance structure. In the case of dark matter capture, the energy of the capture state is below the threshold for neutron emission, i.e., it is in the region of bound states. The exact contribution of the resonant capture to the cross section depends on how well the state resulting from the capture of a dark matter particle with a given kinetic energy matches the energy of the nuclear state. Consequently, the dark matter capture cross section will depend on the velocity distribution of the particles being captured, on the density of states in the capture energy region, and on the widths of the nuclear states.

For bound states the only available de-excitation channels are through photon emission, which implies that the states have widths ~ 1 eV or smaller. In later discussion, we mainly focus on argon and xenon as our target elements. The excitation energies range from 5 to 8 MeV, and the densities of nuclear states are on the order of 40/MeV for ^{41}Ar (the spacing between the states is ~ 25 keV) and 20 000/MeV for ^{132}Xe (the spacing is ~ 50 eV).

Regarding the kinetic energy of the dark matter particle χ being captured, within the Standard Halo Model [19] the dark matter velocity distribution with respect to the center of the Milky Way can be approximated by the Maxwell-Boltzmann distribution

$$f(\vec{v}) = \begin{cases} \frac{1}{N_{\text{esc}}} \left(\frac{3}{2\pi\sigma_v^2} \right)^{3/2} \exp\left(-\frac{3\vec{v}^2}{2\sigma_v^2}\right) & \text{for } |\vec{v}| < v_{\text{esc}}, \\ 0 & \text{for } |\vec{v}| > v_{\text{esc}}, \end{cases} \quad (13)$$

with σ_v being the root mean square velocity dispersion and $N_{\text{esc}} = \text{Erf}(v_{\text{esc}}/v_0) - 2v_{\text{esc}} \exp(-v_{\text{esc}}^2/v_0^2)/(\sqrt{\pi} v_0)$, where $v_0 = \sqrt{2/3}\sigma_v \approx 235$ km/s is the most probable velocity and $v_{\text{esc}} \approx 500 - 600$ km/s is the escape velocity.

Since the Solar System is moving around the center of the Milky Way with a speed $v_\oplus \approx 220$ km/s [20], the expected dark matter velocity distribution measured on Earth is given by $f(\vec{v} - \vec{v}_\oplus)$. The most probable dark matter velocity shifts to ~ 320 km/s. Translating this into the dark matter kinetic energy distribution, the most probable kinetic energy of χ (measured on Earth) is ~ 530 eV and the half-width of the distribution is ~ 1100 eV.

An additional contribution to the kinetic energy distribution of χ comes from the Earth's motion within the Solar System. The Earth's velocity around the Sun is ~ 30 km/s and its orbit is inclined at an angle $\sim 60^\circ$ relative to the plane of the Milky Way. Therefore, the kinetic energy distribution of χ in experiments on Earth will shift by $\sim \pm 50$ eV over the course of the year. For a detailed discussion of the resulting annual modulation of dark matter signals, see Ref. [21].

Given the large width of the dark matter kinetic energy distribution and the narrowness of the nuclear bound states, the resonant χ capture does not dramatically increase the total capture cross section above its nonresonant value. To make a rough estimate of this enhancement, consider the case of xenon. Assuming a spacing between the states of ~ 50 eV, their width ~ 1 eV and an average enhancement within a resonance over the nonresonant contribution of ~ 1000 (which corresponds to the enhancement theoretically predicted for neutron capture by ^{131}Xe at $E_k \sim 500$ eV), the resonant contribution to the dark matter cap-

ture cross section is only ~ 10 times larger than the nonresonant contribution.

A nontrivial resonant structure would also lead to a varying cross section for dark matter capture throughout the year due to the shift in the kinetic energy distribution of χ , thus an annual modulation of the signal would be expected. We leave a detailed analysis of these effects for future work and limit ourselves to qualitative estimates for the annual modulation.

