

Impact of interacting dark sector on Cosmic Microwave Background Radiation

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Abstract. It has been shown earlier that an arc-like pattern found on the Cosmic Microwave Background Radiation (CMBR) may result from the decay of dark matter particles initiating near particle horizon in the Q -phase of the interacting cosmological constant (ICC) model. In the present work, an investigation is made into how the corresponding decay of such dark matter particles might influence these signatures, in view of the recent data on CMBR and the diffuse glow of the anomalous microwave radiation. We also discuss the constraints on such decay imposed by the interaction of the cosmological constant with the background. These predictions made in the ICC model can be verified in the concordance space of multiple observations.

1. Introduction

Dark energy has emerged as a major component of the universe following the observations of type Ia supernovae (SNe Ia) [1, 2] and the acoustic peak at the Legendre multipole $l_{\text{peak}} \sim 197$ in the Cosmic Microwave Background Radiation (CMBR) power spectrum [3].

We have shown previously [4] that in a three-phase universe λ must vary due to interaction (then we call it a parameter) in the intermediate phase II known as Q phase. In phase I and III, however, it remains a constant and does not interact, where it causes the required acceleration. Thus in Q phase we have $T_{(n);k}^{ik} = -T_{(\lambda);k}^{ik} = Q^i$ and only the combined stress energy tensor is conserved. The interaction pulls down the energy density of the cosmological constant ρ_λ quite fast, while the fall of matter ρ_n occurs much more slowly than without interaction.

In the following Section 2 we mention the background results for the λ -matter interaction and the evolution of their relative energy densities in Q phase. In Section 3 we discuss a scenario where in this phase of the universe the energy transfer from the decaying cosmological constant would generate the dark matter particles with wide mass spectrum whose signatures must be observed in form of the mass dependent arcs on the CMBR power spectrum among other observable influences on structure formation [5].

2. Decay of cosmological constant through interaction in Q phase

The model for interacting cosmological constant is in part motivated by the simple consideration of the effective curvature parameter $k_{\text{eff}} = k - \lambda a^2/3$ in the Friedmann equations. The CMBR observations suggest $k_{\text{eff}} = 0$ that yields a^{-2} dependence of λ with the time dependence $\lambda \propto H^{4/3}$ in matter dominated (MD) and $\lambda \propto H$ radiation dominated (RD) universe. Several authors have discussed forms of λ dependent on H [6, 7]. With $L_\lambda/L_P \sim 10^{61}$ (where L_λ and L_P are the



λ - and the Planck length scales respectively), a plausible way to remove the inconsistency of $\lambda L_P^2 \sim 10^{-123}$ would be to find a physical mechanism that allows a swift decline of λ . We have attempted to achieve this by means of the interaction between λ and matter components.

Against the commonly prevalent view among the cosmologists, we emphasize that our local observations, by virtue of their nature, would always indicate a spatially flat universe, and it is not clear how we relate it to the global geometry [8, 9]. Any motivation to “solve the flatness problem” as claimed by inflationary theories appear to be actually silent, contrary to the common opinion, on the curvature parameter *per se* since the entire k/a^2 term is thrown out like “baby with the water” from the Friedmann equations because of inflating scale factor, and not because one ‘gets’ $k = 0$. (This term might haunt again in the contracting phase of a closed universe !)

In the phase I of the 3-phase model, the universe can undergo transition from RD to λ dominated universe or remain in the λ throughout this phase, depending on the initial constant value of ρ_λ . However, in the next Q phase, we have energy conservation as

$$\dot{\rho}_n + 3H\rho_n(1 + w_n) = -\dot{\rho}_\phi - 3H\rho_\phi(1 + w_\phi) = Q \quad (1)$$

with ρ_n , w_n standing for dominant background and ρ_ϕ , w_ϕ for the dark energy scalar field. The scalar fields with or without interaction have been discussed by several authors e.g. [10]. However, we consider here an interacting cosmological constant. In λ CDM model $w_\phi = w_\lambda = -1$. Three alternative courses exist here. One, $w_\lambda = -1$ is held fixed throughout the evolution of the universe with $\rho_\lambda = -\dot{Q}$, as we consider in this paper. Then it is found that interaction of λ will reduce its energy density with the consequence of change in any one or more parameters of λ , c or G . Two, if all these (λ , c , G) are truly constants of time, $w_\lambda = -Q/3H\rho_\lambda - 1$ and $Q = 0$. Three, both w_λ and ρ_λ may vary like dynamical quintessence.

The scalar field dynamics is given as

$$\ddot{\phi} + 3H\dot{\phi} + V'(\phi) = -\frac{Q}{\dot{\phi}}. \quad (2)$$

The solutions of (1) provide $\rho_n \propto f a^{-3(1+w_n)}$ with $f(\lambda/M_P) = a^{Q/H\rho_n}$ as the coupling function ($M_P = (8\pi G)^{-1/2}$ is the reduced Planck mass) and keeping $Q/H\rho_n$ constant, while $\rho_\lambda \propto a^{-Q/H\rho_\lambda}$. The interaction parameter clearly depends on the variation of the coupling function as

$$Q \propto \dot{f}/a^{3(1+w_n)}. \quad (3)$$

In the absence of interaction $f = 1$ and ρ_n declines as $a^{-3(1+w_n)}$ as expected. Regardless of the background being matter or radiation dominated, $f \propto Qt^3$ for constant Q . The fast increase of coupling function with time during Q phase is mainly responsible for conversion of energy of dark cosmological constant into dark matter particles as also for slowing down the rate of fall of ρ_n in the concave-like manner. Thus, with a non-zero Q we have higher matter density at the end of Q phase than without interaction.