The rate of dark matter capture is calculated as

$$R = n \int d^3\vec{v} f(\vec{v} - \vec{v}_\oplus) \Phi_\chi(v) \sigma_\chi(v), \quad (14)$$

where $f\Phi_\chi$ is the dark matter flux and n is the number of target nuclei. Given the dark matter density in the vicinity of the Solar System of ~ 0.3 GeV/cm³, for a dark matter with mass in the range given in Eq. (5), the flux (governed by the velocity distribution of χ) has a maximal value $\Phi_\chi(v_0) \approx 10^7$ cm⁻² s⁻¹. The number of target nuclei n for a detector filled with a substance of mass M and mole density ρ_{mol} is $n = N_0 M / \rho_{\text{mol}}$, where N_0 is the Avogadro number.

4. Signals in experiments

Signatures of dark matter capture can be searched for in large volume neutrino experiments and dark matter direct detection experiments which are capable of recording MeV photon energy depositions. We first investigate the possible constraints on χ capture by hydrogen arising from the Borexino experiment [22]. We then discuss the unique signatures of χ capture by argon at the future Deep Underground Neutrino Experiment (DUNE) [23] and investigate the prospects of discovering χ capture by xenon at PandaX [24], XENON1T [25] and LUX [26]. Finally, we comment on the sensitivity of other existing and upcoming neutrino and dark matter experiments, and propose detector materials that would increase the dark matter capture signal.

Borexino

Borexino recorded the photon spectrum of low-energy neutrinos interacting with electrons [27]. The detector contained ~ 70 tons of pseudocumene (made up of hydrogen, carbon and oxygen) and the exposure time was ~ 1300 days.

The cross sections for radiative neutron capture by those constituents at energies $E_k \sim 530$ eV are [28]: $\sigma_n(^1\text{H}) \approx 2$ mb, $\sigma_n(^{12}\text{C}) \approx 20$ μb and $\sigma_n(^{16}\text{O}) \approx 4$ μb . Since the rate of dark matter capture by hydrogen is several orders of magnitude larger than on carbon and oxygen, in the following analysis of Borexino data we focus solely on hydrogen (we estimate its mass in the detector to be ~ 7 tons).

Neutron capture by hydrogen and subsequent de-excitation of deuterium lead to the emission of a single photon with energy $E_\gamma^{(n)}(^1\text{H}) = 2.225$ MeV [29]. This implies that the signature of dark matter capture by hydrogen,

$$\chi + ^1\text{H} \rightarrow ^2\text{H}^* \rightarrow ^2\text{H} + \gamma, \quad (15)$$

consists of a single monochromatic photon with energy

$$E_\gamma^{(\chi)} = 2.225 \text{ MeV} - \Delta m, \quad (16)$$

where $\Delta m = m_n - m_\chi$. The photon energy is thus uniquely determined by the dark matter mass. Equation (5) implies

$$0.653 \text{ MeV} < E_\gamma^{(\chi)} < 1.443 \text{ MeV}. \quad (17)$$

There is no resonant behavior of the capture cross section in the case of hydrogen. Using Eqs. (8) and (9), one arrives at

$$\sigma_{\chi}({}^1\text{H}) = \sigma_n({}^1\text{H}) \frac{W\left(\frac{\sqrt{2m_{\chi}E_{\gamma}^{(\chi)}}}{\hbar} R\right)}{W\left(\frac{\sqrt{2m_n E_{\gamma}^{(n)}}}{\hbar} R\right)} \frac{\varepsilon^2}{(m_n - m_{\chi})^2}, \quad (18)$$

where $W(x)$ is defined in Eq. (10). Adopting the values of parameters for hydrogen, i.e., $R \approx 0.84$ fm [30,31] and $a_s \approx -3.74$ fm [32], the ratio $\sigma_n^{\text{NR}}/\sigma_n^{\text{NR}}$ falls in the range $0.19 - 0.55$, with the exact value depending on the mass of χ . Substituting in Eq. (14) the velocity distribution from Eq. (13) and the cross section from Eq. (18), one obtains, upon setting $\sigma_n^{\text{NR}}/\sigma_n^{\text{NR}} \sim 0.4$, the expected rate of dark matter capture at Borexino of

$$R_B \sim 3 \times 10^{-5} \left[\frac{\varepsilon}{m_n - m_{\chi}} \times 10^{10} \right]^2 \text{ events/day}. \quad (19)$$

The Borexino photon spectrum, which is in good agreement with solar neutrino and radioactive backgrounds, can be used to derive upper bounds on $\varepsilon/(m_n - m_{\chi})$. Following the method adopted in Ref. [9], we performed a Z-test for the statistical significance of a possible dark matter capture signal in the publicly available Borexino data [33].

For photon energies ~ 1 MeV, the detector resolution is ~ 50 keV [27]. We combined the data bins ($\sim 2 - 3$ keV each) into 150 keV segments, each with an assigned energy corresponding to the middle bin in the segment. This way, assuming the dark matter capture signal in the detector is a Gaussian of width equal to the resolution, each segment contains most of the signal corresponding to the energy of its central bin. We then calculated, for each segment (labeled by i) separately, the number of standard deviations between the data (with the theoretical background subtracted) and the predicted signal according to the Z-test formula $\Delta\sigma = (d_i - s_i)/\sqrt{(\Delta d_i)^2 + (\Delta s_i)^2}$, where d_i is the number of photon counts per day with the background subtracted (provided in the publicly available Borexino data [33]), s_i is the signal count per day, and Δd_i , Δs_i are the corresponding uncertainties obtained using Poisson statistics.

Taking into account the detector's efficiency of 25% [22], we find that the parameter region $\varepsilon/(m_n - m_{\chi}) \lesssim 7 \times 10^{-9}$ is not constrained by our analysis of the Borexino data, thus the bounds from the direct search for $n \rightarrow \chi\gamma$ [9] and the hydrogen lifetime [11] are not improved.

A much larger experiment which, a priori, could look for dark matter capture by hydrogen, is Super-Kamiokande [34]. It consists of 50 kilotons of water and records events via Cherenkov radiation. However, the energy threshold for event selection is 3.5 MeV [35], which eliminates the potential photon signal of dark matter capture by hydrogen. Similar arguments apply to the Sudbury Neutrino Observatory (SNO) [36].

DUNE

The DUNE detector will be based on liquid argon time projection chamber (LArTPC) technology, which offers unprecedented accuracy in reconstructing energies and positions of photon energy depositions. A study of the reach for MeV photons in a LArTPC environment was conducted in the context of the ArgoNeUT experiment and demonstrated very promising results [37].

DUNE is planned to contain 40 kilotons of liquid argon. At kinetic energies $E_k \sim 530$ eV the radiative neutron capture cross section is $\sigma_n({}^{40}\text{Ar}) \approx 5$ mb [28]. The dark matter capture by ${}^{40}\text{Ar}$ would lead to

$$\chi + {}^{40}\text{Ar} \rightarrow {}^{41}\text{Ar}^* \rightarrow {}^{41}\text{Ar} + \gamma_c, \quad (20)$$

Table 1

The photon energies for the dominant cascades from dark matter capture by ${}^{40}\text{Ar}$, along with their branching fractions (Br) and the range of possible values of the ratio $\sigma_n^{\text{NR}}({}^{40}\text{Ar})/\sigma_n^{\text{NR}}({}^{40}\text{Ar})$ for each cascade separately; $\Delta m = m_n - m_{\chi}$.