For a spatially flat universe with $\Omega_n^0 + \Omega_\lambda^0 = 1$, then we have

$$\frac{Q}{\dot{\phi}} \sim H_0^2 M_P^2 (1+z)^{3(1+w_n)} f'(\phi) \quad (4)$$

$$V'(\phi) = -\frac{Q}{\dot{\phi}}. \quad (5)$$

It is found that for the observed range of $w_\phi \simeq -1$, ($w = -0.969 \pm 0.061$ (stat) ± 0.065 (sys)) $\dot{\phi} \approx 0$ and thus the right hand side of the (2) and (5) will be too high even for small value of Q .

This shows the important role played by the interaction even if there is no cosmological constant proper with $w_\lambda = -1$, but instead, a set of quintessence fields having $w_\phi \approx -1$. Thus the range of present observational uncertainties in determining EOS of dark energy is consistent with the concept of interaction presented in this work. The coupling function shows the time dependence of the ratio of energy densities $r = \rho_\lambda/\rho_n$ as $rf^{1+1/r} \propto a^4$ in RD and $\propto a^3$ in MD universe.

3. Diagnostics for coupling and dark matter generation

It was shown earlier that the third derivatives, and not the second derivatives, of scale factor of the universe provide effective state finders [11] for pinning down the interaction among the components as $u = \ddot{\alpha}/(aH^3)$ and $s = (u - 1)/[3(q - 1/2)]$ where $\alpha = \dot{a} = aH$. In our case of cosmological constant this set becomes

$$u = 1 - \frac{9}{2} \left(\frac{Q}{3H(\rho_n + \rho_\lambda)} \right) \quad (6)$$

$$s = \frac{Q}{3H\rho_\lambda}. \quad (7)$$

In terms of (7) we have the coupling function as $f = a^{3rs}$. In non-interacting phases (initial phase I and the presently ongoing phase III) no mechanism to generate dark matter can operate and so $s = 0$. However, such mechanism can clearly work in Q phase with $s \neq 0$ (though not constant). In this way the universe has greater matter density ρ_n at the end of Q phase, thus alleviating the problem of producing light nuclei like deuterium-2 and lithium-7 in the standard Friedmann-Robertson-Walker (FRW) model. In addition to it, the part of the decaying λ energy is used up in generation of dark matter. These processes must have a deep influence on the structure formation in the subsequent course of evolution.

We envisage a scenario to generate particles of dark matter (with Lorentz factor γ) each with energy

$$\xi = \int \Theta^{00} d\tau = \gamma mc^2 \quad (8)$$

where Θ^{00} is contributed by the T^{00} of λ field. For the same given ξ the more massive particles would travel more slowly than the lighter particles. Thus more massive particles would settle further away from the light cone envelopes, while the lighter particles traveling with relativistic speeds remain closer to the particle horizon. This particle mass distribution would in effect generate a multiple arc-like pattern on the Last Scattering Surface (LSS) around the recombination epoch ($z \approx 1000$) frozen on the CMBR, (and other observations [12]), each arc identifiable with a characteristic mass of definite dark matter particle (or its decay product releasing net energy $\tilde{\xi}$). Equation (3) shows the decay rate depends on the strength of interaction and as coupling grows with time as $\propto t^3$ during Q phase greater is the amount of λ energy used up to create the dark matter particles of broad mass spectrum.

It is therefore expected to have an isotropic streaming of particles, produced simultaneously (the vertex of the light cone lies at this epoch) but leaving their signatures in form of multiple arcs at LSS. For a vertex at a given epoch prior to the LSS, the arcs of smaller radii correspond to more massive particles and those of large radius towards the horizon R_H (horizon size ≈ 0.18 Mpc at $t_{\text{recombination}}$ now blown to ≈ 200 Mpc) correspond to the lighter particles. The size of an arc $L \sim R_H v/c$ for a typical particle having streaming velocity v can be related to its mass as

$$m = \xi(1 - L^2/R_H^2)^{-1/2}. \quad (9)$$

Thus, the mass density may be estimated from the vacuum considerations [13] with a suitable cut-off as

$$\Pi = \frac{\hbar}{4\pi^2} \int_0^{k_m} \omega_k k^2 dk (1 - L^2/R_H^2)^{-1/2}. \quad (10)$$

Putting the present energy scales $\sim 10^{-52} GeV^4$, $\rho_\lambda^0 \simeq 4 GeV$. This must include the uncertainty in mass given as

$$\Delta m \geq \hbar/\Delta L \quad (11)$$

where putting the maximum value of the present horizon $R_H^0 \approx 3000$ Mpc, we obtain $\Delta m \sim 10^{-32}$ eV.

We expect the intensive dark matter search experiments, like XENON100, DAMA/LIBRA, CDMS II, as well as future LUX or XMASS etc. would look for the lighter dark matter particles below ~ 10 GeV as the decay products of more massive particles, and at the same time for the signatures of decay energy ξ released in form of γ radiation.

However, the particle production being a continuous process throughout Q phase, an overlapping of orders seems inevitable for different sets of particles produced at different epochs. In this scenario it seems that the more massive particles would be the most suitable sites for the star formation in proto-galaxies due to stronger self-gravity, and so would yield age for the youngest stars. The above conclusions of the interacting cosmological constant would be accepted only at the touchstone of such future observations.

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