Br	Photon energies [MeV]	$\sigma_n^{\text{NR}}/\sigma_n^{\text{NR}}$
31%	4.745 $-\Delta m$, 1.187, 0.167	0.64 – 0.82
9%	5.583 $-\Delta m$, 0.516	0.68 – 0.84
6%	4.745 $-\Delta m$, 0.838, 0.516	0.64 – 0.82
6%	2.772 $-\Delta m$, 2.811, 0.516	0.42 – 0.70
4%	3.701 $-\Delta m$, 1.044, 1.187, 0.167	0.55 – 0.77

where γ_c denotes cascades of photons. The neutron separation energy for the ${}^{41}\text{Ar}$ nucleus is $S(n) = 6.099$ MeV [38], which implies the total energy per cascade

$$E_c = 6.099 \text{ MeV} - \Delta m, \quad (21)$$

where $\Delta m = m_n - m_{\chi}$. It follows from Eq. (5) that

$$4.527 \text{ MeV} < E_c < 5.317 \text{ MeV}. \quad (22)$$

Provided the decay pattern of the χ capture state is similar to that of the neutron capture state, the energies of photons in the cascades can be determined from the known data for neutron capture by ${}^{40}\text{Ar}$ [39–41]. The only difference is that the energy of the first photon in each cascade from dark matter capture will be smaller by Δm than the energy of the corresponding photon in a neutron capture process.

The energies of photons in the five dominant cascades from dark matter capture by ${}^{40}\text{Ar}$ are summarized in Table 1 along with their branching fractions. Cascades that are not listed have branching fractions less than 2%. Only two ${}^{41}\text{Ar}$ excited states have known lifetimes: 0.167 MeV ($\tau = 0.32$ ns) and 0.516 MeV ($\tau = 0.26$ ns); lifetimes of the other excited states of ${}^{41}\text{Ar}$ are likely also shorter than 1 ns. Table 1 provides also the suppression factors $\sigma_n^{\text{NR}}({}^{40}\text{Ar})/\sigma_n^{\text{NR}}({}^{40}\text{Ar})$ due to the lower energy of the first photon in each cascade compared to the neutron capture case. To compute those factors, an analogous formula to the one in Eq. (18) was applied to each channel of dark matter capture by argon, with the adopted parameter values for argon $R \approx 3.43$ fm [42] and $a_s \approx 1.83$ fm [32].

Although there are no resonances for neutron capture by argon at energies ~ 530 eV, a resonance behavior may arise in the case of dark matter capture. However, as discussed earlier, due to the expected narrowness and low density of nuclear states in argon, as well as the broad velocity distribution of dark matter, the resonant contribution will not change the cross section significantly.

Given the comparable sizes of the suppression factors $\sigma_n^{\text{NR}}/\sigma_n^{\text{NR}}$ for each cascade, in order to estimate the rate of dark matter capture at DUNE we assume $\sigma_n^{\text{NR}}({}^{40}\text{Ar}) \approx 0.7 \sigma_n^{\text{NR}}({}^{40}\text{Ar})$ for the total cross section. This gives

$$R_{\text{DUNE}} \approx 0.02 \left[\frac{\varepsilon}{m_n - m_{\chi}} \times 10^{10} \right]^2 \text{ events/day}. \quad (23)$$

Based on the constraints shown in Fig. 1, the largest allowed value of $\varepsilon/(m_n - m_{\chi})$ is $\sim 5 \times 10^{-10}$, which yields a dark matter capture rate at DUNE of ~ 0.5 events/day.

Since the spacing between nuclear states in ${}^{41}\text{Ar}$ at energies relevant for dark matter capture is ~ 25 keV, and the half-width of the dark matter kinetic energy distribution is ~ 1 keV, it is likely that there will be no resonant enhancement of the signal – with a high probability the energy of the nuclear state will lie outside of the excitation energy range populated in dark matter capture. A numerical calculation, e.g., using the TALYS code [43], is required to arrive at a more definite answer, but it is beyond the scope of this paper. If there exists a state falling within the energy range

populated in dark matter capture, an annual modulation of the signal will occur. Assuming that: (i) the width of the resonance is ~ 1 eV, (ii) the resonance enhances the nonresonant contribution by a factor of ~ 1000 , (iii) the dark matter kinetic energy distribution shifts by ± 50 eV due to the Earth's motion around the Sun, the signal will have variations of only up to 30% over the course of the year.

A dark matter capture rate of ~ 0.5 events/day with its unique signature is very nontrivial. Whether it can be distinguished from the background depends on the progress of several ongoing research and technological development efforts at DUNE and the final experimental setup. In particular, the optimal photon detection strategy is still under investigation.

Also, of great importance is the reduction of the neutron background in this case. This background comes mainly from the neutron capture process where the neutrons originate from the environment. They can be produced either by radioactivity in the rock (radiogenic) or by muons passing through the rock (cosmogenic). Those two backgrounds were estimated in Ref. [44] and can be significantly reduced by implementing a hermetic passive water shield. Such a shield with a thickness of 1 m would reduce the radiogenic neutron capture rate down to ~ 0.01 events/day/kton. The cosmogenic neutron background can be efficiently removed by identifying the accompanying electromagnetic showers that enter the detector [44,45]. Therefore, with a ~ 1 m of shielding, this neutron capture background would be comparable to the dark matter capture rate. Given the difference in the photon cascade energies in the neutron capture case versus the dark matter capture case, the resolution needed for the detection of dark matter capture would be $\sim 10\%$.

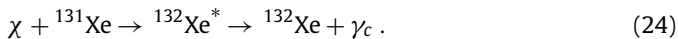
The above considerations do not take into account other possible contributions to the neutron capture background, e.g., from neutrons produced internally from radioactivity in the detector. The level of this type of background, however, can be controlled by the purification procedures of liquid argon. In any case, a detailed experimental analysis is needed before making definitive claims regarding the detectability of dark matter capture at DUNE.

PandaX, XENON1T, LUX

Several experiments with detectors based on xenon time projection chamber technology are searching for dark matter nuclear recoils. They are using natural xenon, for which the most abundant isotopes are: ^{129}Xe (26.4%), ^{131}Xe (21.2%) and ^{132}Xe (26.9%). In their final runs PandaX [24] and LUX [26] had fiducial volumes of ~ 300 kg, whereas XENON1T [25] had ~ 2 tons of xenon.

Let us consider the isotope ^{131}Xe . In the kinetic energy region $E_k \sim 530$ eV, the neutron capture cross section, with a nonresonant contribution $\sigma_n^{\text{NR}}(^{131}\text{Xe}) \sim 1$ b, exhibits a resonant behavior. According to theoretical predictions, it can reach up to ~ 700 b at the resonance peak, but the width of the resonances is small [28].

The dark matter capture by ^{131}Xe would result in cascades of photons through the process



The neutron separation energy for the final state ^{132}Xe is $S(n) = 8.937$ MeV, thus the total energy of the cascade would be $E_c = 8.937$ MeV $- \Delta m$, i.e.,

$$7.365 \text{ MeV} < E_c < 8.155 \text{ MeV}. \quad (25)$$

In the case of xenon, the structure of the photon cascades is very complex. To arrive at a rough estimate for the dark matter capture rate, we approximate $\sigma_n^{\text{NR}}(^{131}\text{Xe}) \approx \sigma_n^{\text{NR}}(^{131}\text{Xe}) \sim 1$ b. The nonresonant contribution leads to

$$R_{\text{Xe}} \sim 10^{-5} \left[\frac{\varepsilon}{m_n - m_\chi} \times 10^{10} \right]^2 \frac{\text{events}}{\text{ton} \times \text{day}}, \quad (26)$$

which, for the largest experimentally allowed value of $\varepsilon/(m_n - m_\chi) \sim 5 \times 10^{-10}$, gives only $\sim 10^{-4}$ events/ton/day. The spacing between the xenon nuclear states at energies relevant for dark matter capture is ~ 50 eV and their width is ~ 1 eV, while the half-width of the dark matter kinetic energy distribution is ~ 1 keV. Since the resonant spectrum is so dense in that region, and since the dark matter capture cross section multiplied by the dark matter flux is velocity independent (see, Eq. (9)), the annual modulation of the signal is weak. Similar conclusions apply to XENON1T and LUX.

At PandaX-II a detailed evaluation of the neutron background has been performed [46]. It was determined for two low-background runs: Run 9 (26-ton-day) and Run 10 (28-ton-day) [46]. The estimated neutron capture rate was ~ 0.7 events/ton/day in Run 9 and ~ 0.4 events/ton/day in Run 10. Those signals came from neutron capture mostly on ^{129}Xe and ^{131}Xe , since those two isotopes have the largest cross sections for thermal (~ 25 meV) neutron capture: ~ 20 b and 100 b, respectively. The corresponding total energies of the resulting photon cascades (the current PandaX detector is capable of measuring only the total energy deposition and cannot distinguish each photon), equal to the neutron separation energies in ^{130}Xe and ^{132}Xe , are 9.256 MeV and 8.937 MeV. The capture by other, less abundant xenon isotopes produce cascades with energies < 6.5 MeV.

The current energy resolution at PandaX is about 5–10%. The nonresonant dark matter capture by xenon gives a signal rate which is a few orders of magnitude lower than the neutron capture background, making it difficult to tell them apart. However, the energy resolution is planned to be improved at PandaX-II and XENON1T in the ~ 10 MeV energy regime. In addition, as estimated in Sec. 3, the resonant contribution to the capture cross section can be comparable or even larger (up to a factor of ~ 10) than that of the nonresonant channel, making it easier to distinguish the signal from the background. Combining a better energy resolution and an enhanced capture cross section from resonance effects with a much larger amount of data to be collected, future experiments, such as PandaX-4T, XENONnT and LUX-Zeplin, may provide promising probes of the dark matter capture process. On the other hand, neutrinoless double beta decay experiments with excellent energy resolution, like Kamland-Zen [47] and EXO-200 [48], which use ^{136}Xe as the detector material and were shown to be sensitive to the formation of weak-scale dark matter-nucleus bound states [49], would give a much lower rate for dark matter capture due to the small neutron capture cross section for ^{136}Xe .

5. Discussion

We have proposed a new method to search for a particular kind of dark matter particle based on its capture by atomic nuclei. The signature of such a process is a single photon or a cascade of photons resulting from the de-excitation of the capture state. The energy of this single photon or the total energy of the cascade (and of the first photon in each cascade) distinguishes the signal from the neutron capture background.

The ideal places to look for dark matter capture are large volume neutrino experiments and dark matter direct detection experiments, especially those with scintillation detectors and time projection chamber technology. Whether the signal can stand out from the background, mainly from neutron capture, depends on the detailed design of each experiment. The ongoing research and development at DUNE, and future improvements at various xenon experiments, show great promise in searching for dark matter capture. A detailed analysis of the event reconstruction efficiency,

object identification capability and detector resolution are still needed to reach a concrete conclusion.

In some liquid scintillator experiments, like KamLAND, which measures the antineutrino flux from nuclear reactors, a delayed coincidence cut is imposed to reduce the background [50]. This cut removes the potential signal from dark matter capture. It would be interesting to perform an analysis of the photon background with this cut abandoned.

It might be worth exploring the possibility of using other detector materials with large neutron capture cross sections, e.g., iodine and samarium [28]. Out of those, iodine is already widely used as a scintillator in the form of sodium iodine or cesium iodine. Further examples of nuclei with large neutron capture cross sections are terbium and gadolinium, which could also be used to construct future detectors with increased sensitivity to dark matter capture. A more economical solution would be to dope existing detectors with such substances. This has been considered in a different context for the case of gadolinium at Super-Kamiokande [51,52]. Given the expected nontrivial resonance structure of the nuclear states, an annual modulation of the signal may be observed, although only for light-element materials and with a rather small amplitude.

Finally, we would like to emphasize that the dark matter capture explored in this letter offers an alternative way to look for the neutron dark decay channel $n \rightarrow \chi \gamma$, with a similar or better sensitivity than direct searches.

Declaration of competing interest

The authors declare that they have no known competing financial interests or personal relationships that could have appeared to influence the work reported in this paper.

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